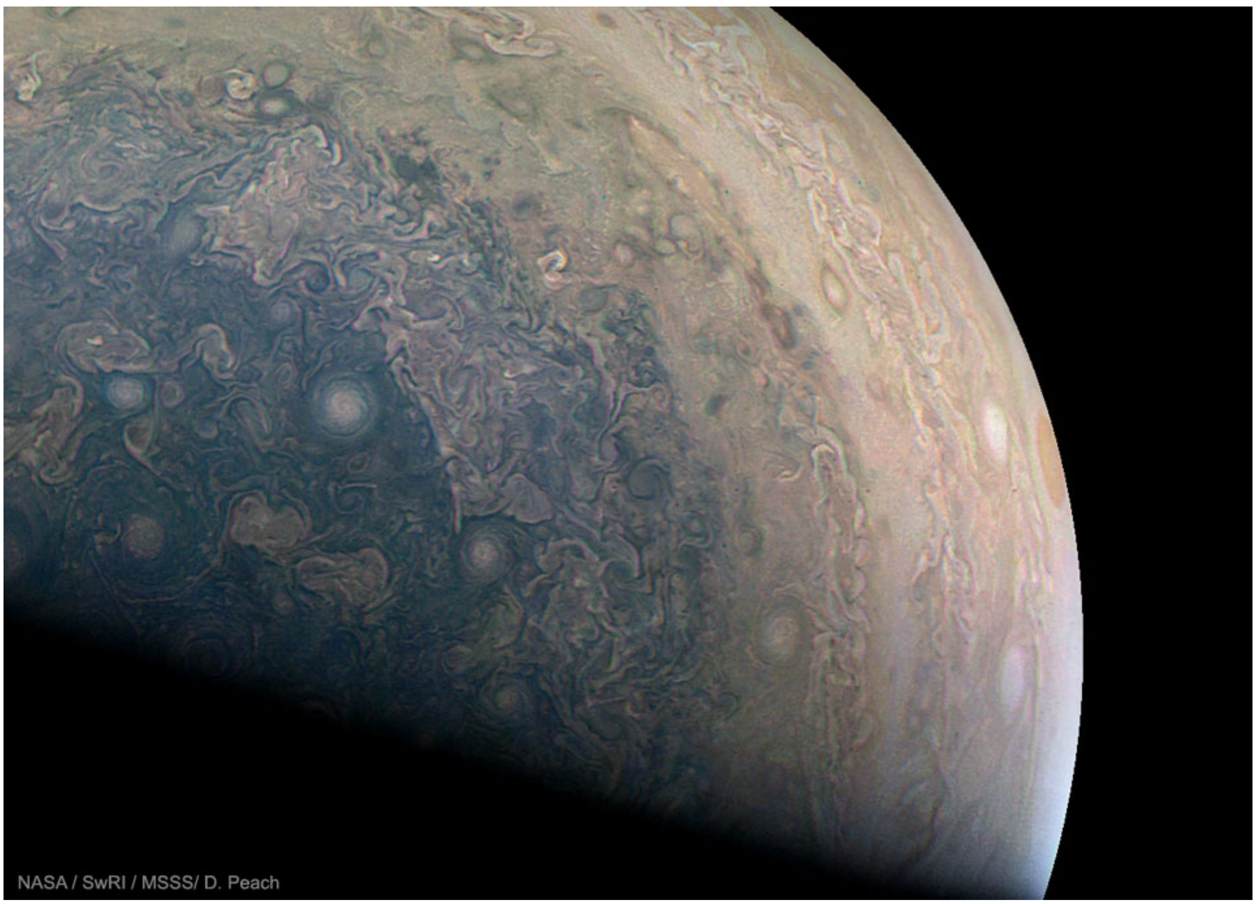
UNIVERSITY OF COLORADO AT BOULDER
DEPARTMENT OF ASTROPHYSICAL AND PLANETARY SCIENCES

ASTR1030 Laboratory
Accelerated Introduction to Astronomy

Fall 2019



IMPORTANT INFORMATION

NAME:
PHONE:
EMAIL:

COURSE TITLE: ASTR 1030 | Accelerated Introduction to Astronomy I

Please fill in the blanks below with information from the syllabus and/or lectures.

LECTURE

Class Time: Monday, Wednesday, Friday 9:00-9:50am

Location: Duane G130 (and sometimes Fiske Planetarium)
Instructor: Dr. Zach Berta-Thompson
Office Location:
Office Hours:
E-mail: Zach.BertaThompson@colorado.edu

Lecture TA: Daniel Everding
Office Location:
Office Hours:
E-mail: Daniel.Everding@colorado.edu

LABORATORY

Section Number:
Lab Time:
Lab Location: Sommers-Bausch Observatory Room S-175
Lab TA (+ LA):
Office Location(s):
Office Hours:
E-mail(s):

NIGHT OBSERVING SESSIONS

Dates and Times:

Front Cover: Cloud swirls around Southern Jupiter from Juno
Back Cover: Planetary Travel Poster for the Trappist-1 system from NASA Jet Propulsion Laboratory

TABLE OF CONTENTS

* = *Clear skies are required for this exercise*
° = *Clear skies are needed for a portion of the exercise*

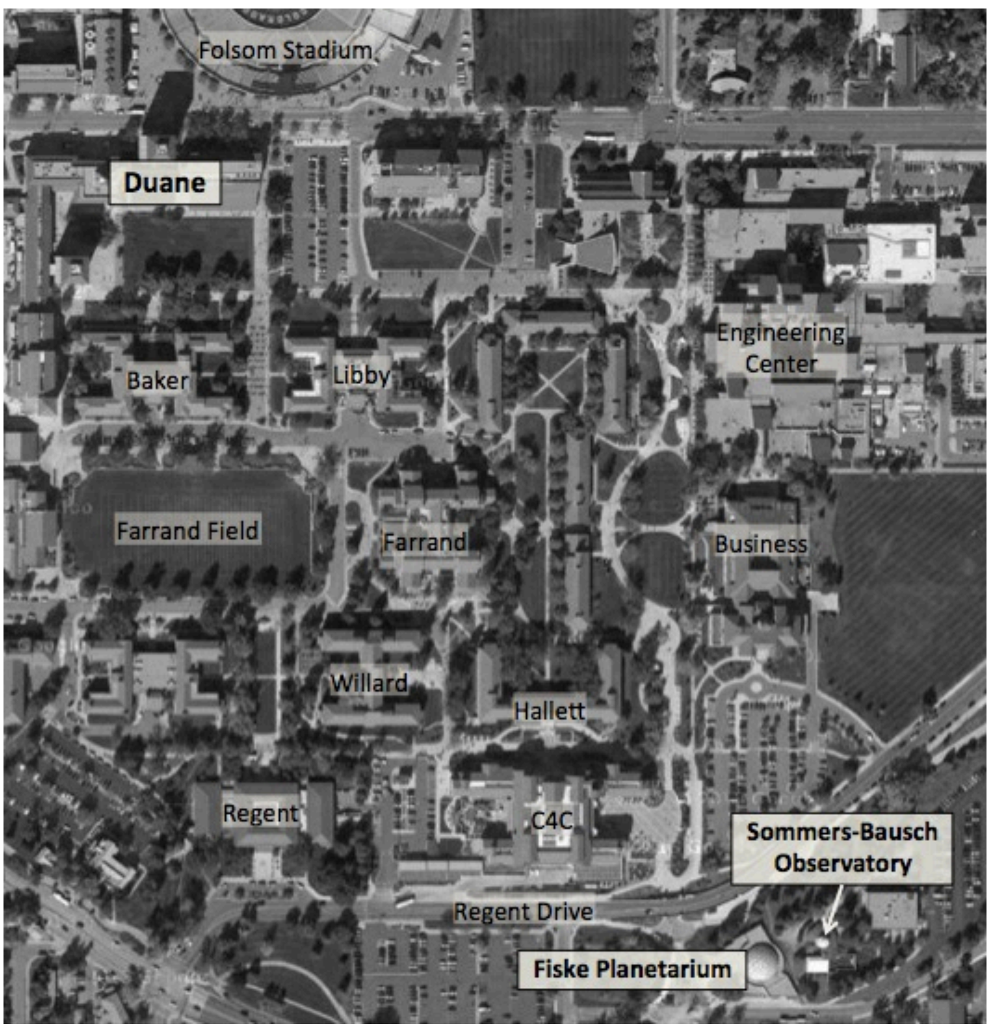
GENERAL INFORMATION

You *must* enroll for *both* the lecture section *and* a specific laboratory section. You cannot switch lab sections within the week.

Your **lecture** section will usually be held in the classroom in Duane Physics & Astrophysics Building (just south of Folsom Stadium). Several lectures will be held instead at the Fiske Planetarium (at the intersection of Regent Drive and Kittridge Loop).

Your **laboratory** section will meet once per week during the daytime in Room S175 at Sommers-Bausch Observatory (just east of the Fiske Planetarium). Follow the walkway around the south side of Fiske and up the hill to the Observatory.

You will also have **nighttime observing sessions** using the Sommers-Bausch Observatory telescopes to view and study the constellations, the moon, planets, stars, and other celestial objects.



MATERIALS

The following materials are needed:

- **ASTR 1030 Astronomy Lab Manual** (*this booklet*), available from the CU Bookstore. Replacement copies are available in Acrobat PDF format downloadable from the SBO website, under Education → Courses.
- **Calculator**. All students should have access to a scientific calculator that can perform scientific notation, exponentials, and trig functions (sines, cosines, etc.).
- **A 3-ring binder** to hold this lab manual, your lab notes, and lab write-ups.

THE LABORATORY SECTIONS

Your laboratory session will meet for 1 hour and 45 minutes in the daytime once each week in the Sommers-Bausch Observatory (SBO), Classroom S175. Each lab section will be run by a lab instructor, who will also grade your lab exercises and assign you scores for the work you hand in. Your lab instructor will give you organizational details and information about grading at the first lab session.

The lab exercises do not exactly follow the lectures or the textbook. We concentrate on *how* we know what we know; thus, we spend more time making and interpreting observations. Modern astronomers, in practice, spend almost no time at the eyepiece of a telescope. They work with photographs, satellite data, computer images, or computer simulations. In our laboratory, we will explore both the traditional and more modern techniques.

You are expected to attend all lab sessions. Most of the lab exercises can *only* be done using the equipment and facilities in the SBO classroom. Therefore, if you do not attend the daytime lab sessions, you *cannot* complete those experiments and cannot get credit. Most of the observational exercises can *only* be done at night using the Observatory telescopes. If you do not attend the nighttime sessions, you *cannot* complete these either.

NIGHTTIME OBSERVING

You are expected to attend some of the nighttime observing labs. These are held in the evenings approximately every second or third week at the Sommers-Bausch Observatory, Mondays through Thursdays. Your lab instructor will tell you the dates and times. Write the dates and times of the nighttime sessions on your calendar so you do not miss them. If you have a job or other obligations that conflict with the nighttime sessions, it is *your* responsibility to make arrangements with your instructor to attend at different times.

The telescopes are not in a heated area, so dress warmly for the night observing sessions!

THE LABORATORY WRITE-UPS

You are expected to turn in lab assignments of *collegiate* quality. This means that they must be neat and easy to read, well-organized, and demonstrating a mastery of the English language (grammar and spelling), as well as mastery of the subject matter. For most labs, you will turn in your lab report by simply removing the appropriate pages from this manual and turning in your answers at the end of the lab period.

HONOR CODE

The University of Colorado Honor Code will be strictly enforced. Plagiarism *will not be tolerated* and can result in academic and/or non-academic sanctions. Specifically, we point out the following guideline regarding your laboratory assignments:

All work turned in must be your own. You should understand all work that you write on your paper. We encourage you to work in groups if it is helpful, but you must not copy the work of someone else. We encourage you to consult friends for help in understanding problems. However, if you copy answers blindly, it will be considered a breach of the Honor Code.

SOMMERS-BAUSCH OBSERVATORY

Sommers-Bausch Observatory (SBO, http://www.colorado.edu/sbo), on the University of Colorado campus, is operated by the Department of Astrophysical and Planetary Sciences (APS). SBO provides hands-on observational experience for CU undergraduate students, and research opportunities for University of Colorado astronomy graduate students and faculty. Telescopes include two 20-inch Cassegrain telescopes installed in 2017, a 24-inch Ritchey-Cretien Cassegrain telescope in its own dome, and a 10.5-inch aperture heliostat. In its teaching role, the Observatory is used by approximately 1500 undergraduate students each year to view celestial objects that might otherwise only be seen on the pages of a textbook or discussed in classroom lectures.

The 10.5-inch aperture heliostat is equipped for viewing sunspots, measuring the solar rotation, implementing solar photography, and studying the solar spectrum. A unique optical system called SCRIBES permits simultaneous observations of the photosphere (using white light) and the solar chromosphere (using red light from hydrogen atoms, and ultraviolet light from calcium atoms which absorb and emit light within the upper solar atmosphere).

Open Houses for free public viewing through the 20-inch telescopes are held every Friday evening that school is in session. Students are encouraged to attend. Call 303-492-5002 for starting times; call 303-492-6732 for general astronomical information.

FISKE PLANETARIUM

The Fiske Planetarium and Science Center (http://www.colorado.edu/fiske) is used as a teaching facility for classes in astronomy, planetary science, and other courses that can take advantage of this unique audiovisual environment. The star theater seats 210 under a 62-foot dome that serves as a projection screen, making it the largest planetarium between Chicago and California.

Astronomy programs designed to entertain and to inform are presented to the public on Fridays and Saturdays, and to schoolchildren on weekdays. Laser-light shows rock the theater late Friday nights as well. Following Friday evening star-show presentations, visitors are invited next door to view the celestial bodies at Sommers-Bausch Observatory, weather permitting. The Planetarium provides students with employment opportunities to assist with show production and presentation, and in the daily operation of the facility.

ASTRONOMICAL WEBSITES

Celestial Objects

US Naval Observatory http://www.usno.navy.mil/USNO/astronomical-applications/
Lunar Phases https://svs.gsfc.nasa.gov/4537
Meteor Showers http://www.theskyscrapers.org/meteors/
Night Sky Viewing: http://www.seasky.org/astronomy/astronomy-calendar-current.html
Eclipses and Transits of Solar System Objects http://eclipse.gsfc.nasa.gov/eclipse.html
Sky & Telescope At a Glance http://www.skyandtelescope.com/observing/ataglance
Constellations http://www.hawastsoc.org/deepsky/constellations.html

Planets, Near and Far

Solar System Simulator http://space.jpl.nasa.gov/
Super Planet Crash http://www.stefanom.org/spc/
Planet Hunters https://www.planethunters.org/
Exoplanets https://exoplanets.nasa.gov/

Some Current (or Recent) Missions

Mars Reconnaissance Orbiter http://mars.jpl.nasa.gov/mro/
Mars Exploration Rovers http://mars.nasa.gov/mer/
Mars Odyssey http://mars.jpl.nasa.gov/odyssey/
MAVEN (at Mars) http://lasp.colorado.edu/home/maven/
Dawn (to Vesta and Ceres) http://dawn.jpl.nasa.gov/
Cassini-Huygens Mission (to Saturn) http://saturn.jpl.nasa.gov/
New Horizons Mission (to Pluto) http://pluto.jhuapl.edu/
Juno (to Jupiter) http://missionjuno.swri.edu
Hubble Space Telescope http://hubblesite.org/
Spitzer Space Telescope http://www.spitzer.caltech.edu/
Kepler Space Telescope http://kepler.nasa.gov/
Transiting Exoplanet Survey Satellite http://tess.mit.edu

Space News

Astronomy Picture of the Day: https://apod.nasa.gov
The Planetary Society: http://www.planetary.org
Bad Astronomy: http://www.syfy.com/tags/bad-astronomy
Sky & Telescope: http://www.skyandtelescope.com
Space Weather: http://www.swpc.noaa.gov
NASA TV: http://www.nasa.gov/multimedia/nasatv
Daily Space News: http://www.space.com

UNITS AND CONVERSIONS

Modern science uses the **metric system** (the SI, *Systeme International d'Unites*, internationally agreed upon system of units) with the following fundamental units:

- The **meter** (m) for length.
- The **kilogram** (kg) for mass.
- The **second** (s) for time.

Since the primary units are the **m**eter, **k**ilogram, and **s**econd, this is sometimes called the **mks system**. (Astronomers often also use another metric system with centimeters, grams, and seconds as its fundamental units, called the **cgs system**.)

All of the unit relationships in the metric system are based on multiples of 10, so it is very easy to multiply and divide. This system uses prefixes to make multiples of the units. All of the prefixes represent powers of 10. The table below provides prefixes used in the metric system, along with their abbreviations and values.

Metric Prefixes

Prefix	Abbreviation	Value
deci-	d	10^{-1}
centi-	c	10^{-2}
milli-	m	10^{-3}
micro-	m	10^{-6}
nano-	n	10^{-9}
pico-	p	10^{-12}
femto-	f	10^{-15}

Prefix	Abbreviation	Value
decka-	da	10^{1}
hecto-	h	10^{2}
kilo-	k	10^{3}
mega-	M	10^{6}
giga-	G	10^{9}
tera-	T	10^{12}
peta-	P	10^{15}

The United States is one the few countries in the world that has not yet made a complete conversion to the metric system. Even Great Britain has adopted the SI system; so, what are officially called "English" units are now probably better termed "American." As a result, Americans must often convert between English and metric units, because all science and international commerce is transacted in metric units. Some common conversions are

Imperial to metric		
1 inch	=	2.54 cm
1 mile	=	1.609 km
1 lb	=	0.4536 kg
1 gal	=	3.785 liters

metric to Imperial		
1 m	=	39.37 inches
1 km	=	0.6214 mile
1 kg	=	2.205 pound
1 liter	=	0.2642 gal

Strictly speaking, the conversion between kilograms and pounds is valid only on the Earth, because kilograms measure *mass* while pounds measure *weight*. However, since most of you will be remaining on the Earth for the foreseeable future, we will not yet dwell on this detail here. (Strictly, the unit of weight in the metric system is the *newton*, and the unit of mass in the English system is the *slug*.)

Using the "Well-Chosen 1"

Many people have trouble converting between units because, even with the conversion factor at hand, they are not sure whether they should *multiply* or *divide* by that number. The problem becomes even more confusing if there are multiple units to be converted, or if there is need to use intermediate conversions to bridge two sets of units. We offer a simple and foolproof method for handling the problem.

We all know that any number multiplied by 1 equals itself, and also that the reciprocal of 1 equals 1. We can exploit these simple properties by choosing our 1's carefully so that they will perform a unit conversion for us, so long as we remember to always include our units.

Suppose we wish to know how many kilograms a 170-pound person weighs. We know that 1 kg = 2.205 pounds, and can express this fact in the form of 1's:

$$1 = \frac{1\text{ kg}}{2.205\text{ pounds}} \qquad \text{or its reciprocal} \qquad 1 = \frac{2.205\text{ pounds}}{1\text{ kg}}$$

Note that the 1's are *dimensionless*. In other words, the quantity (number *with* units) in the numerator is exactly equal to the quantity (number *with* units) in the denominator. If we took a shortcut and omitted the units, we would be writing nonsense: of course, without units, neither 1 divided by 2.205, nor 2.205 divided by 1, equals "1"! Now we can multiply any other quantity by these 1's, and the quantity will remain unchanged (even though it will look considerably different).

In particular, we want to *multiply* the quantity "170 pounds" by 1 so that it will still be equivalent to 170 pounds but expressed in kg units. But which "1" do we choose? Very simply, *if the unit we want to "get rid of" is in the numerator, we choose the "1" that has that same unit appearing in the denominator* (and vice versa), so that the unwanted units will cancel. In our example, we can write:

$$170\text{ lbs} \times 1 \quad = \quad 170\text{ lbs} \times \frac{1\text{ kg}}{2.205\text{ lbs}} \quad = \quad \frac{170 \times 1}{2.205} \times \frac{\cancel{\text{lbs}} \times \text{kg}}{\cancel{\text{lbs}}} \quad = \quad 77.1\text{ kg}$$

Be certain not to omit the units, but multiply and divide them just like ordinary numbers. If you have selected a "well-chosen" 1 for your conversion, then your units will nicely cancel, assuring you that the numbers themselves will also have been multiplied or divided properly. This is what makes this method foolproof: if you accidentally used a "poorly-chosen" 1, the expression itself will immediately let you know about it:

$$170\text{ lbs} \times 1 \quad = \quad 170\text{ lbs} \times \frac{2.205\text{ lbs}}{1\text{ kg}} \quad = \quad \frac{170 \times 2.205}{1} \times \frac{\text{lbs} \times \text{lbs}}{\text{kg}} \quad = \quad 375 \times \frac{\text{lbs}^2}{\text{kg}} \quad !$$

Strictly speaking, this is not really incorrect: 375 lbs^2/kg *is* the same as 170 lbs, but this is not a very useful way of expressing this, and it is certainly not what you were trying to do…

Example: As a passenger on the Space Shuttle, you notice that the inertial navigation system shows your orbital velocity to be 8,042 meters per second. You remember from your astronomy course that a speed of 17,500 miles per hour is the minimum needed to maintain an orbit around the Earth. Should you be worried?

$$8042\,\frac{\text{m}}{\text{s}} \quad = \quad \frac{8042\text{ m}}{1\text{ s}} \times \frac{1\text{ km}}{1000\text{ m}} \times \frac{1\text{ mile}}{1.609\text{ km}} \times \frac{60\text{ s}}{1\text{ min}} \times \frac{60\text{ min}}{1\text{ hr}}$$

$$= \frac{8042 \times 1 \times 1 \times 60 \times 60}{1 \times 1000 \times 1.609 \times 1 \times 1} \times \frac{\text{m x km x mile x s x min}}{\text{s x m x km x min x hr}}$$

$$= 17{,}993 \ \frac{\text{miles}}{\text{hour}}$$

Your careful analysis using "well-chosen 1's" indicates that you are fine, and so you will be able to perform more unit conversions!

Temperature Scales

Scales of temperature measurement are often referenced to the freezing point and boiling point of water. In the United States, the Fahrenheit (F) scale is the most common; water freezes at 32 °F and boils at 212 °F. In Europe, the Celsius system is usually used; water freezes at 0 °C and boils at 100 °C. In scientific work, it is common to use the Kelvin temperature scale. The Kelvin degree is exactly the same "size" increment as the Celsius degree, but it is based on the idea of absolute zero, the unattainable temperature at which all random molecular motions would cease. Absolute zero is defined as 0 K, water freezes at 273 K, and water boils at 373 K. Note that the degree mark is not used with Kelvin temperatures, and the word "degree" is commonly not even mentioned: we say that "water boils at 373 Kelvin."

To convert among these three systems, recognize that 0 K = -273 °C = -459 °F and that the Celsius and Kelvin degree is larger than the Fahrenheit degree by a factor of 180/100 = 9/5. The relationships between the systems are:

$$K = °C + 273 \qquad °C = 5/9\ (°F - 32) \qquad °F = 9/5\ K - 459$$

Energy and Power: Joules and Watts

The SI metric unit of energy is called the joule (abbreviated J). Although you may not have heard of joules before, they are simply related to other units of energy with which you may be more familiar. For example, 1 food Calorie is 4,186 joules. House furnaces are rated in btu (British thermal units), indicating how much heat energy they can produce: 1 btu = 1,054 joules. Thus, a single potato chip (with an energy content of about 9 Calories) could be said to possess 37,674 joules or 35.7 btu of energy.

The SI metric unit of power is called the watt (abbreviated W). Power is defined to be the rate at which energy is used or produced. Power is measured as energy per unit time. The relationship between joules and watts is:

$$1 \text{ watt} = 1 \ \frac{\text{joule}}{\text{second}}$$

For example, a 100-watt light bulb uses 100 joules of energy (about 1/42 of a Calorie or 1/10 of a btu) each second it is turned on. One potato chip contains enough chemical energy to operate a 100-watt light bulb for over 6 minutes!

You might be familiar with the unit of power called the horsepower; one horsepower equals 746 watts, which means that energy is consumed or produced at the rate of 746 joules per second. You can calculate (using unit conversions) that if your car has "fifty horsepower" under the hood, they need to be fed 37,300 joules, or the equivalent energy of one potato chip every second, in order to pull you down the road.

To give you a better sense of the joule as a unit of energy (and of the convenience of scientific notation, our next topic), some comparative energy outputs are listed here:

Energy Source	Energy (joules)
Big Bang	10^{68}
Supernova	10^{44} - 10^{46}
Sun's radiation for 1 year	10^{34}
U.S. annual energy consumption	10^{20}
Volcanic explosion	10^{19}
H-bomb (20 megaton)	10^{17}
Earthquake	10^{16}
Thunderstorm	10^{15}
Fission of 1 kg of Uranium-235	5.6×10^{13}
Lightening flash	10^{10}
Burning 1 liter of oil	10^{7}
Daily energy needs of average adult	10^{7}
Kinetic energy of a car at 60 mph	10^{6}
Energy expended by a 1 hour walk	10^{6}
Solar energy at Earth (per m^2 per sec)	10^{3}
Baseball pitch	10^{2}
Hitting keyboard key	10^{-2}
Hop of a flea	10^{-7}

Labeling Units

In 1999, a NASA spacecraft, the Mars Climate Orbiter traveled to the planet Mars carrying instruments intended to map the planet's surface and profile the structure of the atmosphere. Unfortunately, while it was trying to maneuver itself into orbit, the orbiter burned up in Mars' atmosphere. In the end, a rather simple problem was discovered to have caused the accident: software on-board the spacecraft reported a critical value in pounds (English units) rather than the newtons (metric unit) that the scientists were expecting. This little error caused a big difference in the calculations of the scientists and resulted in the loss of the $125 million spacecraft. A number without units is meaningless. Always label your units!

SCIENTIFIC NOTATION & SIG FIGS

What Is Scientific Notation?

Astronomers deal with quantities ranging from the truly microcosmic to the hugely macrocosmic. It would be very inconvenient to always have to write out the age of the universe as 14,000,000,000 years or the distance to the Sun as 149,600,000,000 meters. For simplicity, powers-of-ten notation is used, in which the exponent tells you how many times to multiply by 10. For example, $10 = 10^1$, and $100 = 10^2$. As another example, $10^{-2} = 1/100$; in this case the exponent is negative, so it tells you how many times to divide by 10. One slightly tricky one is to remember that $10^0 = 1$ (see the section on **Powers and Roots** below). Using powers-of-ten notation, the age of the universe is 1.4×10^{10} years and the distance to the Sun is 1.496×10^{11} meters.

- The general form of a number in scientific notation is **$a \times 10^n$**, where **a** (called the coefficient) is a number between 1 and 10, and **n** (called the exponent) is an integer.

  Correct examples of scientific notation: 6×10^2, 4.8×10^5, 8.723×10^{-3}.

  *In*correct examples of scientific notation: 34×10^5, $4.8 \times 10^{0.5}$, 0.2×10^3.

- If the number is between 1 and 10, so that its coefficient would be multiplied by 10^0 (=1), then it is not necessary to write the power of 10. For example, the number 4.56 already is in scientific notation (it is not necessary to write it as 4.56×10^0, but you could write it this way if you wish).

- If the number is already a power of 10, then it is not necessary to write that it is multiplied by 1. For example, the number 100 can be written in scientific notation either as 10^2 or as 1×10^2. (Note, however, that the latter form may be necessary for entering numbers on a calculator.)

The use of scientific notation has several advantages, even for use outside of the sciences:

- Scientific notation makes the expression of very large or very small numbers much simpler. For example, it is easier to express the U.S. federal debt as $\$2 \times 10^{13}$ rather than as $20,000,000,000,000.

- Because it is so easy to multiply powers of ten in your head (by adding the exponents), scientific notation makes it easy to do "in your head" estimates of answers.

- Use of scientific notation makes it easier to keep track of significant figures; that is, does your answer really need all of those digits that pop up on your calculator?

Converting from "Normal" to Scientific Notation:

Place the decimal point after the first non-zero digit and count the number of places the decimal point has moved. If the decimal place has moved to the *left* then multiply by a positive power of 10; to the right will result in a negative power of 10.

Example: To write 3040 in scientific notation, we must move the decimal point 3 places to the left, so it becomes 3.04×10^3.

Example: To write 0.00012 in scientific notation, we must move the decimal point 4 places to the right: 1.2×10^{-4}.

Converting from Scientific to "Normal"" Notation:

If the power of 10 is positive, then move the decimal point to the right; if it is negative, then move it to the left.

Example: Convert 4.01×10^2. We move the decimal point two places to the right, making 401.

Example: Convert 5.7×10^{-3}. We move the decimal point three places to the left, making 0.0057.

Addition and Subtraction with Scientific Notation:

When adding or subtracting numbers in scientific notation, their powers of 10 must be equal. If the powers are *not* equal, then you must first write the numbers so that they all have the same power of 10.

Example: $(6.7 \times 10^9) + (4.2 \times 10^9) = (6.7 + 4.2) \times 10^9 = 10.9 \times 10^9 = 1.09 \times 10^{10}$. (Note that the last step is necessary in order to put the answer into proper scientific notation.)

Example: $(4 \times 10^8) - (3 \times 10^6) = (4 \times 10^8) - (0.03 \times 10^8) = (4 - 0.03) \times 10^8 = 3.97 \times 10^8$.

Multiplication and Division with Scientific Notation:

It is easy to multiply or divide just by rearranging so that the powers of 10 are multiplied together.

Example: $(6 \times 10^2) \times (4 \times 10^{-5}) = (6 \times 4) \times (10^2 \times 10^{-5}) = 24 \times 10^{2-5} = 24 \times 10^{-3} = 2.4 \times 10^{-2}$. (Note that the last step is necessary in order to put the answer in scientific notation.)

Example: $(9 \times 10^8) \div (3 \times 10^6) = \dfrac{9 \times 10^8}{3 \times 10^6} = (9/3) \times (10^8/10^6) = 3 \times 10^{8-6} = 3 \times 10^2$.

Approximation with Scientific Notation:

Because working with powers of 10 is so simple, use of scientific notation makes it easy to estimate *approximate* answers. This is especially important when using a calculator since, by doing mental calculations, you can verify whether your answers are reasonable. To make approximations, simply round the numbers in scientific notation to the nearest integer, then do the operations in your head.

Example: Estimate $5{,}795 \times 326$. In scientific notation, the problem becomes $(5.795 \times 10^3) \times (3.26 \times 10^2)$. Rounding each to the nearest integer makes the approximation $(6 \times 10^3) \times (3 \times 10^2)$, which is 18×10^5, or 1.8×10^6. (The exact answer is 1.88917×10^6.)

Example: Estimate $(5 \times 10^{15}) + (2.1 \times 10^9)$. Rounding to the nearest integer this becomes $(5 \times 10^{15}) + (2 \times 10^9)$. We can see that the second number is nearly $10^{15}/10^9 = 10^6$, or one million, times smaller than the first. Thus, it can be ignored in the addition, and our approximate answer is simply 5×10^{15}. (The exact answer is 5.0000021×10^{15}.)

Significant Figures

Numbers should be given only to the accuracy that they are known with certainty, or to the extent that they are important to the topic at hand. For example, your doctor may say that you weigh 130 pounds, when in fact at that instant you might weigh 130.16479 pounds. The discrepancy is unimportant and anyway will change as soon as you drink a glass of water.

If numbers are given to the greatest accuracy that they are known, then the result of a multiplication or division with those numbers cannot be determined any better than to the number of digits in the *least* accurate number.

Example: Find the circumference of a circle measured to have a radius of 5.23 cm using the formula: $C = 2\pi R$. Because the value of *pi* stored in your calculator is probably 3.141592654, the calculator's numerical solution will be

$$(2 \times 3.141592654 \times 5.23 \text{ cm}) = 32.86105916 = 3.286105916 \times 10^1 \text{ cm}.$$

If you write down all 10 digits as your answer, you are implying that you know, with absolute certainty, the circle's circumference to an accuracy of one part in 10 billion! That would require that your measurement of the radius was in error by no more than 0.000000001 cm. That is, its actual value was at least 5.229999999 cm, but no more than 5.230000001 cm.

In reality, because your measurement of the radius is known to only three decimal places, the circle's circumference is also known to only (at best) three decimal places as well. You should round the fourth digit and give the result as 32.9 cm or 3.29×10^1 cm. It may not look as impressive, but this is an honest representation of what you know about the figure.

Because the value of "2" was used in the formula, you may wonder why we are allowed to give the answer to *three* decimal places rather than just one: 3×10^1 cm. The reason is because the number "2" is *exact* - it expresses the fact that a diameter is *exactly* twice the radius of a circle - no uncertainty about it at all. Without any exaggeration, the number could have been represented as 2.0000000000000000000, but the shorthand "2" is used for simplicity. This does not violate the rule of using the number that is least accurately known.

USEFUL MATH FOR ASTRONOMY

Dimensions of Circles and Spheres

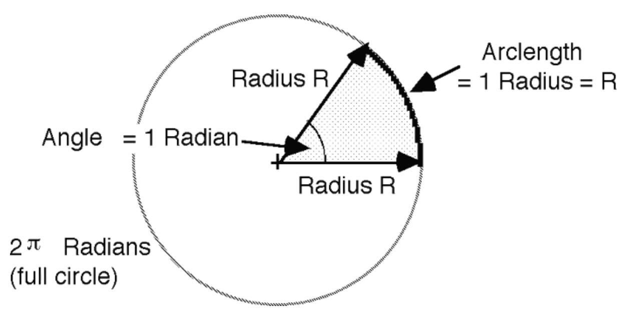
- The circumference of a circle of radius R is $2\pi R$.
- The area of a circle of radius R is πR^2.
- The surface area of a sphere of radius R is given by $4\pi R^2$.
- The volume of a sphere of radius R is $4\pi R^3/3$

Notice that the units will make sense if we propagate them through these equations. If we know the radius in units of meters, then the circumference will also have units of meters. Because they both involve R^2, the area of a circle and the surface area of a sphere will have units of meters^2. The volume will have units of meters^3, appropriate for talking about a three-dimensional volume.

Measuring Angles - Degrees and Radians

- There are 360° in a full circle.
- There are 60 arcminutes in one degree. The shorthand for arcminute is the single prime ('), so we can write 3 arcminutes as 3'. By converting units, we can see there are 360° x (60'/1°) = 21,600' in a full circle.
- There are 60 arcseconds in one arcminute. The shorthand for arcsecond is the double prime ("), so we can write 3 arcseconds as 3".) Therefore, there are 360° x (60'/1°) x (60"/1') = 1,296,000" in a full circle. In astronomy we often talk about things that are very far away, so very tiny units of angles can be very useful!

We also often express angles in units of *radians* instead of degrees. If we were to take the radius (length R) of a circle and bend it so that it conformed to a portion of the circumference of the same circle, the angle covered by that radius is defined to be an angle of *one radian*.



Because the circumference of a circle has a total length of $2\pi R$, we can fit exactly 2π radii (6 full lengths plus a little over 1/4 of an additional length) along the circumference. Thus, a full 360° circle is equal to an angle of 2π radians. In other words, an angle in radians equals the arclength of a circle intersected by that angle, divided by the radius of that circle. If we imagine a *unit* circle (where the radius = 1 unit in length), then an angle in radians equals the actual curved distance along the portion of its circumference that is "cut" by the angle.

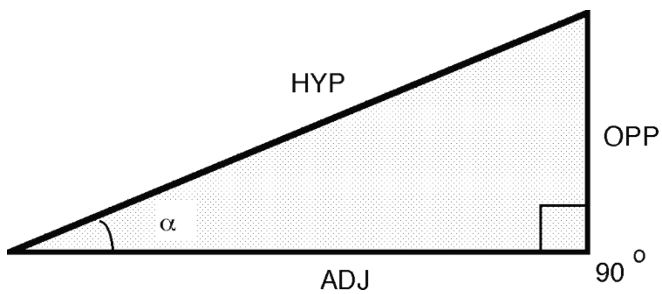
The conversion between radians and degrees is

$$1 \text{ radian} = \frac{360}{2\,\pi} \text{ degrees} = 57.3° \qquad 1° = \frac{2\,\pi}{360} \text{ radians} = 0.01745 \text{ radians}$$

Trigonometric Functions

In this course, we will make occasional use of the basic trigonometric (or "trig") functions: sine, cosine, and tangent. Here is a quick review of the basic concepts.

In any *right triangle* (where one angle is 90°), the longest side is called the *hypotenuse*; this is the side that is opposite the right angle. The trigonometric functions relate the lengths of the sides of the triangle to the other (i.e., not the 90°) enclosed angles. In the right triangle figure below, the side *adjacent* to the angle α is labeled "adj," the side *opposite* the angle is labeled "opp." The hypotenuse is labeled "hyp."



- The Pythagorean theorem relates the lengths of the sides of a right triangle to each other:

$$(\text{opp})^2 + (\text{adj})^2 = (\text{hyp})^2 \quad .$$

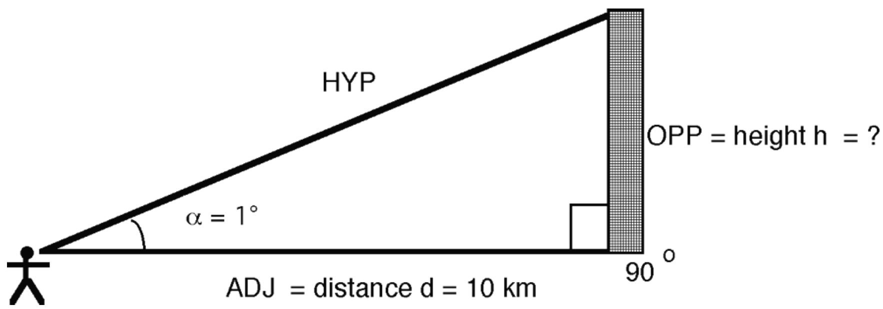
- The trig functions are just ratios of the lengths of the different sides:

$$\sin\alpha = \frac{(\text{opp})}{(\text{hyp})} \qquad \cos\alpha = \frac{(\text{adj})}{(\text{hyp})} \qquad \tan\alpha = \frac{(\text{opp})}{(\text{adj})} \quad .$$

Angular Size, Physical Size and Distance

The *angular* size of an object (the angle it "subtends," or appears to occupy from our vantage point) depends on both its true *physical* size and its *distance* from us. For example, if you stand with your nose up against a building, it will occupy your entire view; as you back away from the building it will cover a smaller and smaller angular size, even though the building's physical size is unchanged. Because of the relations between the three quantities (angular size, physical size, and distance), we need know only two in order to calculate the third.

Suppose a tall building has an angular size of 1° (that is, from our location its height appears to span one degree of angle), and we know from a map that the building is located precisely 10 km away. How can we determine the actual physical size (height) of the building?



We imagine that we are standing with our eye at the apex of a triangle, from which point the building covers an angle $\alpha = 1°$ (greatly exaggerated in the drawing). The building itself forms the *opposite* side of the triangle, which has an unknown height that we will call *h*. The distance *d* to the building is 10 km, corresponding to the *adjacent* side of the triangle.

Because we want to know the opposite side, and already know the adjacent side of the triangle, we only need to concern ourselves with the *tangent* relationship:

$$\tan\alpha = \frac{(\text{opp})}{(\text{adj})} \qquad \text{or} \qquad \tan 1° = \frac{h}{d}$$

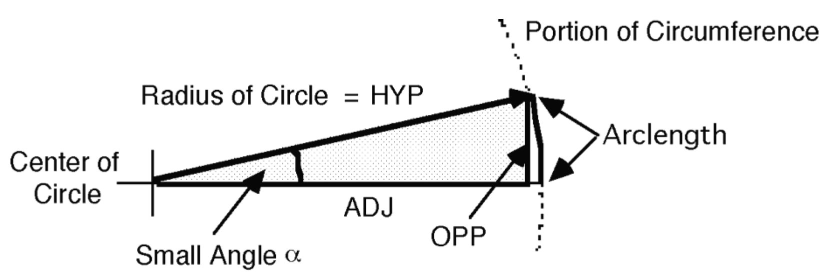
which we can reorganize to give

$$h = d \times \tan 1° \qquad \text{or} \qquad h = 10 \text{ km} \times 0.017455 = 0.17455 \text{ km} = 174.55 \text{ meters.}$$

Small Angle Approximation

We used the *adjacent* side of the triangle for the distance instead of the *hypotenuse* because it represented the smallest separation between the building and us. It should be apparent, however, that because we are 10 km away, the distance to the top of the building is only very slightly farther than the distance to the base of the building. A little trigonometry shows that the hypotenuse in this case equals 10.0015 km, or less than 2 meters longer than the adjacent side of the triangle.

In fact, *the hypotenuse and adjacent sides of a triangle are always of similar lengths whenever we are dealing with angles that are "not very large."* Thus, we can substitute one for the other whenever the angle between the two sides is small.



Now imagine that the apex of a small angle α is located at the center of a circle that has a radius equal to the hypotenuse of the triangle, as illustrated above. The *arclength* of the circumference covered by that small angle is only very slightly longer than the length of the corresponding straight ("opposite") side. In general, then, *the opposite side of a triangle and its corresponding arclength are nearly equal whenever we are dealing with small angles.*

Now we can go back to our equation for the physical height of our building:

$$h \;=\; d \times \tan\alpha \;=\; d \times \frac{(\text{opp})}{(\text{adj})} \quad .$$

Because the angle α is small, the opposite side is approximately equal to the "arclength" covered by the building. Likewise, the adjacent side is approximately equal to the hypotenuse, which is in turn equivalent to the radius of the inscribed circle. Making these substitutions, the above (exact) equation can be replaced by the following (approximate) equation:

$$h \;\approx\; d \times \frac{(\text{arclength})}{(\text{radius})} \quad .$$

But remember that the ratio (arclength)/(radius) is the definition of an angle expressed in *radian* units rather than degrees, so we now have the very useful **small angle approximation**:

> For small angles, the physical size h of an object can be determined directly from its distance d and angular size in radians by
>
> $$h \approx d \times (\text{angular size in radians}) \quad .$$
>
> Or, for small angles, the physical size h of an object can be determined from its distance d and its angular size in degrees by
>
> $$h \approx d \times \frac{2\pi}{360°} \times (\text{angular size in degrees}) \quad .$$

Using the small angle approximation, the height of our building 10 km away is calculated to be 174.53 meters high, an error of only about 2 cm (less than 1 inch)! And best of all, *the calculation did not require trigonometry*, just multiplication and division!

When can the approximation be used? Surprisingly, the angles do not really have to be very small. For an angle of 1°, the small angle approximation leads to an error of only 0.01%. Even for an angle as great as 10°, the error in your answer will only be about 1%.

Powers and Roots

We can express any *power* or *root* of a number in *exponential notation*, in which we say that b^n is the "*n*th power of *b*", or "*b* to the *n*th power." The number represented here as *b* is called the *base*, and *n* is called the *power* or *exponent.*

The basic definition of a number written in exponential notation states that the base should be multiplied by itself the number of times indicated by the exponent. That is, b^n means *b* multiplied by itself *n* times. For example: $5^2 = 5 \times 5$; $b^4 = b \times b \times b \times b$.

From the basic definition, certain properties automatically follow:

- **Zero Exponent:** Any nonzero number raised to the zero power is 1. That is, $b^0 = 1$.

  Examples: $2^0 = 10^0 = -3^0 = (1/2)^0 = 1$.

- **Negative Exponent:** A negative exponent indicates that a *reciprocal* is to be taken. That is,

$$b^{-n} = \frac{1}{b^{n}} \qquad \frac{1}{b^{-n}} = b^{n} \qquad \frac{a}{b^{-n}} = a \times b^{n}.$$

Examples: $4^{-2} = 1 / 4^2 = 1/16$; $10^{-3} = 1 / 10^3 = 1/1000$; $3 / 2^{-2} = 3 \times 2^2 = 12$.

- **Fractional Exponent:** A fractional exponent indicates that a *root* is to be taken.

$$b^{1/n} = \sqrt[n]{b}\,; \qquad b^{m/n} = \sqrt[n]{b^m} = \left(\sqrt[n]{b}\right)^m$$

Examples: $8^{1/3} = \sqrt[3]{8} = 2$ $\qquad 8^{2/3} = \left(\sqrt[3]{8}\right)^2 = 2^2 = 4$

$2^{4/2} = \sqrt{2^4} = \sqrt{16} = 4$ $\qquad x^{1/4} = \left(x^{1/2}\right)^{1/2} = \sqrt{\sqrt{x}}$

PROPORTIONALITY
or How to use the "∝" to simplify math

The ∝ symbol means "is proportional to". The use of proportionalities in astronomy is extremely common, for good reasons. This shorthand way of working with equations saves time, prevents calculator mistakes, and helps you quickly check that your answers make sense. Often the main thing you care about is how changing one variable affects your result; proportionalities allow you to answer that question very quickly.

Example: Planet Alphabet has a radius that is twice the radius of Planet Boomerang. How do their volumes compare?

The Long Way: The volume V of a sphere of radius R is $V = 4/3\pi R^3$ so we might consider calculating the exact volumes of each planet directly, and then taking their ratios:

$$V_B = 4/3\pi R_B^3$$

$$V_A = 4/3\pi R_A^3 = 4/3\pi(2R_B)^3 = 32/3\pi R_B^3$$

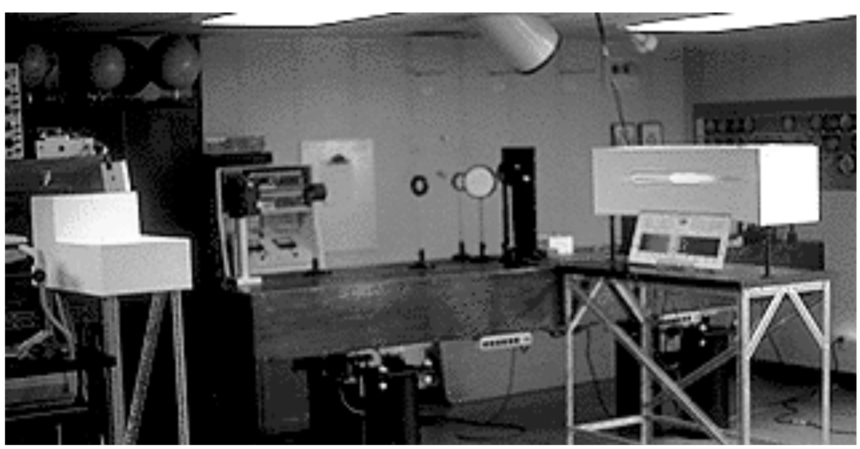
$$V_A / V_B = [32/3\pi R_B^3]/[4/3\pi R_B^3] = [32/4] = 8$$

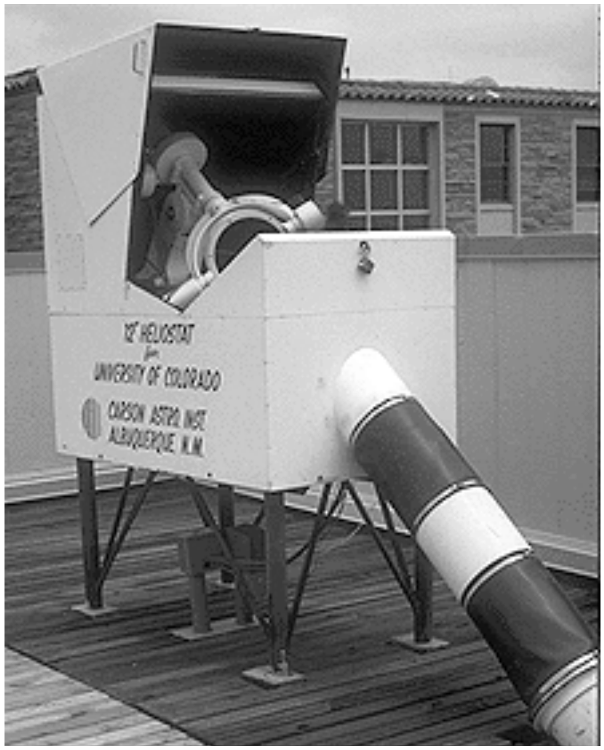
The Short Way: Instead, we could recognize that the volume of a sphere is "proportional" to the cube of the radius $V \propto R^3$ so we can simply write that the ratio of the volumes will be

$$V_A / V_B = (R_A/R_B)^3 = (2/1)^3 = 8$$

Using proportionalities is really helpful when the answer that you care about in the end is a ratio. You can often take a moment before you start doing math to think about what parts of an equation will cancel out in a ratio, and just do a calculation with the parts you care about. Astronomers use ratios and proportionalities very frequently!

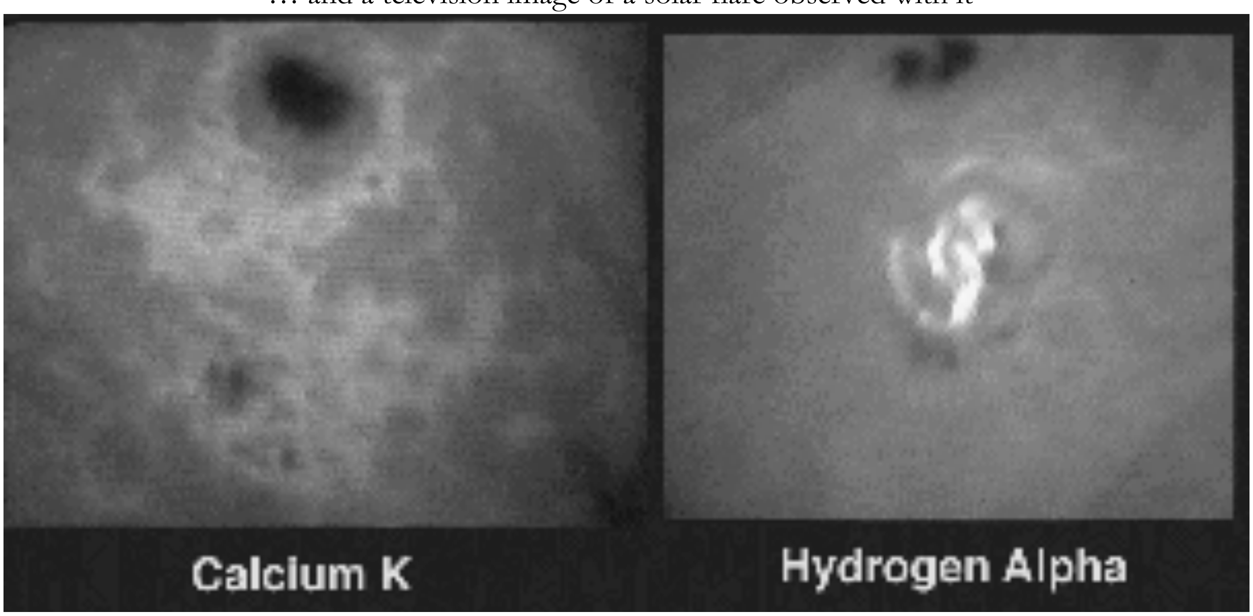
DAYTIME LABORATORY EXPERIMENTS





The SBO Heliostat …

… and a television image of a solar flare observed with it



MOTIONS OF THE EARTH & MOON

Name: ______________________________

Pre-Lab = please complete *before* coming to lab

Read briefly through the entire lab. Check this box ☐ to indicate that you have done so. Jot down any questions that occurred to you while skimming through the lab description.

1) What is the difference between latitude and longitude? What is the latitude of Boulder?

2) At noon today in Boulder, generally where should you look to see the Sun in the sky? (In the northern, eastern, southern, or western part of the sky?)

3) How can the Sun and the Moon have the same apparent angular size in our sky?

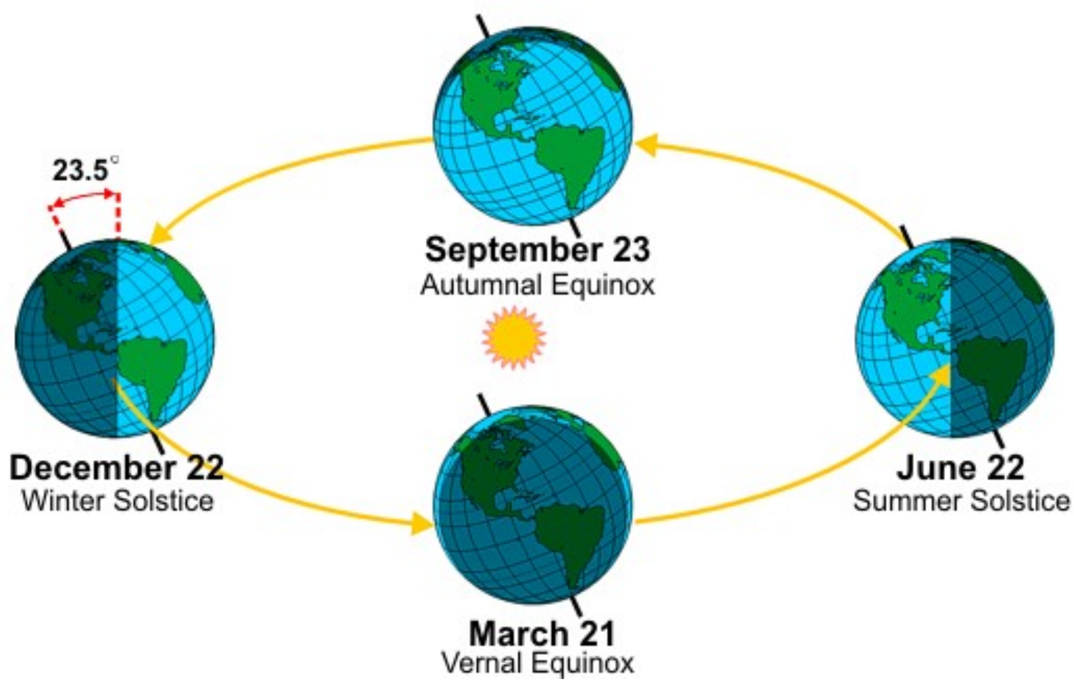
MOTIONS OF THE EARTH AND MOON

SYNOPSIS: The goal is to investigate some of the apparent motions of the Sun & Moon in the sky. You will then use these results to explain why Earth experiences both seasons and phases of the Moon.

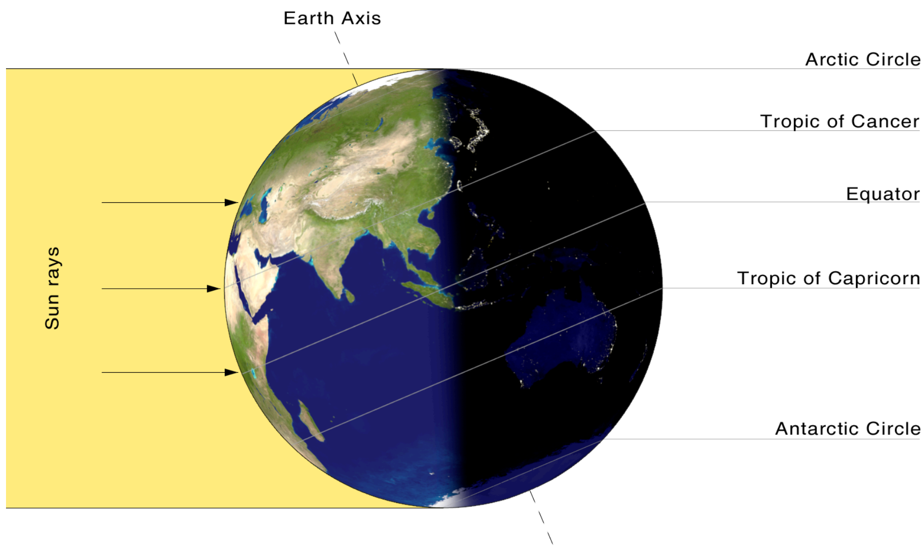
EQUIPMENT: A globe of the Earth, a bright light, foam Moon ball, space to move around.

Part I. The Annual Motion of the Sun

The position of the Sun in the sky appears to change throughout the year as the Earth orbits around the Sun. *This motion is not to be confused with the daily motion of the Sun (rising and setting).* If you think of the plane of the Earth's orbit -- the ecliptic plane -- being horizontal (parallel to the ground or table top) then the Earth's spin axis does not point directly upwards but is tilted 23.5º. This tilt is fixed in direction in space (always pointing towards Polaris) but as the Earth orbits the Sun, the tilt affects the angle of the North pole to the Sun (as in the diagram below). As a consequence, different parts of the Earth receive different amounts of sunlight depending on where the Earth is in its orbit. (Note: The Earth's orbit is nearly circular, but it appears very non-round in this diagram only because we are viewing at an angle across the plane of the Earth's orbit.)



(Image courtesy of NOAA)



(Image courtesy of User:Blueshade / Wikimedia Commons / CC-BY-SA-2.0)

I.1 We know the Sun (like the stars) looks like it rises in the east and sets in the west. Use your globe to determine which direction the Earth spins. When viewed from above the North Pole, does it rotate clockwise or counterclockwise?

The illustration above question I.1 shows the Earth at the summer solstice for the Northern Hemisphere. Position your globe such that Boulder is at noon *and* it is the summer solstice. At noon, the Sun is high in the sky in Boulder and shining (almost) directly down on us, so it is summer for us. Let's look at a few other locations on the globe and compare them to Boulder.

I.2 Find a city somewhere in Asia at the same latitude as Boulder. List the city (and country) that you've chosen. What season is this part of the world experiencing?

I.3 The **Tropic of Cancer** (23.5° north latitude) marks the latitude where the Sun is at the **zenith** (directly overhead) at noon on the summer solstice. Can the Sun ever be seen at the zenith (at noon) from here in Boulder? If so, when? If not, why not? Use your globe to confirm your answer.

I.4 Is there anywhere in all of the fifty United States where you could see the Sun at the zenith at some time of the year? If so, where?

I.5 The **Arctic Circle** lies at a latitude of 66.5° north. North of the Arctic Circle, the Sun is above the horizon for 24 continuous hours at least once per year and below the horizon for 24 continuous hours at least once per year. Find the town of Barrow, Alaska. On the summer solstice at what times will the Sun rise and set in Barrow?

Without changing the orientation (tilt) of the globe, rotate it so it is noon in Australia. Study what is happening "down under" in Melbourne, Australia.

I.6 On the Northern Hemisphere's summer solstice is the Sun in the northern or southern portion of the sky, as seen from Melbourne?

I.7 Is it high in their sky or low?

I.8 What season is Melbourne (and all of Australia) experiencing? Explain.

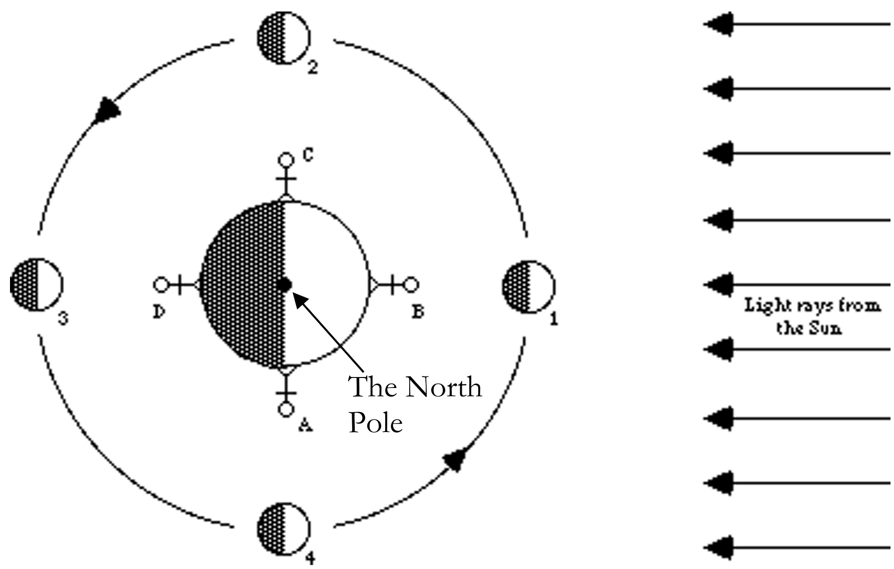
I.9 The **Antarctic Circle** is the southern equivalent of the Arctic Circle. On our summer solstice what time does the Sun rise at the South Pole?

I.10 The **Tropic of Capricorn** lies at 23.5° south latitude. Find Lake Disappointment in the Great Sandy Desert of Australia. Is the Sun ever directly overhead there? If so, when?

I.11 In Australia, do the stars still rise in the east and set in the west? (Hint: Use your globe to verify!) Explain the reason for your answer.

Part II. The Moon's Orbit

Now let's think about the Moon's movement around the Earth. The diagram below shows the overhead view of the Earth and Moon with the Sun off to the right. Depending on where you are sitting in the lab room, the Sun (bright light in the center of the classroom) may be coming from a different direction, but the Moon still orbits the Earth ***counterclockwise*** when viewed from above the North Pole. *For this activity, you can assume that sunrise is at 6am and sunset is at 6pm.*



Note that half of the Moon is always illuminated (just like the Earth); even though it may not appear to be from our view here on Earth.

II.1 How long (roughly) does it take the Moon to orbit the Earth once? How many times does the Earth rotate in that same period?

Imagine that your head is the Earth and that you live on the tip of your nose. Position yourself such that your head is pointing directly at the "Sun" (the light).

II.2 What time is it for the mini-you living on the tip of your nose?

Slowly turn your head around counterclockwise to simulate the daily cycle: sunrise, noon, sunset, midnight, and back to sunrise. This should give you a feel of what direction in space you are looking during the different times of the day (and entertain your labmates).

II.3 Assuming your head is the Earth and the United States stretches from your right eye to your left eye, which eye represents the East Coast?

Hold your "Moon" foam ball out at arm's length. Start with the Moon pointed in the direction of the Sun. While holding the Moon out in front of your nose, spin counterclockwise (to your left). Keep your eyes on the Moon the entire time. Stop turning when the portion of the Moon <u>you can see</u> is more than a sliver but not quite half lit. We call this a **crescent moon**.

II.4 Which side of the Moon (that you can see) is illuminated (right or left)?

II.5 Draw the alignment between the Moon (foam ball), the Earth (your head) and the Sun (light) as seen from above. It may look similar to the diagram at the start of this section, but there should only be one Moon and one person (and it is likely not in any of the places shown on that diagram).

II.6 What is the (approximate) angle between the Moon, the Earth and the Sun: 0°, 45°, 90°, 135°, 180°?

II.7 What time is it for mini-you on the tip of your nose? (This is the time you would see that Moon phase highest in the sky.) Explain your reasoning.

If the illuminated portion of the Moon is getting bigger as it progresses through its phases, we call it "**waxing**." If it is getting smaller, we call it "**waning**."

II.8 Are you looking at a waxing crescent or waning crescent moon?

Continue to turn slowly to your left until the illuminated portion of the Moon you can see, is exactly half lit. This phase is called **first quarter.**

II.9 Why do you think this phase is called a **quarter** moon, when it looks **half** lit?

Continue to turn slowly to your left until the Sun is directly behind you. (You may have to lift the Moon above your head a bit so your head doesn't block the sunlight.)

II.10 What do we call this phase of the Moon?

If you lower the Moon a little you will probably move it behind the shadow of your head. This is a **lunar eclipse,** when the shadow of the Earth (in this case, the shadow from your head) covers the Moon. In Part IV, you'll explore why there isn't a lunar eclipse every time there is a full moon.

Continue to turn counterclockwise, holding the Moon at arm's length. Stop when the illuminated portion of the Moon you can see is no longer full but not yet half illuminated. This is known as a **"gibbous" moon**.

II.11 From your perspective, which side of the Moon is now illuminated?

II.12 Draw the alignment between the Moon, the Earth, and the Sun as seen from above.

II.13 Is this a waxing gibbous or waning gibbous moon? Explain your answer.

II.14 When you are facing directly towards your moon, what time is it for mini-you on the tip of your nose? (This is the time you would see that moon phase highest in the sky.)

Part III. Moonrise and Moonset

In reality, the Moon doesn't orbit as fast as the Earth rotates. So, when you've been moving your head *with* the Moon, you've been making the Moon orbit way too fast. (A day is 24 hours, how long was the Moon's orbit?) Let's investigate moonrise and moonset.

If your mini-you (living on the tip of your nose) had a horizon, it could be simulated as an imaginary plane that cuts down through your head and goes through both ears. North on the horizon would be the top of your head and south on the horizon would be below your chin.

III.1 Which ear represents the western point of the horizon (your right or left)?

While keeping your arm in the same place it was for QII.14 (i.e. don't move the Moon in its orbit) continue to turn your head to the left (simulating the Earth's rotation). Stop when the Moon is even with your *right* ear.

III.2 The Moon is about to cross your horizon. Is this rising or setting?

III.3 What time is it for mini-you on the tip of your nose?

III.4 Use everything you have learned in this section, draw a picture of the alignment of the Moon, the Earth, and Sun as seen from above for a **waning crescent moon**. Also, draw a mini-you on the Earth that would currently be seeing the **moon rising.**

Even though it is usually regarded as a nighttime object, some phases of the Moon can be seen during the day. Go outside to see if the Moon is visible right now (your TA/LA can help). If it is up (and it is sunny), hold up your foam Moon ball at arm's length in the direction of the real Moon.

III.5 Should the foam Moon ball have the same phase as the real Moon? (If weather and lunar phase permit), does it? Explain your thinking; drawings might help.

Part IV. Solar and Lunar Eclipses

Because the Sun and Moon have roughly the same angular size (as viewed from Earth), it is possible for the Moon to block out the Sun from our view, causing a **solar eclipse**.

IV.1 Hold the foam Moon ball at arm's length and move it through its phases. There is only one phase of the Moon when it is possible for it to block your view of the Sun, causing a solar eclipse. Which phase is this?

IV.2 Now simulate this arrangement with the Moon, the Earth globe, and the Sun. In your lab setup, is it possible for one person to see a solar eclipse and someone far away on Earth to see the Sun not being eclipsed? Explain your answer.

IV.3 The lab model you are using is accurate in some ways and inaccurate in others. In what ways is it inaccurate? How might those inaccuracies affect predictions for solar eclipses?

The Moon's orbit is not quite circular. The Earth-Moon distance varies between 28 and 32 Earth diameters. If a solar eclipse occurs when the Moon is at the point in its orbit where the distance from Earth is "just right," then the Moon's apparent size can *exactly* match the Sun's apparent size.

IV.4 If the Moon were at its *farthest* point to the Earth during a solar eclipse, would its angular size appear bigger or smaller than the Sun as seen from Earth?

IV.5 Draw and/or describe what the Sun would look like from the Earth during a solar eclipse, if the Moon is at its farthest possible distance.

It is also possible for the shadow of the Earth to block the sunlight reaching the Moon, causing a **lunar eclipse.**

IV.6 Once again, move the Moon through its phases around your head and find the one phase where the shadow from the Earth (your head) can fall on the Moon, causing a lunar eclipse. What lunar phase is this? (Or what phase was it *just* before the eclipse?)

IV.7 Is it possible for one person to see a lunar eclipse and someone far away on Earth to see the Moon not being eclipsed? Explain your answer.

Many people mistakenly think that a lunar and solar eclipse should occur every time that the Moon orbits the Earth. This misconception is due to the fact that we usually show the Moon *far* closer to the Earth than it actually is, making it appear that an eclipse is unavoidable. However, the Moon's actual distance is roughly 30 Earth-diameters away, and the Moon's orbit is tilted slightly (~5°) to the ecliptic plane.

Your TA may have toys set up in the lab to demonstrate the Moon's orbit around the Earth. Use these demos to visualize the scale and/or tilt of the Moon's orbit.

Position the foam Moon ball in the full phase and hold it (as close as you can to) a properly scaled distance from the globe of the Earth (30 Earth-diameters away).

IV.8 Compared to when you held the foam Moon at arm's length from the globe Earth, is it easier or more difficult to align the Moon, Earth, and Sun to cause a lunar eclipse?

IV.9 First using the foam Moon ball, and then by making a sketch below showing the view of the Sun-Earth-Moon system from the side (edge-on), show how the Earth's shadow can miss the Moon, so that a lunar eclipse does not occur.

IV.10 Calculate the radius of Earth's shadow at the orbital distance of the Moon during a lunar eclipse. Here are some things you will need:

- Radius of Earth: 6400 km - Earth-Sun distance: 1.5×10^8 km
- Radius of the Sun: 696,000 km - Earth-Moon distance: 380,000 km
- A good diagram that takes advantage of carefully labeled similar triangles

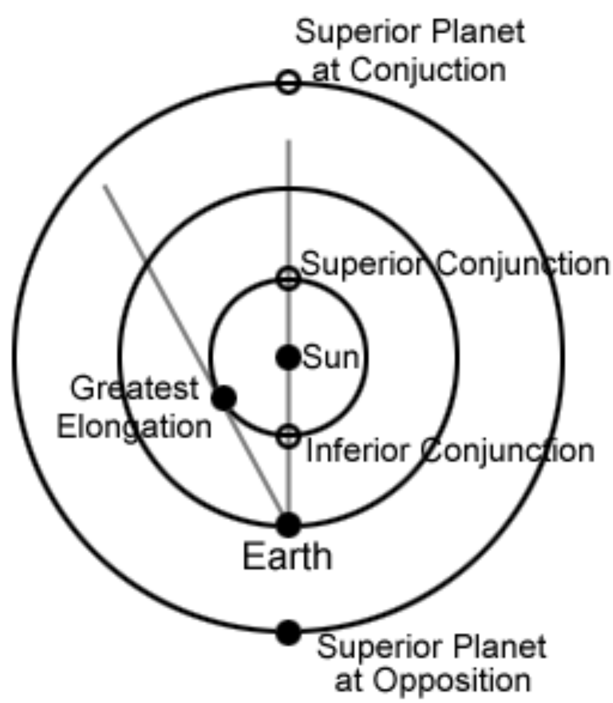
IV.11 Make a drawing of a lunar eclipse (seen from Earth), comparing your calculated size for Earth's shadow to the 1700km radius of the Moon.

Name: ___________________________

MOTIONS | Post-Lab = please hand in at the start of next lab

Five planets are easy to find with the naked eye: Mercury, Venus, Mars, Jupiter, and Saturn. Like the Sun and the Moon, the planets appear to move slowly through the constellations of the zodiac. The word planet comes from the Greek for "wandering star." However, although the Sun and Moon always appear to move eastward relative to the stars, the planets occasionally reverse course and appear to move westward through the zodiac. Because Venus and Mercury are always observed to be fairly close to the Sun in the sky (we only see them near sunrise or sunset), they must have orbits interior to the Earth's. We call them inferior planets. Likewise, because Mars, Jupiter, and Saturn can be seen in the middle of the night, they must be further from the Sun than the Earth. These planets are called superior planets.

We can use geometry to define certain positions for the planets in their orbits. Take a look at the diagram. When a planet appears lined up with the Sun as seen from Earth it is called **conjunction**. Note that the inferior planets have two configurations called conjunction: superior conjunction when the planet is behind the sun as seen from Earth and inferior conjunction when the planet is between the Sun and Earth.



1) Although Venus is as close to Earth as possible at inferior conjunction, it would not be a good time to observe the planet. Why not?

2) The angle between the Sun and a planet as seen from the Earth is called **elongation**. When a planet is at conjunction, its elongation is 0°. When elongation is at its maximum value, we say that the planet is at greatest elongation. At greatest elongation, what is the Earth-Venus-Sun angle (how far apart in angle are the Earth and the Sun, seen from Venus)?

3) Calculate the angle of greatest elongation for Venus. That is, what is the Venus-Earth-Sun angle (how far apart in angle are Venus and the Sun, seen from Earth)? *Use whatever information you need from the tables at the back of The Cosmic Perspective.*

4) Repeat the same calculation for Mercury, instead of Venus.

5) From these calculations, explain why Venus is visible in the night sky more often and for longer periods of time than Mercury.

ERATOSTHENES CHALLENGE

Name: ______________________

Pre-Lab = please complete *before* coming to lab

Read briefly through the entire lab. Check this box ☐ to indicate that you have done so. Jot down any questions that occurred to you while skimming through the lab description.

1) If the Sun was directly above the town of Syene at solar noon on the summer solstice, at what latitude must Syene be located?

2) In your own words, explain the difference between **precision** and **accuracy**.

3) Estimate the order-of-magnitude distance from Sommers-Bausch Observatory to Baseline Road. This can be a **very** rough estimate, and you will not be penalized if you are wrong.

Name: ______________________________

THE ERATOSTHENES CHALLENGE

SYNOPSIS: The purpose of this observing project is to measure the circumference of the Earth in your paces and then in yards and miles using the ancient methods of Eratosthenes. We will use the results to have you discuss why measurement errors are *not* mistakes and why systematic errors sometimes *are* mistakes. If you want to make this a more accurate historical re-enactment, we will give you the opportunity to calibrate your paces in the *CU football stadium* (since the ancient Greeks measured distances in *stadia*).

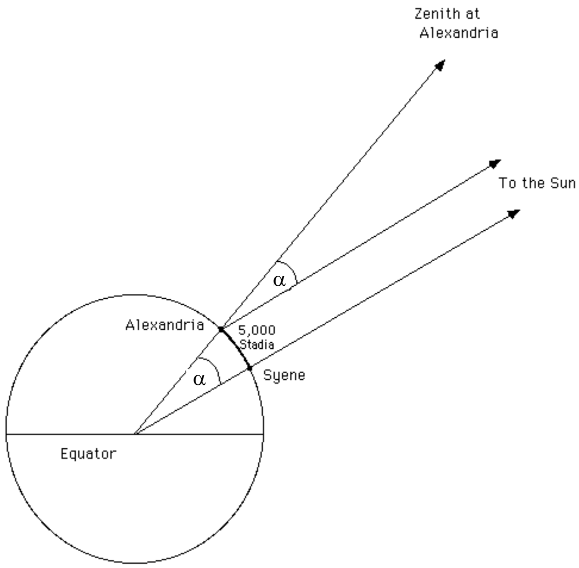
Background: Eratosthenes of Cyrene

Born: 276 BC in Cyrene, North Africa (now Shahhat, Libya).
Died: 194 BC in Alexandria, Egypt

Eratosthenes was student of Zeno (founder of the Stoic school of philosophy), invented a mathematical method for determining prime numbers, and made the first accurate measurement for the circumference of the Earth!

Details were given in his treatise "On the Measurement of the Earth" which is now lost. However, some details of these calculations appear in works by other authors. Apparently, Eratosthenes compared the noon shadow at Midsummer (June 21st) between Syene (now Aswan on the Nile in Egypt) and Alexandria, 500 miles to the North on the Mediterranean Sea. He assumed that the Sun was so far away that its rays were essentially parallel, and then with a knowledge of the distance between Syene and Alexandria, he gave the length of the circumference of the Earth as 250,000 stadia (1 stadium = the length of a Greek stadium).

We still do not know how accurate this measurement is because we still do not know the exact length of a Greek stadium. (Units matter!) However, scholars of the history of science have suggested a value for the stadium and estimate that Eratosthenes' measurement was 17% too small. Unfortunately, in Renaissance times, the length of a Greek stadium was under-estimated as well, yielding an even smaller circumference for the Earth. This small value led Columbus to believe that the Earth was not nearly as large as it is, so when he sailed to the New World, he was quite confident that he had sailed far enough to reach India.



This figure shows how Eratosthenes made his measurement. He had heard that on the **summer solstice** the Sun at noon stood directly over Syene, at the **zenith**, so that the Sun's light penetrated all the way down to the bottom of a well at Syene casting no shadow. Eratosthenes measured the angle of the Sun off the zenith (called the **zenith angle**; angle "α" in the figure) from Alexandria on that same day. (Unfortunately, his measurement of α was ~6% too small.) As shown in the figure, α is also the difference in **latitudes** of these two locations.

The angle α is to 360 degrees (a full circle) as the distance between Alexandria and Syene is to the full circumference of the Earth. Eratosthenes had a measurement for the distance between Syene and Alexandria of 5000 stadia. Mathematically:

$$\frac{\alpha}{360^\circ} = \frac{5{,}000 \text{ Stadia}}{\text{Circumference of the Earth}}$$

and so (rearranging):

$$\text{Circumference of the Earth} = \frac{360^\circ}{\alpha} \times 5{,}000 \text{ Stadia}$$

What do we need to know to make a modern "Eratosthenes measurement"?

We need to know the equivalent of the two measurements Eratosthenes had:

1. The difference in latitude between two locations on Earth.
2. The difference in distance between these two locations in an **exactly** north-south direction.

Eratosthenes measured #1 and had obtained from others a value for #2. We will measure #2 (in paces, then in yards and miles) and obtain a value from others for #1.

Conveniently, we have two nearby locations with well-known latitudes.

- Location 1: When Colorado was surveyed in the 1800s, Baseline Road was determined to be at precisely 40 degrees North Latitude.
- Location 2: More recently than that, an astronomical measurement at the **Sommers-Bausch Observatory** (SBO) 24-inch telescope (located here at CU just north of Baseline Road) determined the latitude of SBO to be:
    - **40.00372 degrees (+40° 00' 13.4") North**

Part I – Measuring the Distance to Baseline Rd.

Unfortunately, in recent years due to traffic control necessity, the course of Baseline Road has been altered just south of SBO. As shown in the photograph from space overleaf, Baseline curves gently north between Broadway Blvd and 30th Street. The white line is our best estimate for exactly 40 degrees North latitude based upon the course of Baseline Road east and west of this bend. Perhaps realizing that they had altered a geographically (and astronomically) important landmark, the city of Boulder (or maybe RTD?) has painted a red line on the sidewalk near the bus stop in order to mark the exact location of the 40^{th} parallel. (Notice how it splits the rock to the east.)

To measure the north-south distance between SBO and Baseline Road, you will need to:

- Plan a route.
- Walk that route, carefully counting your paces.
- Calibrate the size of your paces using a tape measure.

Each member of your lab group must make these measurements (both pacing between SBO and Baseline and "calibrating" their paces by stepping off 100 yards). Each participant will then use the Eratosthenes Equation to determine how many paces you would need to walk to get all the way around the Earth. By calibrating your paces you will then determine the circumference of the Earth.

Do not use your smart phones to accomplish this task. Eratosthenes did not have a phone.

That's IT! That's all we are going to tell you, but if you need help be sure to ask the LAs or TA for some pointers. Each individual in each group must make their own measurements using the method agreed to by the group. Good luck. Keep thinking and stay safe! Especially when crossing Baseline and other streets…the cars do not know that you are conducting an historical reenactment

.East

North

South

West

I.1 Describe the route you took. You may want to mark your route on the map above as part of your description, although *you must still explain your route in words.*

I.2 There are many possible routes you could have taken. How did your group decide on this particular route? What procedures did you use to make sure your measurements were as accurate as possible?

Part II. Explaining your Measurements

The central equation for this lab is $$\frac{\alpha}{360^\circ} = \frac{Distance}{Circumference}$$

II.1 Explain in words why the fraction on the left side of the equation ($\alpha/360^o$) must equal the fraction on the right side of the equation (Distance/Circumference). Include a drawing.

II.4 Calculate the circumference of the Earth in meters and kilometers, based on your count of the number of paces between SBO and Baseline. Show all your work.

II.5 Write down what the other members of your group calculated for the circumference based on their paces. *You do not need to show the calculations again.*

Name:________________________ Circumference:_________________________

Name:________________________ Circumference:_________________________

Name:________________________ Circumference:_________________________

Name:________________________ Circumference:_________________________

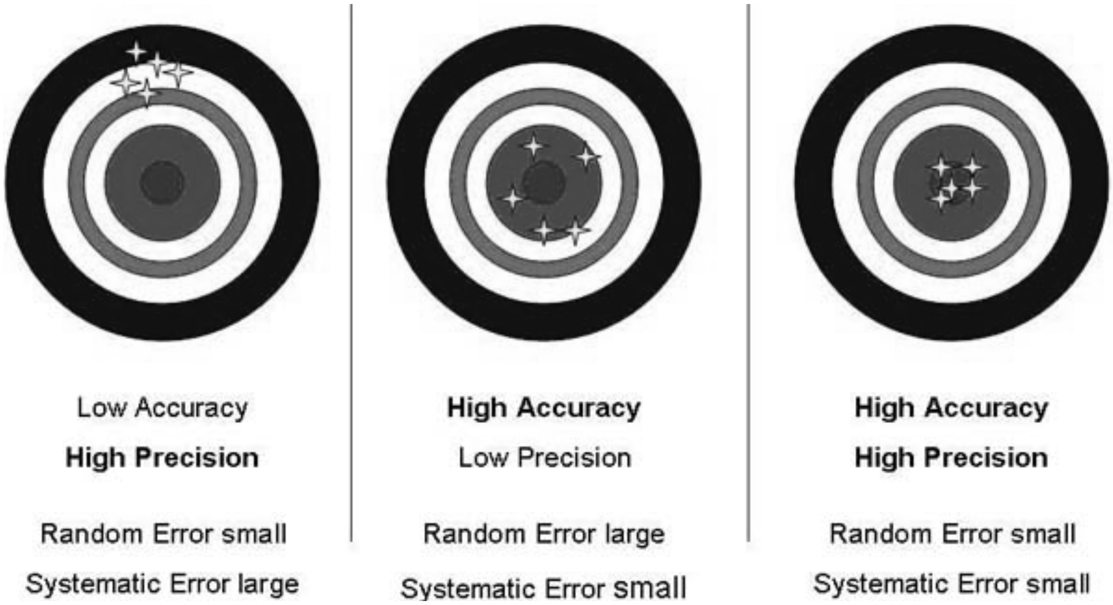
II.6 Calculate your group's **average** value for the Earth's circumference. Show all your work.

Part III. Errors

Webster's dictionary defines **error** as "the difference between an observed or calculated value and the true value". We don't know the true value; otherwise there would be no reason to make the measurement. We wish our measurements to be both ACCURATE and PRECISE.

Accuracy relates to how closely the results of the experiment are to the true result. Thus, accuracy speaks to whether our chosen methods actually work to allow a measurement of the quantity we seek to determine, whether all assumptions have been accounted for and whether these assumptions do not compromise the measurement. Errors in setting up an accurate experiment are called **systematic errors**, and more and more precise measurements cannot reduce these types of errors.

Precision, on the other hand, refers to the actual measurement process itself. Greater precision in measurement can be accomplished by using a more accurate measuring device or by repeating measurements several times. Uncertainties in precision are called **measurement** or **random uncertainties** and repeated measurement can reduce these uncertainties (e.g., independent measurements by equally precise measuring tools or people) but never eliminate them. However, be warned, precise measurements do not yield an accurate result if the experimental setup is inaccurate; i.e., systematic and measurement errors are independent of one another and both must be dealt with to obtain the best value for the true result.



Any scientific measurement has inherent uncertainties and errors (precision in measurement and errors in experimental setup) that limit the ultimate precision of the result. All scientific experiments have these limitations, which must be quoted with the result (for example, even political polling reports results and uncertainties… 54% with an uncertainty of 3 percentage points…but beware, systematic errors are not reported and can be much larger in some cases – what if only men were polled? only rich people were polled?) In this experiment, think about the experimental setup, the specific methods that you and your group employed and the uncertainties and errors which may have limited the ultimate precision of your result.

III.1 Calculate the **standard deviation** of your group's measurements. This is one estimate of the precision of your group's average measurement. Show all your work.

III.2 Typically, an experimental result is listed as: [value obtained] ± [precision] (for example, this might look like 25,000 km ± 1000 km for the circumference of the Earth). Using your group's average as an estimate of the value and the standard deviation as an estimate of your precision, list your group's result in this format

Circumference = ________________ ± ________________

III.3 Another useful quantity to report is the **fractional precision** of your measurement. This is your precision, expressed as a fraction or percentage of the value. Are you off by 0.01%, off by 1%, or off by 100%? Different values could mean qualitatively different interpretations! What is your group's fractional precision?

III.4 Share your group's data with your lab TA/LA, so they can put them into a spreadsheet. Look carefully at the results of the whole class. What do you notice about the precision and the accuracy of your measurements as a whole?

III.5 Random errors often tend to average together to make a more precise measurement. If we repeated your measurements millions of times, sometimes it'd be a little too high and sometimes a little too low, but they'd average out to be correct. Think back on how you got your data – what kinds of errors do you think might have been **random**, and how big are they? *(List at least two.)*

III.6 Systematic errors are ones that we cannot average away. No matter how many times we repeat the measurement, we will be consistently and systematically inaccurate. Think back on how you got your data – what kinds of errors do you think might have been **systematic**, and how big are they? *(List at least two.)*

Name: ______________________________

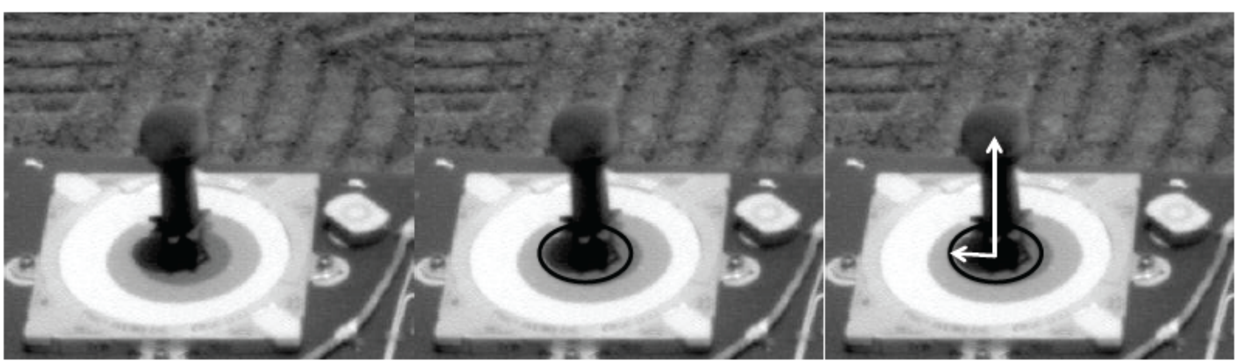
ERATOSTHENES | Post-Lab = please hand in at the start of next lab

Eratosthenes on Mars

We can perform the same experiment on Mars that Eratosthenes did on the Earth. The Mars rovers, Spirit and Opportunity, landed on the Martian surface in early 2004. Each rover was built with a Sundial that was used also as a photometric calibration device, but we will use them as gnomons, measure their shadows, and calculate the angle of the Sun at the two locations of the Mars rovers. Spirit is located at (14.57° S, 175.48° E) and Opportunity at (1.95° S, 354.47° E) and they are 743.7 km apart in the N-S direction. The images in figures below were taken on the same day (March 31, 2004) very close to local noon, so the Sun is near zenith and the shadow from the gnomon is almost completely in the N-S direction. The rovers were not oriented in the N-S direction, so the direction of the shadow in the images is not the same.

(a) Gnomon (b) Projected ring of constant radius (c) Projected shadow length

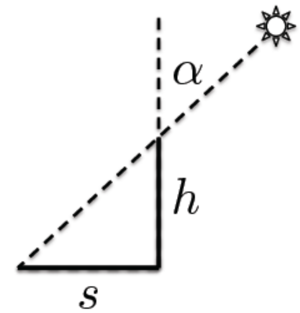
Sundial on Spirit rover at 12.45 PM local Mars time on March 31, 2004



(a) Gnomon (b) Projected ring of constant radius (c) Projected shadow length

Sundial on Opportunity rover at 11.47 AM local Mars time on March 31, 2004

1) Determine the angle of the Sun from zenith α for each rover. You can use the height of the gnomon *h* and the length of the shadow cast *s*, measured with a ruler from the arrows in the above images. Show your work.



α_{Spirit} = ____________________ $\alpha_{\text{Opportunity}}$ = ____________________

2) Based on your estimated angles and the information on the previous page, calculate the circumference of Mars.

3) Explain how you might make your estimate of Mars' circumference more **precise**, compared to the procedure you just completed.

4) Explain how you might make your estimate of Mars' circumference more **accurate**, compared to the procedure you just completed.

SURVIVOR CHALLENGE

Name: ____________________________

Pre-Lab = please complete *before* coming to lab

Read briefly through the entire lab. Check this box ☐ to indicate that you have done so. Jot down any questions that occurred to you while skimming through the lab description.

1) In your daily life around Boulder, what reference points do you often use to estimate which direction you are facing or where you are? *(There's no wrong answer here!)*

2) Have you ever used the sky to navigate or figure out where you are? *(Likewise, no wrong answer!)*

Name: ____________________________

SURVIVOR CHALLENGE

SYNOPSIS: You will be taken to a random location somewhere on the northern hemisphere of the Earth at a random time of the year. Your goal is to figure out where you are located and what month of the year it is using only basic observations of the sky and your ingenuity. You will be graded primarily on the quality of your plan, not your answers.

Real science involves creativity and imagination to figure things out. Part of the goal of this lab is to provide a better understanding for how scientists actually work.

EQUIPMENT: This lab write-up, a pencil, and a white board and marker to share with your group for brainstorming. Globes, laser pointers, and red flashlights are available to share among the whole class.

TOOLS AT YOUR DISPOSAL: Since you will effectively be in control of the planetarium, here are some tools and superpowers you can utilize to help you figure out your time and location on Earth:

- You will have a chance to observe the Boulder sky on the mystery date for 24 hours before you are taken to a mystery location.
- You have an imaginary watch that is always set to Boulder time during this whole activity. At any time you can ask the planetarium operator "What time is it in Boulder?"
- You may request information from the planetarium operators, for instance, "stop turning the sky at sunset and tell us the time when the sun sets." Or, "stop the sky at noon."
- Once you are in the mystery location, you can observe the sky for 24 hours (and stop the sky as you wish). You may also observe any day one month earlier and one month later than the mystery date.
- You DO NOT have superpowers to be able to see lines, grids, degrees, cardinal directions, or other projections on the sky. You must use what you've learned in class and your ingenuity to discover some of these things for yourself!

Planning, Observing, Interpreting

You will be given some time to discuss with your group what kinds of observations you want to make in the planetarium to try to figure out where and when you are. Once you are ready, you will have time in the planetarium to play and explore, working with the planetarium operator. Have fun!

Observations we will make and what they can tell us:

For each item give your answer, what evidence you used *and* how accurate you think it is.

Our **latitude** is ________________ with an approximate uncertainty of ______________
We know this because:

Our **longitude** is ________________ with an approximate uncertainty of ______________
We know this because:

The **date** is ________________ with an approximate uncertainty of ______________
We know this because:

Where do you think you are? (Name or describe the place and what the climate is like.)

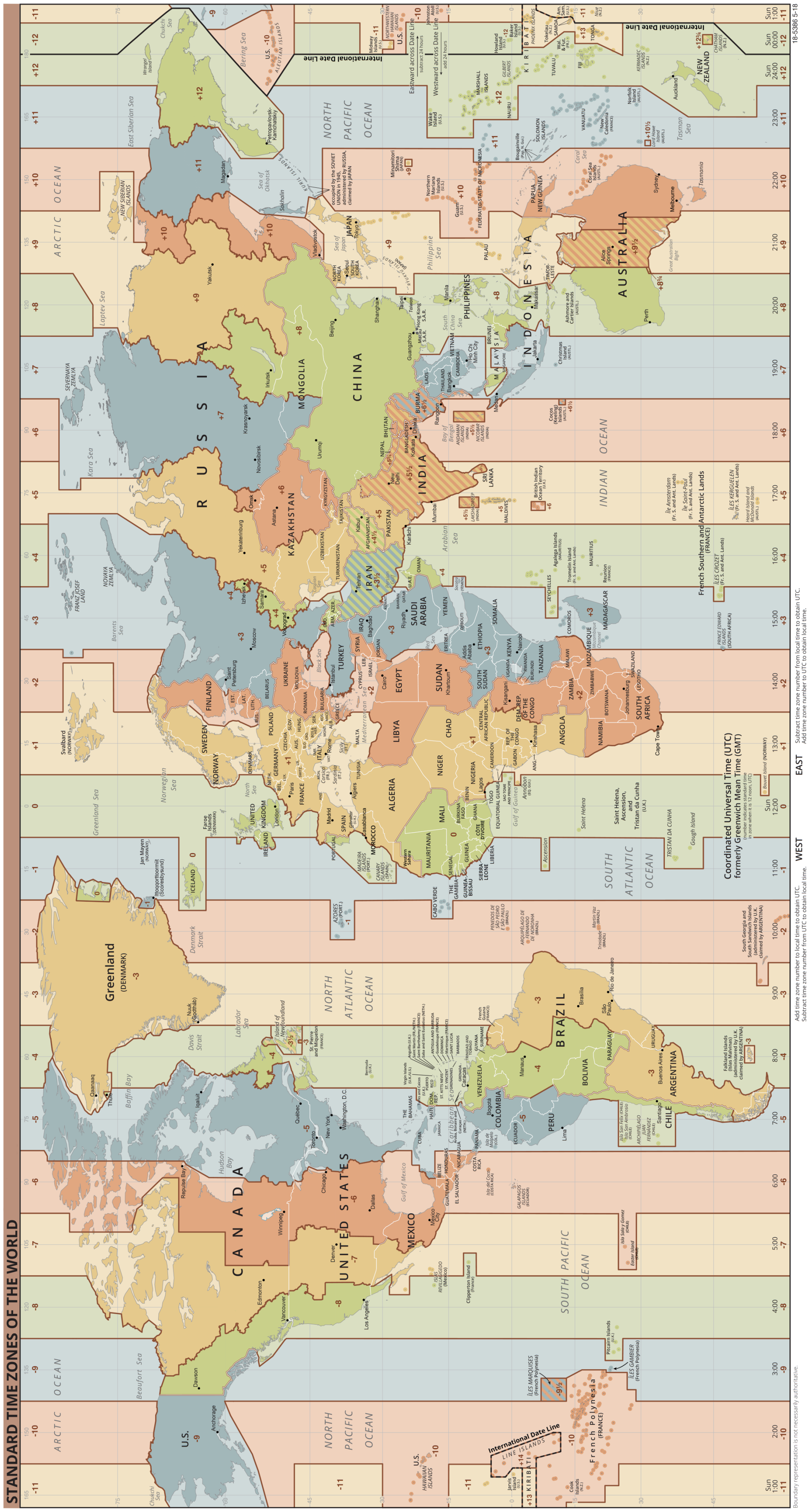


Image courtesy of User:TimeZonesBoy / Wikimedia Commons

Name: ______________________________

SURVIVOR | Post-Lab = please hand in at the start of next lab

1. In the planetarium, you had magically had access to a lot of useful data. Let's imagine you're dropped in a random location on the actual Earth. Somehow all the mapping and GPS apps on your phone are broken, but you can still call or text your friends all over the world. Describe a rough plan of how you might figure out where you are.

2. **Metacognition:** We are always trying to improve the ASTR1030 labs and value your input on where they are succeeding and where they could be improved. Please briefly describe, what did you like and/or not like about this lab? What main idea(s) do you feel you learned from this lab?

LIGHT AND COLOR°

Name: ____________________________

Pre-Lab = please complete *before* coming to lab

Read briefly through the entire lab. Check this box ☐ to indicate that you have done so. Jot down any questions that occurred to you while skimming through the lab description.

1) Explain what the spectroscope does, in general terms (you do not need to explain the quantum mechanics of how the diffraction grating works, just what it does).

2) Explain why a blue shirt looks blue when viewed in white light. What happens to different wavelengths of light interacting with the blue shirt?

3) The y-axis of the spectrum plots throughout this lab are labeled as "intensity." What kinds of different units could "intensity" be measured in? *(There are lots of correct answers here.)*

LIGHT AND COLOR

SYNOPSIS: The only way for astronomers to study distant objects is to examine the light we receive from them. What can we learn about astronomical objects from their light? This lab walks through how we think about light interacting with matter in different ways.

EQUIPMENT: Light bulb, set of 3 filters, spectroscope, colored marbles.

A **spectrum** is the intensity of light at different wavelengths. Normally when we receive light from objects, the wavelengths are all mixed together and we can't tell how much of each wavelength is present. A **spectroscope** allows us to separate the different color components of light, allowing us to examine each wavelength separately. In this way, we can think of a **spectrum** as organized light -- organized according to wavelength.

Part I – White Light & RGB

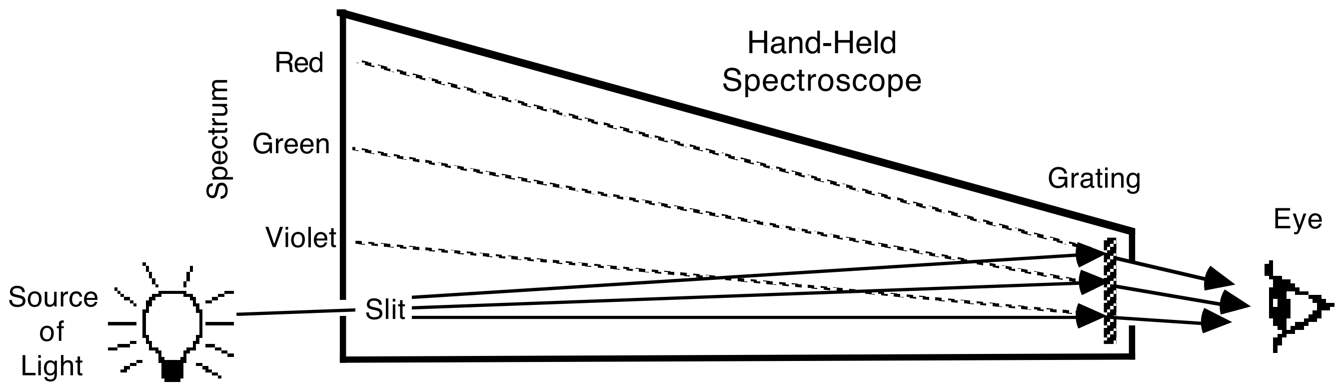
The goal of this section is to become familiar with how white light is a combination of colors. And to learn the "light verbs" that describe all the ways that light interacts with matter – emit, absorb, transmit, and reflect.

Turn on the light bulb.

1.1 List ALL the objects in the room that are ***emitting*** visible light.

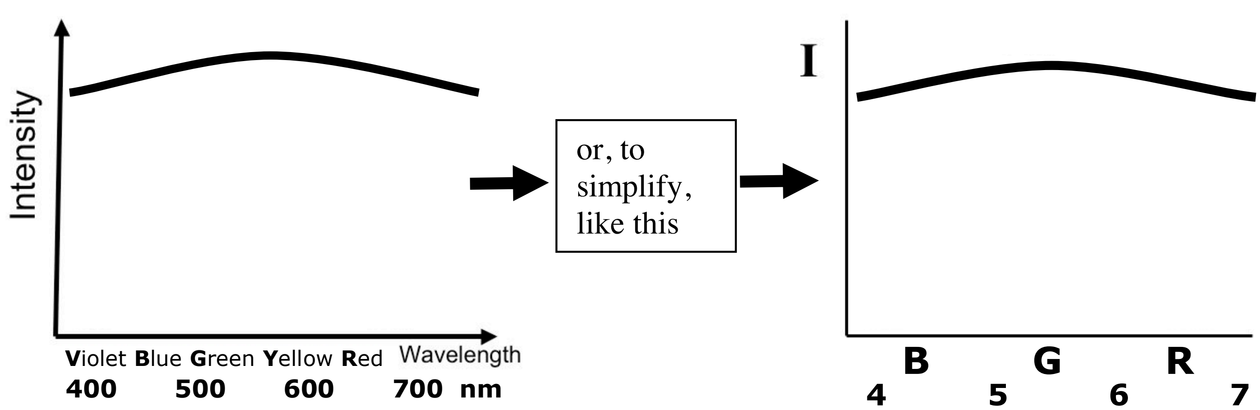
1.2 Can you tell what colors are present within white light by simply looking at the bulb with the unaided eye? → Explain your answer.

A **spectroscope** is a device designed for viewing a spectrum. Light enters the spectroscope through a slit and strikes a diffraction grating made of a material that splits up each color in the light. The grating splits each color apart at a slightly different angle such that each color forms its own separate image of the light from the opening of the spectroscope. A slit is used that the opening to produce narrow images, so that adjacent colors do not overlap each other. The spectroscope has a numerical scale inside of it that measures the wavelength of the light in units of hundreds of nanometers.



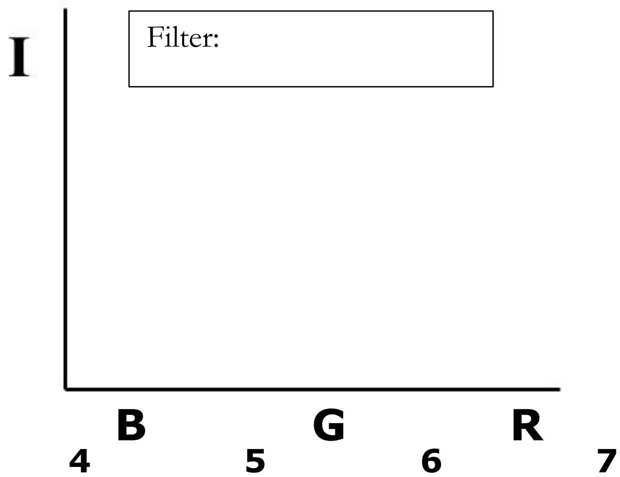
1.3 Use the grating or spectroscope to look at the light bulb. Describe what you see.

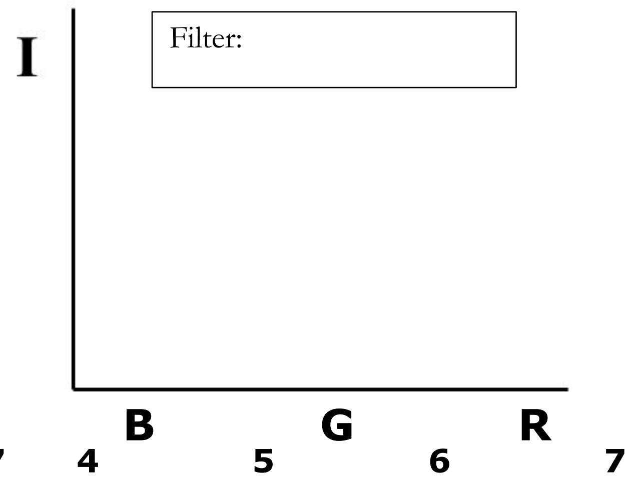
Converting the brightness you see to quantitative intensity, we can sketch the visible portion of a spectrum of white light like this:



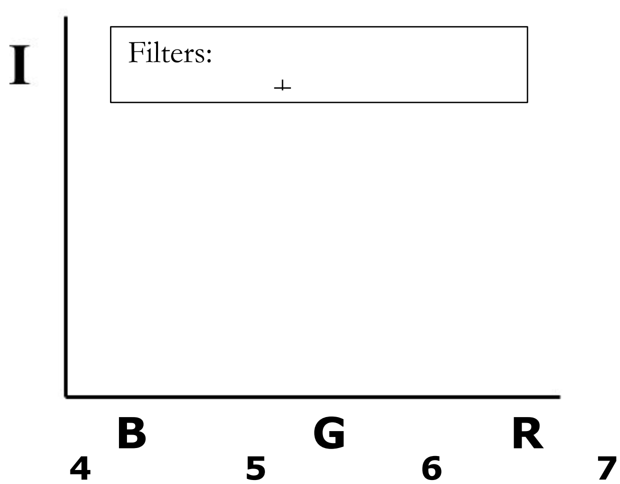
where all colors have a similar value for intensity because all colors are present in white light.

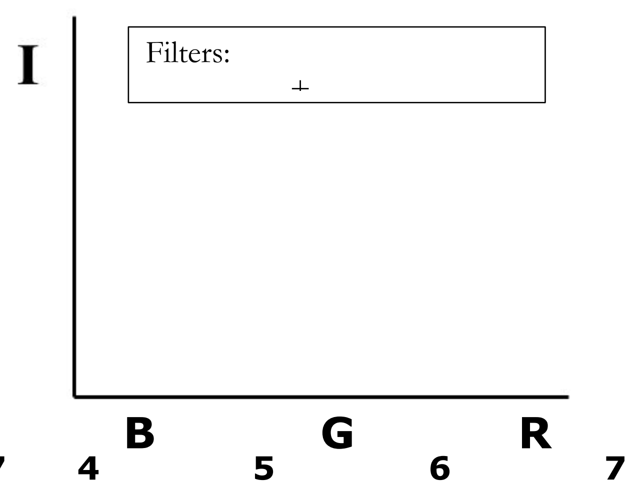
1.4 Using the spectroscope, look at the light through the various filters. Draw the spectra of the light being ***transmitted*** - "let through" - by the different filters. *Make sure the filter is covering the slit of your spectroscope.*



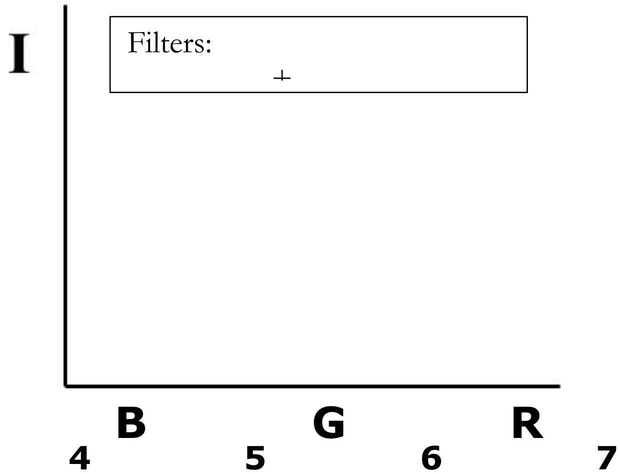


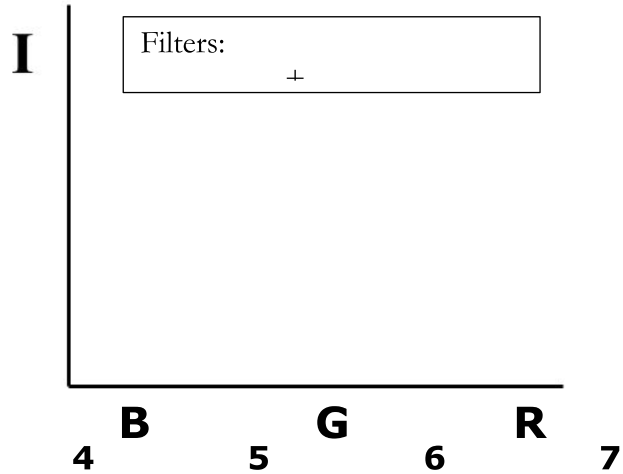
1.5 What happens when the light passes through 2 filters? **Before you try it…**on the graphs below, **predict** what the intensity curve would look like for 2 combinations of the filters. (C*hoose 2 different combinations.)*





1.6 Now actually try out your 2 combinations from above and sketch your results below.





1.7 • Do your predictions from 1.5 match your findings from 1.6? → If not, explain why.

• Explain, in general, what using a filter does to the light coming into the spectroscope.

1.8 • Now think about the white light hitting your shirt/sweater.

• What color(s) is your shirt/sweater?

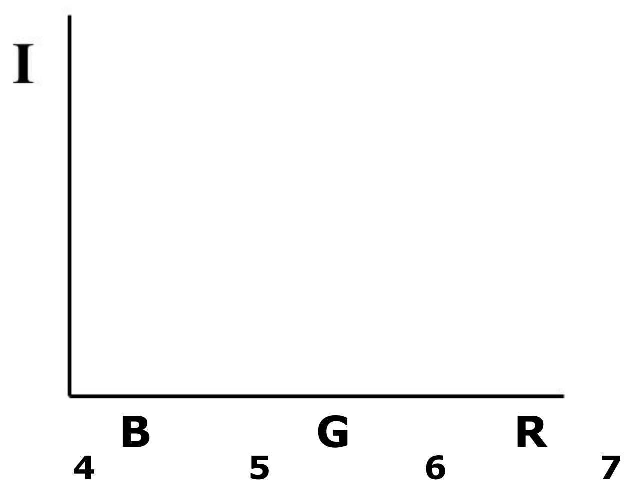
____

• What colors are ***reflected*** by your shirt/sweater? ____________________

• What colors are ***absorbed*** by your shirt/sweater?

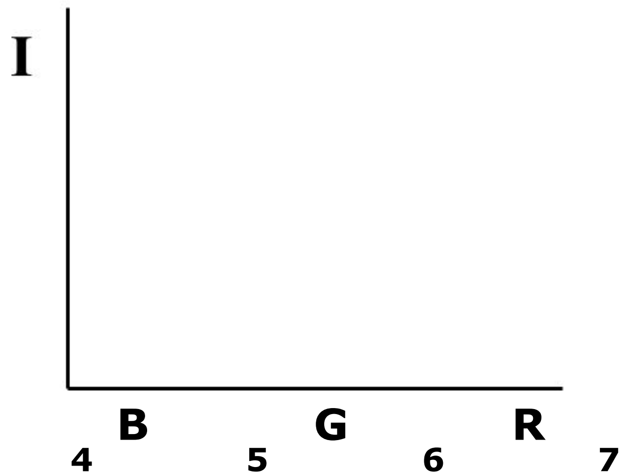
____

• On the graph to the right, draw what you think the spectrum of light **reflected** by your shirt/sweater is.



1.9 On the graph below draw **and label** what you think the reflected spectrum of the following 3 shirt colors would look like (you should have 3 different lines):

- Black
- Gray
- White



Part II – The RGB Room

The goal of this section is to explore the relationship between the color of an object, and how it is affected by the wavelengths of light available.

Up the first stairs you come to as you walk into the Observatory, on the left, there is a room with Red, Green and Blue spotlights. You can control the amount of each color light by sliding the 3 sliders on the white box.

2.1 In this section you will use marbles to study color. *Predict* what will happen if you look at a handful of colored marbles under a single colored light?

2.2 Try it. Take a small handful of marbles - make sure you have an assortment of colors. Now turn off all light except the red light and try to sort the marbles into piles according to color.

- Is it easy to sort out all the colors? → Explain why it is easy/hard to sort the marbles with only the red light on in terms of what is happening with light.

2.3 Once you've sorted the marbles under the red light, turn the white room light back on.

- Did you sort them all correctly?
- List a few colors you mistook for each other. → Explain why these colors were mistaken for each other.

2.4 Mix the marbles up again and now try to sort them using the green light only. Once they're sorted, turn on the white room light and check you work.

- Did you sort them all correctly?
- If not, did you mix up the same colors using the green light only as the red light only?

2.5 Keeping the same sorted piles you made in 2.4, turn off all the lights except the red light and green light.

- With both the red and green light on is it easier to tell which colors have been sorted incorrectly?
- Explain why it is/is not easier to sort with both the red and green lights on in terms of what is happening with light.

2.6 Now turn on all three colored lights.

- Are you able to clearly distinguish all the marble colors now?

→ Explain why or why not.

2.7 Now turn off all lights except the white room light and find the white sheet of paper with the red shape on it.

• **Predict -** will you be able to see the red shape on the paper under the red light only?

2.8 Now turn off the white room light and look at the sheet of paper under the red light only.
• Was your prediction from 2.7 correct? → Explain what is happening with light in this situation.

Before you leave, turn off all the colored lights and turn on the white room light.

Part III – The Yellow Room

Be sure to bring your spectroscope! The goal of this section is to explore the relationship between the appearance of an object and how it is affected by the wavelengths of light available.

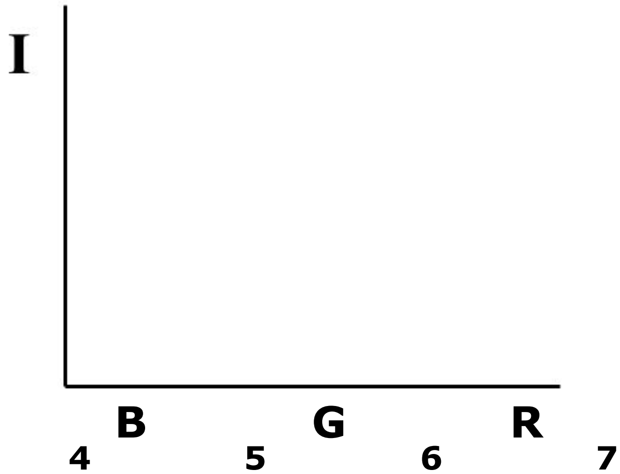
The yellow room is a room illuminated by different types of lights that all emit a very similar yellow color but do it very differently.

3.1 Look at the various objects in the room in each of the following two lights without the spectroscope – *be sure to turn off the room light and only open each light's door one at a time*. Describe how a few objects look under each light

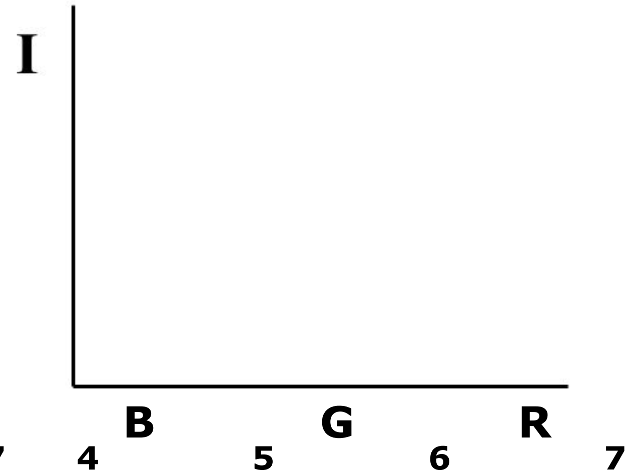
Incandescent: Fluorescent:

3.2 Based on your observations in 3.1, **predict** what each of the lights' spectra will look like **before** you look through the spectroscope. Draw your prediction of the spectra for each bulb below. (You may also use words to clarify your prediction below.)

Incandescent:



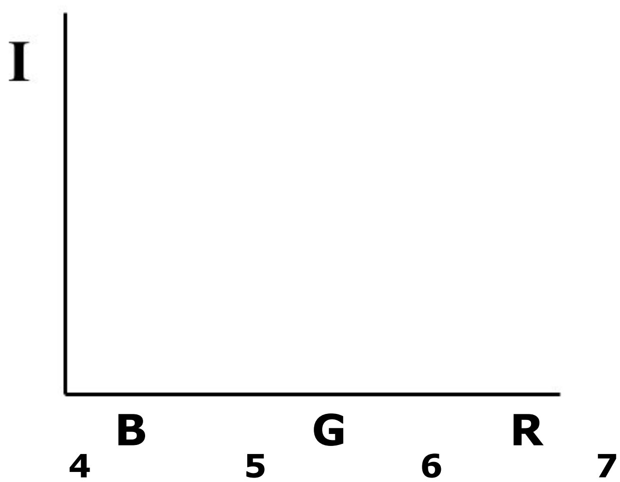
Fluorescent:



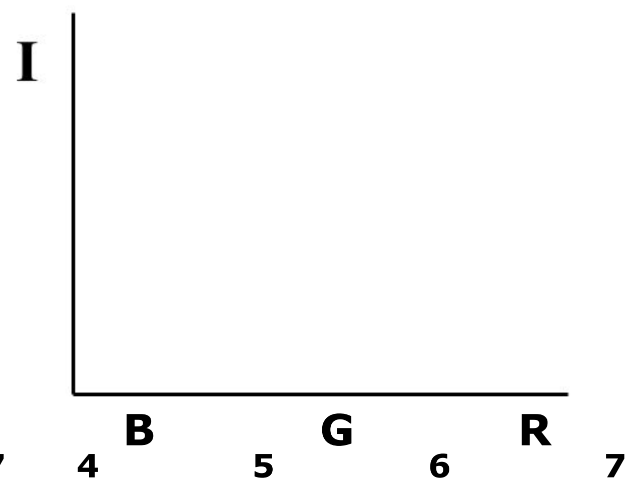
3.3 Now use your spectroscope, to examine the spectrum of each light.

- Were your predictions correct?
- If not, correct your descriptions below on the graphs.

Incandescent:



Fluorescent:



3.4 The next light you will examine is a sodium lamp. The sodium lamp emits yellow light at only <u>one</u> wavelength – this is known as "monochromatic light."

- **Before you look, predict** what the multi-colored sweater will look like under the sodium lamp.

3.5 Close all the light doors and make sure the room light is off. Open the sodium light door.

- How do things look under the sodium light?
- How did the color of the multi-colored sweater change?

3.6 Guess what color some of the other objects are. Then turn on the room light.

- Did you guess all the correct colors? → Explain why or why not.

3.7 Look at the sodium lamp through your spectroscope.

- Explain how the light emitted by the sodium lamp is changing the color you see on the multi-colored sweater and the colored objects in the room.
- Use what you observed through the spectroscope to aid your explanation.

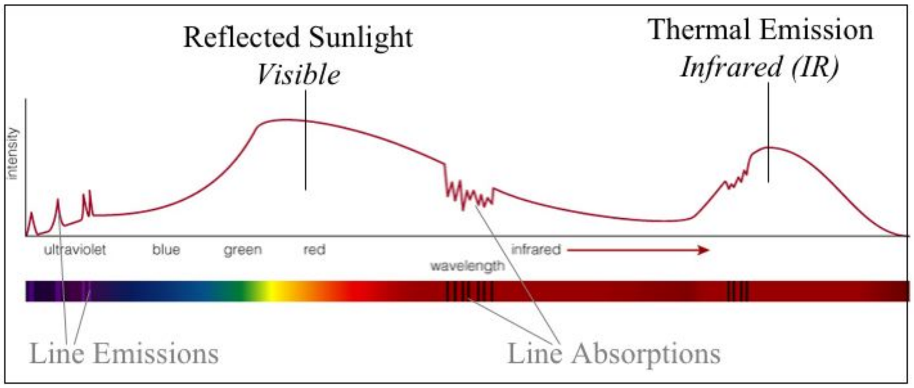
3.8 • Choose one object in the room and explain why it looks the way it does in each of the 3 lights – be sure to say which object you chose in your answer.

Before you leave, make sure the sodium lamp is uncovered and all other lights are covered.
PLEASE DO NOT TURN OFF THE SODIUM LAMP.

Part IV – Planets and People at Infrared Wavelengths

Visible light is only one of several types of light that we can study. The goal of this section is to examine another type of light called infrared (IR) light. By using a camera designed to display IR light, you will discover how IR light contains information that otherwise cannot be seen with visible light alone.

In addition to reflecting visible light, planets also emit infrared, or IR light. The figure the right shows what the spectra of a planet looks like. In this section we'll explore some of the properties of IR light.



4.1 Do you think any of the objects in the room are emitting infrared light?

- Which one(s)?

4.2 Check out the IR camera and TV monitor.

- Stand in front of the camera – What parts of the body look warm?
- What looks cold?

4.3 Examine the black plastic garbage bags as well as the clear plastic square.

- Describe what you see. *Be sure to unfold the garbage bag entirely.*

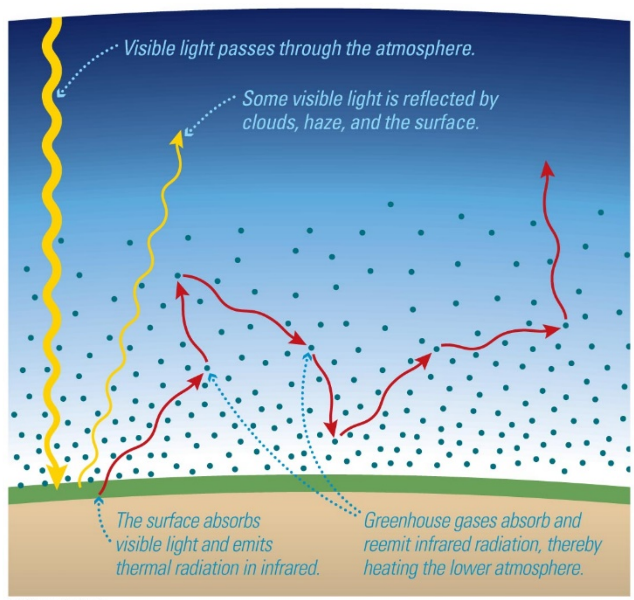
4.4 • Does IR light interact with the black plastic garbage bag in the same way that visible light does? What about for the clear plastic square? → Explain why or why not for both objects in term of what is being **reflected** or **absorbed** and what is **transmitted.**

Garbage Bag:

Plastic Square:

4.5 The phenomena you just examined above is related to the **greenhouse effect** that occurs in Earth's lower atmosphere. (Ask your TA/LA to explain this effect if you have not yet covered it in lecture.)

• Given what you know about the greenhouse effect, which material is the most similar to the Earth's lower atmosphere: the garbage bag or the plastic sheet? → Explain your answer.



Name:

LIGHT AND COLOR | Post-Lab = please hand in at the start of next lab

1. Imagine you've invented the perfect solar panel, that converts 100% of incoming solar radiation and converts it into electricity. Well done, you! What color is it? Explain your answer.

2. Explain how your experiences in lab relate to the phenomenon that wearing a hot on a cold day keeps you warmer.

3. Do incandescent light bulbs emit infrared light? Explain your answer.

D.I.Y. TELESCOPE

Name: ____________________________

Pre-Lab = please complete *before* coming to lab

Read briefly through the entire lab. Check this box ☐ to indicate that you have done so. Jot down any questions that occurred to you while skimming through the lab description.

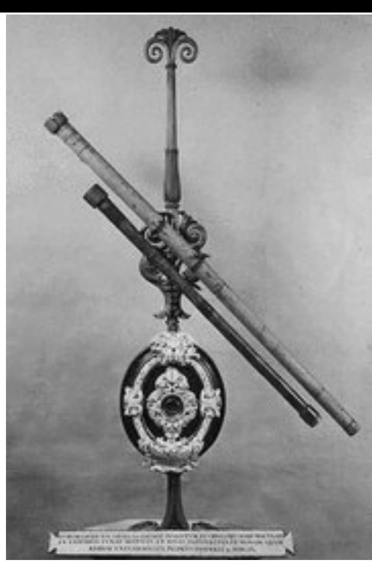
D.I.Y. TELESCOPE

SYNOPSIS: We will be assembling and using refracting telescopes from lenses, foam and cardboard. Not the same materials as Galileo but comparable in size and quality to the telescopes that Galileo used in 1610 to observe the phases of Venus and the Galilean moons.

EQUIPMENT: Telescope kit, resolution targets, mounts.

This lab is intended to give you some historical perspective on early astronomy – the telescopes that we will be assembling and using are very similar to the ones used by Galileo in his early explorations of the solar system and the night skies. Galileo, while he was not the first to invent the telescope (contrary to popular belief), used a similar refracting telescope to make the first observations of the largest Jovian satellites, the phases of Venus, and resolving clusters into individual stars. His telescope used plano-convex and plano-concave lenses (similar to those used here) and had an objective focal length of about 30-40 inches. Our telescopes will be smaller, but the quality of the lenses supplied really does more than make up for what Galileo worked with!

Part I. Build Your Telescope



The kits we have provided include all parts and can be assembled without the need for glue or adhesives of any kind. You should have a set of two sliding cardboard tubes (one within the other), and a small plastic bag containing:

1 red plastic cap
1 cardboard washer
1 large lens (objective)
1 small lens (eyepiece)
1 foam eyepiece holder
1 cardboard space (small tube)

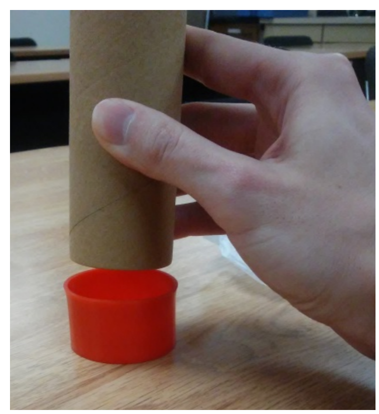
Note before starting – *DO NOT touch either lens with your bare hands, as fingerprints are extremely difficult to get off the lenses! Use a piece of cloth or a tissue at all times.*

Each telescope comprises of these pieces:

- Two red caps (one slightly larger than the other)
- One large lens (primary)
- One small lens (eyepiece)
- One cardboard washer
- One cardboard spacer
- One black foam cylinder
- Two sliding cardboard tubes

1. Place the larger red cap on the table with the lip facing downwards (open end up).

2. Pick up the large lens **BY THE EDGES** and look at its curvature. You should notice that one side is flat, and the other is curved outwards. This is called a plano-convex lens which is used to reduce geometrical aberrations. Place the lens inside the red cap so that the curved side is facing the outside of the telescope.

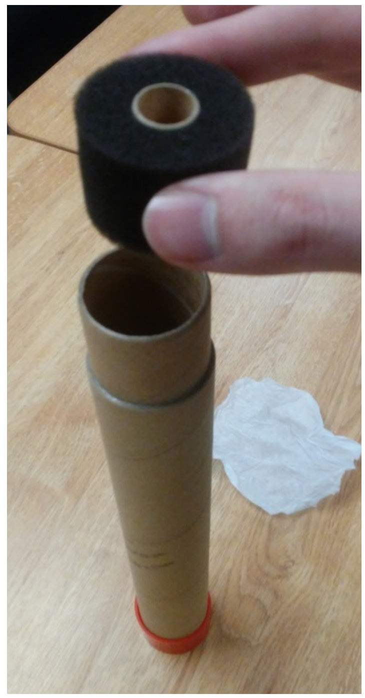


3. Insert the wider end of the two sliding cardboard tubes into the cap. This will hold the lens in place.

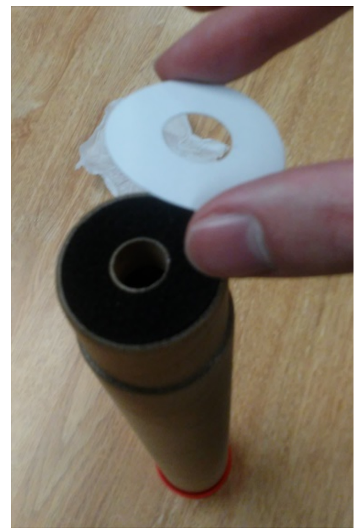
4. Place the black foam cylinder on the table. Put the cardboard spacer inside the hole in the cylinder so it is flush with the table.

5. Using a piece of cloth or tissue, push the small eyepiece lens into hole in the black foam cylinder until the lens rests on the spacer.

6. Push the black foam cylinder into the open end of the cardboard tube. The spacer should be positioned between the eyepiece lens and the outside of the telescope.



7. Place the cardboard washer on top of the eyepiece holder



8. Put the smaller red cap on the end of the telescope. You may wish to tape the caps of your telescope down to keep them from falling off.

There are 3 types of exercises described below that occur at 3 stations: (A) on the observing deck upstairs, (B) in the classroom, and (C) outside. They can be done in any order. Spread yourself out so your whole lab section is not trying to do the same things at the same time.

Part II. Exercise A: Observing from the Deck

II.1 The Flatirons. Take your DIY telescope up onto the deck and check out the flatirons. Note how to focus the telescope by sliding the inner tube relative to the outer tube. Is the image the same as your eyes or inverted? How small an object can you see on the Flatirons? The CU sign on the 3rd flatiron? Trees? Climbers?

CAUTION!
NEVER LOOK DIRECTLY AT THE SUN THROUGH THE TELESCOPE!!!

II.2 The Sun. Scientists for many years have been looking at the sun through a variety of scopes (although not with their own eyes!) in order to study solar behavior. While most images taken of the sun nowadays are multi-wavelength (the Sun emits detectable light from X-ray all the way through radio), visual images are especially useful for viewing sunspots and prominences like flares and coronal mass ejections.

The main features we expect to see on the Sun with our telescope are *sunspots*. Sunspots are caused by variations of the magnetic field lines within the Sun.

Solar observing is normally done with special solar telescopes like our heliostat here at SBO, where the image is projected onto a screen (rather than looking directly at the source). Simple refractors like the one we've just assembled, though, can also be used to observe the sun – we can spare our own eyes from looking at the blinding solar radiation using a method called *eyepiece projection*.

- Focus your telescope on a distant object (such as the crest of the mountains or a far-off building). Make a pencil mark on the smaller tube of your telescope so that you have a reference the next time you want to focus.

II.2.1. Place your sheet of white paper on the ground (or tape it to a wall if the Sun is approaching the horizon). Align your focused telescope with the sun until you see a projected image of the sun on your paper – try to get a nice, crisp edge. Move the telescope back and forth along your line of sight to the Sun. Does the sharpness/focus of the image change? What about the image size?

II.2.2. Once you have the Sun located and focused, have one member of your group trace the Sun on your paper, preserving its size and shape. Be sure to include any individual features that you notice, such as sunspots or flares. Describe what you see.

II.2.3. Estimate the size of the sunspots you sketched using your projection in terms of solar diameters. The diameter of the Sun is approximately 1.4 x 10^6 km. What is the diameter of the sunspots that you measured in Earth diameters?

II.2.4. Can you measure the motion of the Sun? You will need to have your telescope fixed to a stand. How long does the Sun take to move across the field of view of the telescope? Compare with (i) what you know about the rate of motion of the Sun across the sky, or (ii) your measurement in part IV of the angular field of view of the telescope.

Part III. Exercise B: Resolution and Aberration – in the Classroom

The **magnification** of a telescope matters a lot. But it is no good just making objects look bigger – they need to be clear, crisp – well resolved. So, let's measure the **resolution** of some telescopes. The way we measure the resolution of our telescopes is to use targets at a known distance to view the separation between lines.

- Place the telescope on the equatorial mounts provided and attach it firmly with the rubber bands. Point it at the resolution target at the other end of the room (or room along the hall or across the deck – depends where you are located and who's using the other spaces). Focus the telescope.

Focus on the smallest grouping of lines in which you can distinguish individual lines (the white space in between the lines). What is the smallest spacing that you can distinguish with your scope? Use the conversion table on the back of the chart to write down what is the smallest spacing (in mm/line) you can resolve with the DIY telescope.

What is the distance to the target? Measure in meters.

The angular resolution of your telescope (usually measured in arc seconds – 1/3600 of a degree) can be found using the spacing between lines and the distance to the target:

$$\theta = (180 \text{ degrees} / \pi \text{ radians}) * (3600 \text{ arcsec/degree}) * \frac{(\text{line spacing})}{(\text{target distance})}$$

Make sure you convert your line-spacing and target distance to the same units (e.g. meters). What is the angular resolution of your telescope?

Now, let's compare this to the **diffraction limit** – the theoretical maximum of your telescope's resolving power. The diffraction limit of a telescope (in visual wavelengths) is given as $\theta = (120/D)$ arcseconds, where D is the diameter of the aperture measured in millimeters. What is the diffraction limit of your telescope?

Now compare the resolution of your DIY telescope with the resolution and other commercial telescopes in the lab. How does your telescope compare? Of course, you really ought to compare price as well as resolution!

The telescopes we've assembled use plano-convex lenses, which reduce the geometrical aberration in the image. Point the scope again at the resolution target and look carefully toward the edges. Describe the – **aberration** - geometrical distortion (if any) that you see. Can you see any distortion of color – **chromatic aberration**?

Part IV. Exercise C: Solar System – the Colorado Scale Model Solar System

Now, let's go back to the model Solar System outside Fiske Planetarium that we used in the first lab. Typically, some of the brightest objects we see in the sky are right in our own backyard, astronomically speaking - not surprising, considering the range in distance between our Solar System and even the nearest star. Let's simulate a couple of Galileo's classic measurements: phases of Venus and moons of Jupiter.

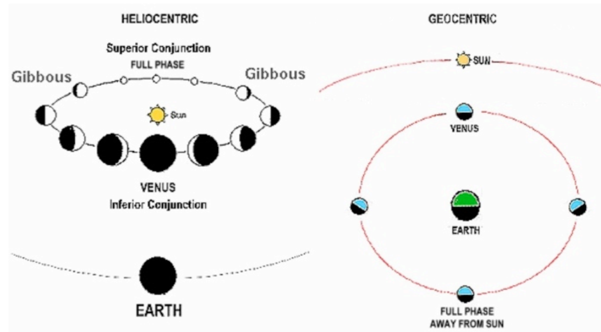
Split up your lab group in two – have one group stay fixed at the Earth with the scope and the other group roam around the Solar System – actually just Venus and Jupiter. We can use the scaling of the Solar System here to predict whether we could observe these objects on a clear night. Remember from the first lab – the model Solar System is scaled so that both distances and sizes are 10^{-10} of the real Solar System.

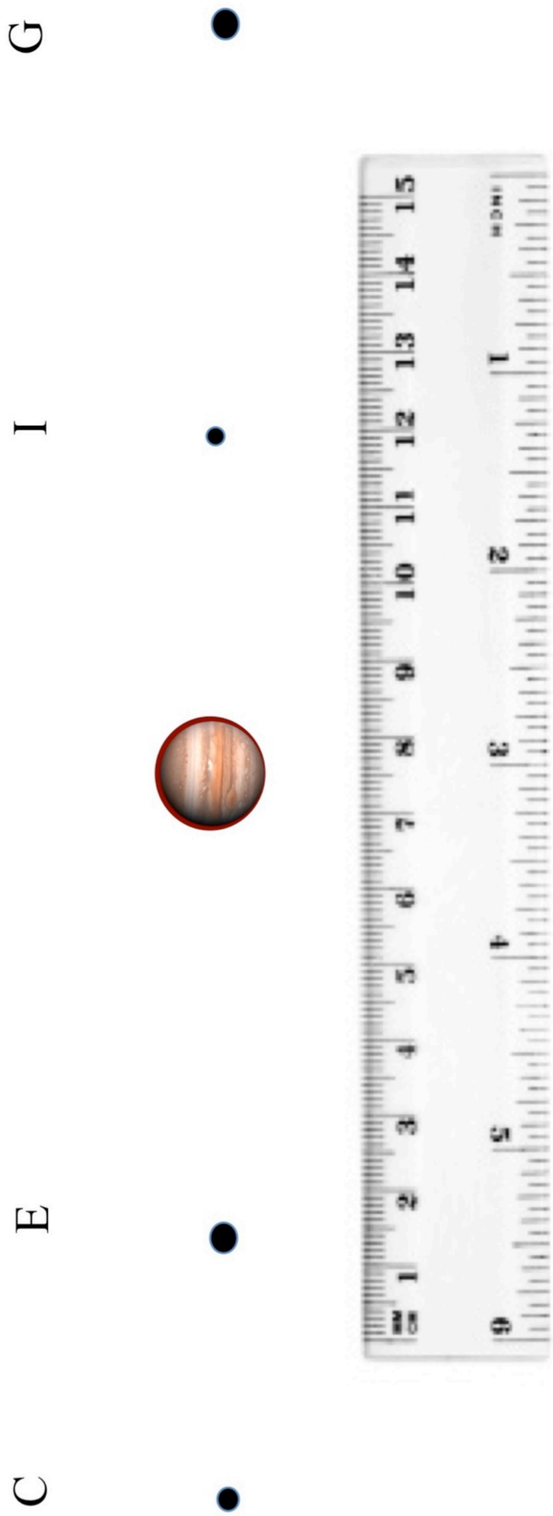
On the back page of the lab exercise are drawn (i) Venus at various phases and (ii) the moons of Jupiter – to the 10-billionth scale. Take turns to hold up the drawing so the people a Venus or Jupiter distance away can look through their telescope at the drawing. It does not matter whether you look from Earth or Jupiter (for example) – the distance is the same between the two from either direction.

VI.1. Have the roaming group go to Venus. Check that the printed drawing of Venus is about the same size as the scaled planet Venus on the plaque. Standing next to the plaque, hold up your images so they can be seen by the group standing at Earth. And vice versa. Can you make out the shapes with your naked eye? With the telescope? You may want to mount the telescope on a plaque to keep it steady.

VI.2. Moving on the other side, head out to Jupiter. Check that the drawing is about the same size as the planet on the plaque. Take turns holding up the drawing of Jupiter and the moons – and looking through the telescope. Can you make out the individual moons? How big does Jupiter appear in your field of view?

VI.3. Back at Earth, take one last look at the Sun. Remember (you can use the "hand calculator" as a check) that the Sun has angular size of about 1/2 degree. Look at the scale Sun through your telescope – based on that, what is the estimated size of your telescope's field of view? (the diameter, in degrees, that your telescope can see in a single viewing).


HELIOCENTRIC
Superior Conjunction
FULL PHASE
Gibbous
Gibbous
Sun
VENUS
Inferior Conjunction
EARTH
GEOCENTRIC
SUN
VENUS
EARTH
FULL PHASE
AWAY FROM SUN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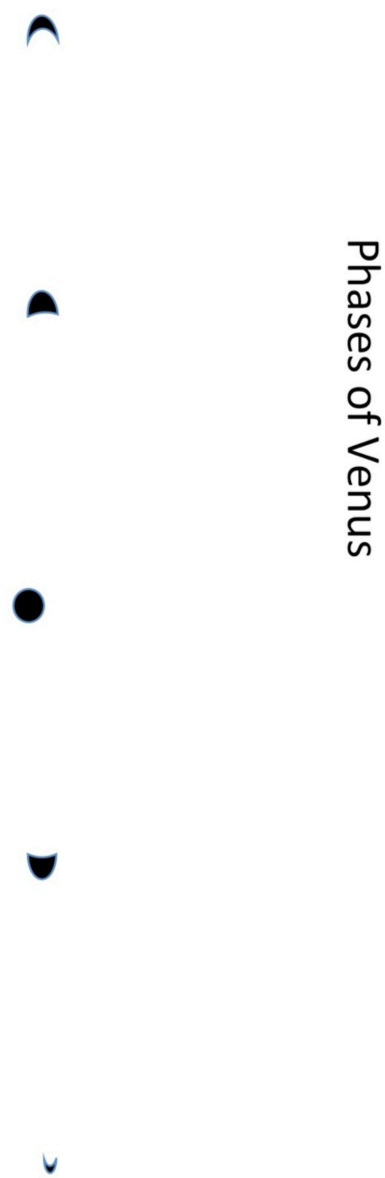

Phases of Venus

Name: ______________________________

D.I.Y. TELESCOPE | Post-Lab = please hand in at the start of next lab

1. **Metacognition:** We are always trying to improve the ASTR1030 labs and value your input on where they are succeeding and where they could be improved. Please briefly describe, what did you like and/or not like about this lab? What main idea(s) do you feel you learned from this lab?

Name: ____________________________

TELESCOPE OPTICS

Pre-Lab = please complete *before* coming to lab

Read briefly through the entire lab. Check this box ☐ to indicate that you have done so. Jot down any questions that occurred to you while skimming through the lab description.

1. In your own words, explain what is meant by the terms **object**, **image**, **focal plane,** and **magnification** as they are used in this lab.

   Object

   Image

   Focal plane

   Magnification

Name: ____________________________

TELESCOPE OPTICS

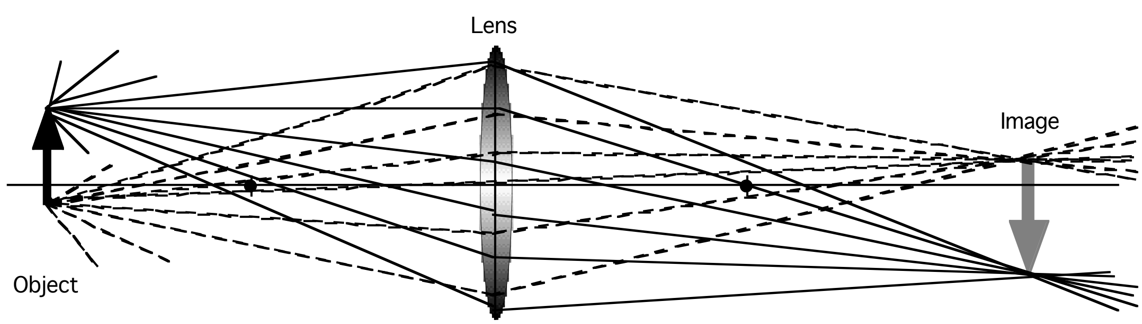
SYNOPSIS: You will explore some image-formation properties of a lens, and then assemble and observe through several different types of telescope designs.

EQUIPMENT: Optics bench rail with 3 holders; optics equipment stand (flashlight, mount O, lenses L1 and L2, image screen I, eyepieces E1 and E2, mirror M, diagonal X); object box.

NOTE: Optical components are delicate and are easily scratched or damaged. Please handle the components carefully and avoid touching any optical surfaces.

Part I. The Camera

In optical terminology, an **object** is any source of light. The object may be self-luminous (such as a lamp or a star) or may simply be a source of reflected light (such as a tree or a planet). If light from an object happens to pass through a **lens**, those rays will be bent (refracted) and will come to a **focus** to form an **image** of the original object.



From *each* point on the object, light rays are emitted in all directions. Any rays that encounter the lens are bent into a new direction, but in such a manner that they all converge through one single point on the opposite side of the lens. Thus, one point on the actual object will focus into one corresponding point on what's called the **focal plane** of the lens. The same thing is true for light rays emanating *from each other point* on the object, although these rays enter the lens at a different angle, and so are bent in a different direction, and again pass through a (different) unique point in the focal plane. The image is composed of an infinite number of points where all of the rays from the different parts of the object converge.

To see what this looks like in "real-life," arrange the optical bench as follows:

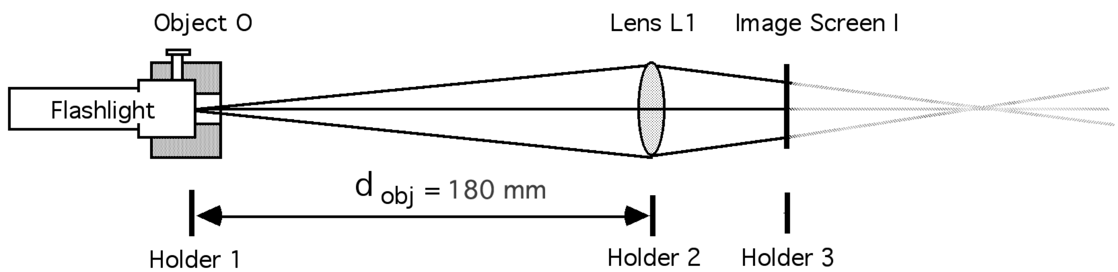
First, loosen the clamping knob of holder #1 and slide it all the way to the left end of the optics rail until it encounters the stop (which prevents the holder from sliding off of the rail). Clamp it in place.

Next, turn on the flashlight (stored in the accessory rack) by rotating its handle, and take a look at its illuminated face. The pattern you see will serve as the physical **object** in our study.

Insert the large end of the flashlight into the large opening of the Object Mount O and clamp it in place with the gold knob. Install the mount and flashlight into the *tall* rod holder of holder #1 so that the flashlight points and down the rail. (Allow the rod to drop fully into the holder so that the mounting collar determines the height of the flashlight; rod collars are used to ensure that all components are positioned at the same height. Please do not adjust the rod collar unless instructed to by your TA.)

The white mark on the backside of each holder indicates the location of the optical component in that holder; thus, measuring the separation between marks is equivalent to measuring the separation of the optical components themselves. Use the meter stick to measure the separation of the white marks so as to position holder #2 at a distance 180 mm (18 cm) from holder #1; clamp it in place. Install lens L1 in the holder so that one of its glass surfaces faces the flashlight.

Finally, put the image screen I (with white card facing the lens) in holder #3 on the opposite side of the lens from the flashlight. Your arrangement should look like this:



The separation between the lens and the object is called the **object distance** d_{object}; because you've positioned the object (flashlight) 180 mm from the center of the lens, the object distance in this case is 180 mm.

On the white screen, you will see a bright circular blob, which is the defocused light from the object that is being bent through the lens. Slowly slide the screen holder #3 back and forth along the rail while observing the pattern of light formed on the screen. At one unique point, the beam of light will coalesce from a fuzzy blob into a sharp image of the object. Clamp the screen at this location where the image is in best **focus**.

I.1 *Predict* what will happen to the image if you swap the positions of the flashlight (object) and the image screen. Explain your reasoning. *(You will not be marked down if your prediction is wrong, so please make an honest prediction before continuing.)*

I.2 Swap the object and image screen (*Hint:* Leave the holders in place so you can return to this arrangement, just take the posts out of the holders.) Was your prediction correct? If not, explain what you see.

Return to your original setup before continuing.

The term **magnification** refers to how many times larger the focused image appears, compared to the actual size of the object:

$$\text{Magnification (definition)} = \frac{\text{Image Size}}{\text{Object Size}} \qquad (1)$$

I.3 What should the units of magnification be?

I.4 What is the magnification produced by this optical arrangement?

Observed magnification = ______________

Explain how you calculated this magnification.

The distance between the lens to the in-focus image is called the **image distance**, d_{image}. In optical terminology, distances are always given in terms of how far things are from the main optical component (in this case, the lens). Instead of directly measuring the magnification, you can also calculate it from the ratio of image distance to the object distance:

$$\text{Magnification (calculated)} = d_{\text{image}} / d_{\text{object}} \qquad (2)$$

I.5 Use the meter stick and the two white marks on holders #2 and #3 to determine the image distance from the lens; record your result to the nearest millimeter:

d_{image} = _________

I.6 Show that Equation (2) gives you (at least approximately) the same value for the magnification that you determined from the image and object sizes:

Calculated magnification = ___________

Now let's find out how things change if the object is a little further from the lens. Unclamp and move holder #2 so that the distance between the object and the lens is somewhat larger than before (say, 200 mm or so). Now move the image screen I to find the new image location.

I.7 (a) When you increased the distance to the object from the lens, did the image distance get closer or farther away from the lens?

(b) Did the magnification increase or decrease?

(c) Move the object a small amount once again and verify that you can still find an in-focus image location on the opposite side. How does the direction of movement of the object relate to the direction of movement of the in-focus image?

Hopefully, you've found that a lens can be used to produce a magnified image of an object, and that the magnification can be varied. But it is also possible to make a *de-magnified* image instead (that is, *smaller* than the original object).

I.8 Move lens L1 (by sliding holder #2) to a position so that the image size is *less* than the original object size.

What is the new object distance? ________

What is the new image distance? ________

What magnification (using Equation 2) does this imply? _________

What is the measured image size? __________

Does the magnification using Equation 1 agree with the magnification you calculated using Equation 2?

Now let's see what will happen if we use a *different* lens. Replace lens L1 with the lens marked L2, but otherwise leave the positions of the holders in exactly the same place.

I.9 Refocus the image.

What is the new image distance with this lens? d_{image} = ______

What is the magnification produced by this arrangement? ________

Indicate whether you determined the magnification by definition (Equation 1) or calculation (Equation 2):

By now you have seen that, for any given lens, and distance of an object from it, there is one (and *only* one) location behind the lens where an image is formed. By changing either the lens or the distance to the object, or both, the location and magnification of the image can also be changed.

The optical arrangement you have been experimenting with is the same as that used in a **camera**, which consists of a lens with a piece of photographic film (or a digital chip) behind it, which records the pattern of light falling onto it. The film/chip is held at a fixed location, which is represented in our optical arrangement by screen I in holder #3.

I.10 Using your experience above (question I.7): if you move closer to an object that you're trying to photograph (smaller object distance), will the lens-to-film/chip distance (image distance) in your camera have to get larger or smaller to keep the image in focus?

I.11 How do you think cameras achieve the proper focus?

In many cameras it is possible to swap lenses, so that the same camera will yield much larger images of distant objects (i.e., a "telephoto" lens).

I.12 Which of the two lenses produced a larger (focused) image under the same conditions?

Which would more likely be considered to be a "telephoto" lens?

Part II. The Lens Equation

Because object and image distances from a lens seem to be related in a predictable manner, you probably won't be surprised to learn that there is a mathematical relationship between the two. It's called the **lens equation,** and for any given lens it looks like this:

$$\frac{1}{f} = \frac{1}{d_{object}} + \frac{1}{d_{image}} \qquad (3)$$

The value f in the formula is called the **focal length** of the particular lens being used. This is a *property of the lens itself* and doesn't change regardless of the location of the object or the image. Notice that the formula actually relates the reciprocal of the values of f, d_{object}, and d_{image}, instead of the actual values themselves.

II.1 Calculate the focal length *f* of lens L1, using your measured image distance d_{image} and object distance d_{object}. You can use either the arrangement from step I.5, or from step I.8, or from both (to see if they give the same answer).

Focal length *f* of Lens L1 = ________________

II.2 What is the focal length of lens L2 (using the information measured in step I.9)?

Focal length *f* of Lens L2 = ____________________

II.3 Which lens, L1 or L2, has the longer focal length? ____________

Considering what you found out in I.9, does a telephoto lens have a longer or shorter focal length than a "normal" lens?

II.4 Using the Lens Equation (3), why was the image you made in I.2 still in focus?

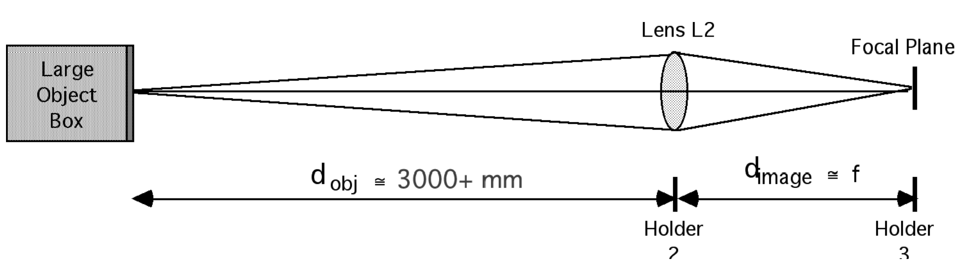
Part III – Observing Distant Objects

In astronomy, we look at objects that are *extremely* far away. For all practical purposes, we can say that the object distance is *infinitely* large. This means that the value $1/d_{object}$ in the lens equation is *extremely* small, and can be said to equal zero. Thus, for the special case where an object is very far away, the lens equation simplifies to:

$$\frac{1}{f} = 0 + \frac{1}{d_{image}} \qquad \text{or simply} \qquad d_{image} = f$$

In other words, when we look at distant objects, images are formed behind the lens at a distance equal to (or very nearly so) its focal length. Because *everything* we look at in astronomy is very far away, images through a telescope are *always* formed at the focal length of the lens.

We can now look at an object that is far enough away to treat it as being "at infinity." Remove the flashlight and mount from holder #1 and replace them in the optics storage rack. Instead, use the large object box *at the opposite end of the table* for your light source. It should be at least 3 meters (10 feet) from your optics bench. Focus the screen. Your arrangement should look like this:



III.1 Measure the image distance from the markings on the optics rail. Confirm that this is fairly close to the value of the focal length of lens L2 that you calculated in II.2.

Measured image distance = ______________________

Your measurement will be slightly larger than the true value of *f,* simply because the illuminated object box isn't *really* infinitely far away; however, since we can't take this telescope outside, we will treat your measurement in III.1 as the actual focal length of the lens.

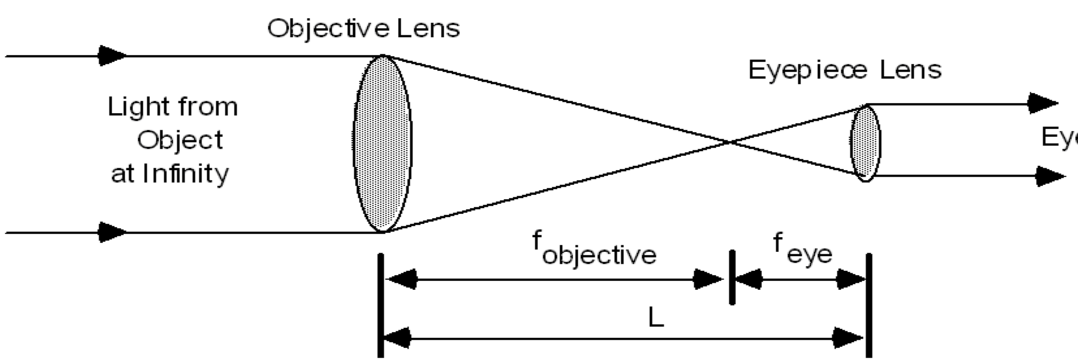
If the screen were not present, the light rays would continue to pass through and beyond the image that forms at the focal plane. In fact, the rays would diverge from the image as if it were a real object, suggesting that *the image formed by one lens can be used as the object for a second lens.*

III.2 Remove the white card from the screen to expose the central hole. Aim the optics bench so that the image passes into the opening.

III.3 From well behind the image screen, look along the optical axis at the opening in the screen. You should be able to see the image once again, "floating in space" in the middle of the opening! If you move your head slightly from side-to-side you should get the visual impression that the image doesn't shift around but instead seems to be fixed in space at the center of the hole. By sliding the screen back and forth along the rail, you can also observe that the opening passes around the image, while the image itself remains stationary *as if it were an actual object.*

Notice that the image is quite tiny compared to the size of the original object (a situation that is always true when looking at distant objects). In order to see the image more clearly, you will need a magnifying glass:

III.4 Remove screen I from its holder and replace it with the magnifying lens E1. Observe through the magnifier as you slowly slide it back away from the image. (*Hint:* Your eye should be right next to the magnifying lens.) At some point, a greatly enlarged image will come into focus. If you run off the end of the rail you will have to move the whole arrangement towards the light. Clamp the lens in place where the image appears sharpest. Your arrangement will be as follows:



You have assembled a **refracting telescope**, which uses two lenses to observe distant objects. The main telescope lens, called the **objective lens,** takes light from the object at infinity and produces an image exactly at its focal length $f_{objective}$ behind the lens. Properly focused, the magnifier lens (the **eyepiece**) does just the opposite: it takes the light from the image and makes it appear to come from infinity. The telescope itself *never forms a final image*; it requires another optical component (the lens in your eye) to bring the image to a focus on your retina.

To make the image appear to be located at infinity (and hence observable without eyestrain), the eyepiece must be positioned behind the image at a distance exactly equal to *its* focal length, f_{eye}. Therefore, the total separation, *L*, between the two lenses must equal the sum of their focal lengths:

$$L = f_{objective} + f_{eye} \tag{4}$$

III.5 Measure the separation between the two lenses (L) and use your value for the objective lens focal length ($f_{objective}$, measured in step III.1) and equation (4) to calculate the focal length of eyepiece E1 (f_{eye}).

L: ____________ f_{eye}:____________

Earlier, we used the term "magnification" to refer to the actual *physical size* of the image compared to the actual *physical size* of the object. This cannot be applied to telescopes, because no final image is formed (and besides, the physical sizes of objects studied in astronomy, like stars and planets, are *huge*). Instead, we use the concept of **angular magnification**: the ratio of the *angular* size of an image appearing in the eyepiece compared to the object's actual *angular* size. In other words, we're referring to how much bigger something *appears* to be, rather than to how big it actually is.

The angular magnification *M* produced by a telescope can be shown to be equal to the ratio of the focal length of the objective to the focal length of the eyepiece:

$$M = \frac{f_{objective}}{f_{eye}}. \tag{5}$$

III.6 Calculate the magnification of the telescope arrangement you're now using (objective lens L2 and eyepiece E1).

Equation 5 implies that if you used a telescope with a *longer* focal length objective lens (make $f_{objective}$ bigger), the magnification would be greater. But the equation also implies that you can increase the magnification of your telescope simply by using an eyepiece with a *shorter* focal length (make f_{eye} smaller).

III.7 Eyepiece E2 has a focal length of 18 mm. Without inserting E2 into the actual system yet, use equation 5 to calculate the magnification resulting from using eyepiece E2 with objective lens L2: ____________

III.8 Replace eyepiece E1 with E2, and refocus. Did the image get bigger or smaller than before?

____

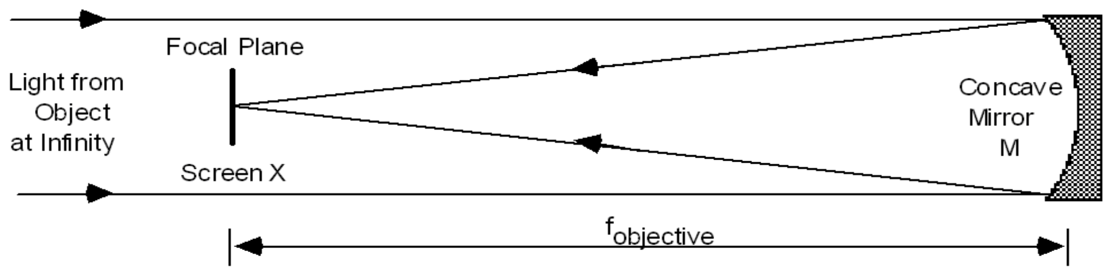
Is your observation consistent with the calculation of III.7?

Consumer Tip: Some inexpensive telescopes are advertised as "high power" (= "large magnification") because most consumers think that this means a "good" telescope. You now know that a telescope can exhibit a large magnification simply by switching to an eyepiece with a very short focal length. Magnification really has nothing to do with the actual quality of the telescope!

Part IV. The Reflecting Telescope

Concave mirrored surfaces can be used in place of lenses to form **reflecting telescopes** (or **reflectors**), rather than refracting telescopes. All of the image-forming properties of lenses also apply to reflectors, except that the image is formed *in front* of a mirror rather than behind. As we will see, this poses some problems! Reflecting telescopes can be organized in a variety of configurations, three of which you will assemble below.

The **prime focus** arrangement is the simplest form of reflector, consisting of the image-forming objective mirror and a flat surface located at the focal plane. A variation on this arrangement is used in a **Schmidt camera** to achieve wide-field photography of the sky.



IV.1 Assemble the prime-focus reflector shown above:

- ° First, return all components from the optical bench to their appropriate locations in the optics storage rack.
- ° Next, slide holder #3 all the way to the end stop away from the light source, and slide holder #2 as far away from the light source as possible until it is touching holder #3.
- ° Install the large mirror M into holder #2, and carefully aim it so that the light from the distant object box is reflected straight back down the bench rail.
- ° Place the mirror/screen X into the *tall* rod holder #1, with the white screen facing the mirror M.
- ° Finally, move the screen back and forth along the rail until you find the location where the image of the object box is focused onto the screen. (Tip: you can use the white card from the image screen to find the image that is being reflected from the mirror, which will let you know if you need to rotate the mirror so that the beam is directed at screen X).

IV.2 Is the image right-side-up, or inverted? ___________________

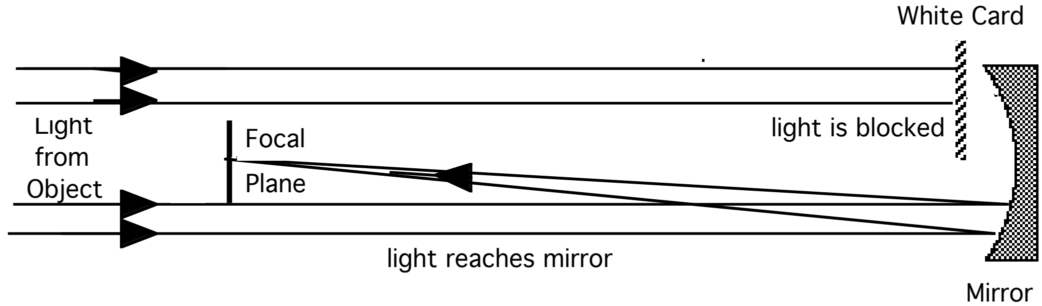
IV.3 Measure the focal length *f* of the mirror M just like you did with the refractor: use your meter stick to measure the distance from the mirror to the image location.

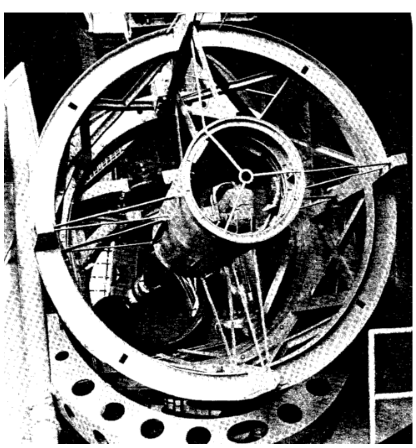
Focal length of the mirror = ____________

Because the screen X obstructs light from the object and prevents it from illuminating the center of the mirror, many people are surprised that the image does not have a "hole" in its middle.

IV.4 Hold the white card partially in front of the objective mirror in order to block a portion of the beam, as shown below. Note that no matter what portion of the mirror you obscure, the image of the distant object box stays fixed in size and location on the screen. This is because *each small portion of the mirror forms a complete and identical image* of the object at the focus! However, as you block more and more of the mirror, what happens to the image?

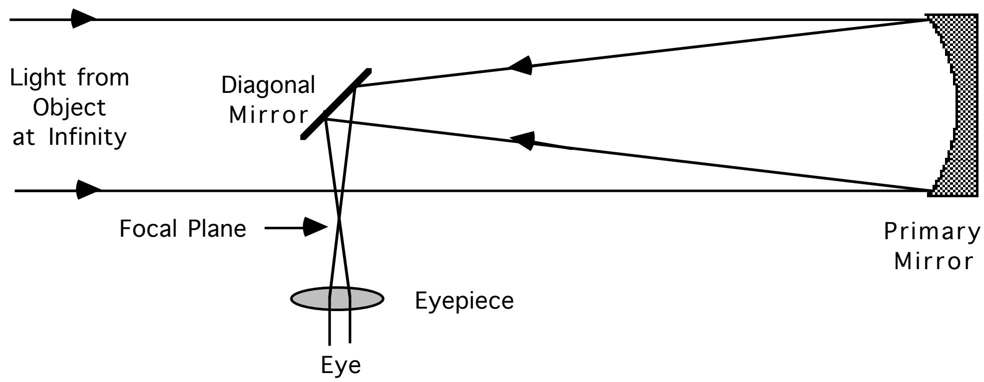
Explain your answer:





The prime focus arrangement cannot be used for eyepiece viewing, because the image falls inside the telescope tube. If you tried to see the image using an eyepiece, your head would also block all of the incoming light. (This, however, is not the case for an extremely large mirror: for example, the 200-inch diameter telescope at Mount Palomar has a small cage in which the observer can actually sit inside of the telescope at prime focus, as shown here!)

Isaac Newton solved the head-obstruction problem for small telescopes with his **Newtonian reflector**, which uses a flat mirror oriented diagonally to redirect the light to the side of the telescope, as shown below. The image-forming mirror is called the **primary** mirror, while the small additional mirror is called the **secondary** or **diagonal** mirror.



IV.5 Convert your telescope to a Newtonian arrangement:

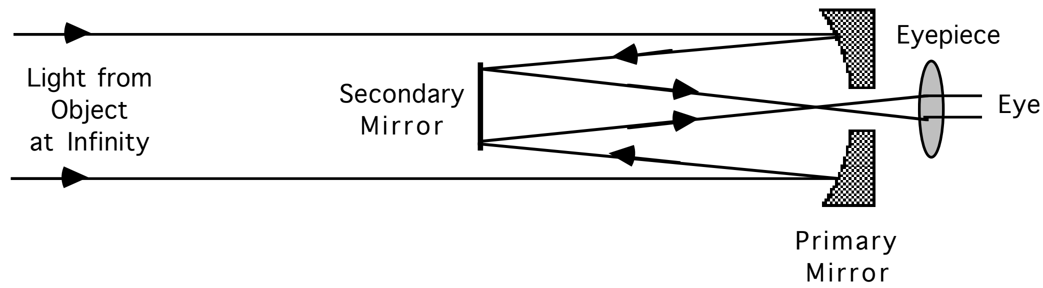
- Reverse the mirror/screen X so that the *mirror* side faces the primary mirror M and is oriented at a 45° angle to it.
- Install eyepiece E1 in the *short* side rod holder of holder #1, and look through the eyepiece at the diagonal mirror (see diagram above).
- Now rotate the diagonal mirror slightly until you see a flash of light that is the bright but out-of-focus image of the distant object. Clamp the diagonal mirror in place.

  Now, while looking through the eyepiece, slowly slide holder #1 towards the mirror until the image comes into sharp focus.

Congratulations! You've constructed a classical Newtonian telescope, one of the most popular forms of telescopes used by amateur astronomers!

IV.6 Use equation 5 and your knowledge of the mirror M's focal length and eyepiece E1's focal length (steps IV.3 and III.5, respectively) to determine this telescope's magnification:

Large reflecting telescopes (including the 20-inch and 24-inch diameter telescopes at Sommers-Bausch Observatory) are usually of the **Cassegrain** design, in which the small secondary redirects the light back towards the primary. A central hole in the primary mirror permits the light to pass to the rear of the telescope, where the image is viewed with an eyepiece, camera, or other instrumentation.



IV.7 Re-arrange the telescope into a Cassegrain configuration as shown above.

° Remove the eyepiece E1 from its short holder in holder #1, and transfer it instead to holder #3 behind the mirror.

° Carefully re-orient the secondary mirror (still in holder #1) so that it reflects light directly back towards the hole in the primary mirror.

° Now slide holder #1 towards the mirror until it is only about 35% of its original separation from the mirror (about 1/3rd of the mirror focal length found in IV.3).

° While looking through the eyepiece, rotate the secondary mirror in holder #1 slightly until you see the flash of light that is the image of the object. Then move holder #1 slightly towards or away from you until you can see an in- focus image.

IV.8 Do you think that the magnification of this image is any different from the magnification you observed with the Newtonian arrangement? Why or why not?

Note: in a *real* Cassegrain telescope, the magnification *would* in fact be different because the secondary mirror is not actually flat, but instead is a convex shape. This permits the size of the secondary mirror to be smaller than the flat secondary you are using here, and thus allows more light to strike the primary mirror.

IV.9 How does the overall length of the Cassegrain telescope design compare with that of the Newtonian style?

Can you come up with one or more reasons why the Cassegrain design might have advantages over the Newtonian? Explain your ideas:

IV.10 Why does the hole in the primary mirror not cause any *additional* loss in the light-collecting ability of the telescope?

IV.11 Compare the *actual* optical and astronomical equipment that has been provided in the lab room (35-mm film **camera**, small **refractor**, **Schmidt** telescope camera, **Newtonian** reflector, **Cassegrain** reflector) with the different telescope styles that you have just assembled. Note especially that, except for having enclosed tubes and more sophisticated controls, each design you assembled is essentially identical to "the real thing"!

Name: ______________________________

TELESCOPE OPTICS | Post-Lab = please hand in at the start of next lab

1. You now know the importance of focal length when choosing a telescope. The other factor to consider is the telescope's light-gathering power which is proportional to the square of the diameter of the lens or mirror. How much more light-gathering power does one of the SBO 20-inch telescopes have compared to the human eye (which has a diameter of about 5 mm)?

2. **Metacognition:** We are always trying to improve the ASTR1030 labs and value your input on where they are succeeding and where they could be improved. Please briefly describe, what did you like and/or not like about this lab? What main idea(s) do you feel you learned from this lab?

MASS OF SATURN

Name: ___________________________

Pre-Lab = please complete *before* coming to lab

Read briefly through the entire lab. Check this box ☐ to indicate that you have done so. Jot down any questions that occurred to you while skimming through the lab description.

MASS OF SATURN

Newton's Version of Kepler's 3rd Law (NVK3L) has three variables: the semi-major axis of the orbit (**a**), orbital period (**p**), and the mass of whatever is being orbited around (**M**). If you know any two of these quantities, you can find the third.

$$p^2 = \frac{4\pi^2}{GM}a^3 \qquad\qquad M = \frac{4\pi^2}{G}\frac{a^3}{p^2}$$

Most usefully, this is how we can use the motion of an object due to gravity to determine mass.

This is the third part of a three-part observational lab. The first two parts composed the night lab **Saturn Observations**. In that lab, students (including you?) observed Saturn with the telescopes at SBO both visually (through the eyepiece) and using the camera to record a detailed digital image. Next, they measured four quantities in the image and uploaded them to the class data archive.

Part 1 - Expectations

Kepler's first law tells us that orbits are elliptical. The moons of planets are no exception to this rule. The rings of Saturn are composed of trillions of chunks of ice so they too can be thought of as tiny moons, all orbiting Saturn on elliptical orbits. In the case of Titan and the ring particles, the orbits are almost perfect circles. Furthermore, the rings and most of the large moons (including Titan) orbit in the equatorial plane of Saturn.

Q1.1: In the space below, draw a sketch of what Saturn (the planet), the rings, and the orbit of Titan would look like **as viewed from directly above Saturn's north pole**. Titan orbits about ten times farther from the planet than the rings do. Recall the observed pattern in our solar system about what direction objects tend to orbit and rotate. Indicate where the north pole is and which direction Titan and the rings are rotating.

As viewed from above the north pole, the Saturnian system should look something like what you drew. However, we are viewing Saturn from within the Ecliptic plane not above the north pole. Saturn, like Earth, has a rotation axis which is tilted with respect to the Ecliptic. Saturn's **axial tilt is 27°** and, in 2019, is experiencing summer in the northern hemisphere (Saturn's north pole is tilted toward the Sun (and the Earth) by 27 degrees). Since a Saturnian year is almost 30 Earth years, northern summer will last for a long time!

Q1.2: Sketch what Saturn, the rings, and Titan's orbit should look like **as viewed from Earth**. Indicate where the north pole is and which direction Titan and the rings are rotating. [hint, think of what Saturn looks like through a telescope on Earth, as shown at the end of this lab.]

Q1.3: In part 2 of the night lab, students measured four quantities: radius of Saturn, radius of the rings, distance to Titan, and Titan's position angle (the compass direction from Saturn to Titan). Which quantities do you expect to vary as Titan orbits Saturn?

Q1.4: Look at the sketch you made for Saturn as seen from Earth. How will the apparent **distance** between Titan and Saturn change with **time**? Sketch a plot of what you expect.

Part 2 - Observational Data

We have taken your images and measurements from the Night Lab and compiled them into a table of the Saturn-Titan data. Measured distances are in **Saturn radii**. Please refer to the data table to answer the questions in this section.

Q2.1: The radius of Saturn and its rings shouldn't change (much) over the course of our observations. How do you explain the variations in measured quantities between groups?

Since observations were taken over a few months, different groups observed Titan at different points in its orbit. By plotting out these points, we can determine the semi-major axis of Titan's orbit. By looking at how long it takes to complete one orbit, we can find Titan's orbital period. Since we are interested in plotting how Titan's distance and position angle vary with time, we have converted observation date into "Day of Semester". This will make plotting quantities versus time a great deal easier than calendar date.

Q2.2: Use the graph paper below to plot the Saturn-Titan distance on each observation date. Time should go on the horizontal axis (x-axis); 2 days per box is a convenient unit. The vertical axis (y-axis) will be D_{Titan}. Pick a value for the tick marks that make the plot use a good range of the y-axis. Be sure to label your axes.

If we had observations every night, it would be easier to see the orbit in distance versus time plot you created above and your plot might look like your expected behavior. Perhaps you can see a pattern in the data, but maybe not. The data are **sparsely-sampled** (there are long gaps between observations). However, Titan's position is **periodic**; it should appear at the same location every orbit. This allows us to analyze the data in a different way.

Q2.3: Plot **distance** versus **position angle** on Figure 2. To visualize Titan's orbit, it is much more convenient to use a polar plot (radius, angle) on the next page rather than a traditional cartesian (x,y) plot. On a polar plot, data is plotted as a distance from the center (radius) and a specific angle from 0 (angle). In order to fit the data onto one graph, let each circle in radius represent a unit of 2 Saturn radii.

1. Mark the cardinal directions outside of the circle. North should be at 0°, East at 90°, etc.

2. Sketch the "ball" of Saturn by drawing a circle with the appropriate diameter (remember that each ring on the graph represents an increase of 2 Saturn radii.)

3. Sketch the **rings** of Saturn as you did the planet. Draw an ellipse with the appropriate diameter in the east-west direction. The north-south dimension should be the same as the ball.

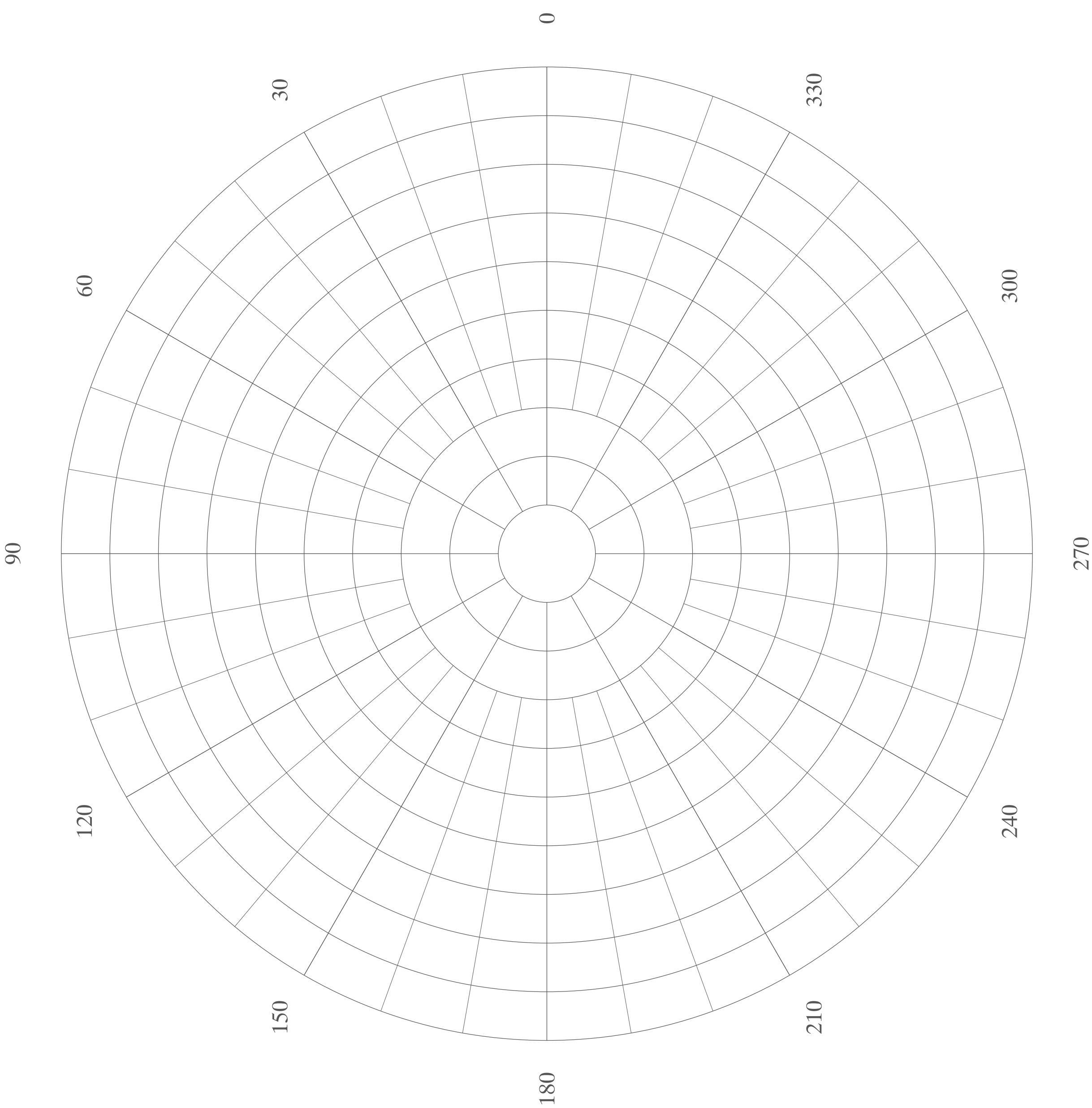
4. Now start plotting the Titan distances from the data table. Plot <u>all</u> the data points. Find the appropriate angle and mark at the appropriate distance from the center. Label each point carefully with the group ID or date so you can relate your sketch to the data in the data table.

5. Compare your sketched points to your expectations from Q1.2. Can you draw an elliptical orbit through the data points which looks like your expectation?

Q2.4: In plotting the data, did any of the data points seem suspicious? Two teams' data have been manipulated to include mistakes in recording their data. Which observations are mistaken? What do you think the mistakes might have been?

Q2.5: What is the **semi-major axis** of Titan's orbit (and what is the unit)?

Q2.6: What is the **period** of Titan's orbit?


0
30
330
60
300
90
270
120
240
150
210
180

Part 3 - Determining the Mass of Saturn

Use the plots you made in part 2 to determine the two quantities you need to solve for Saturn's mass.

Q3.1: To use NVK3L, we will need physical units (km) not Saturn radii. Convert your semi-major axis value from Q2.5. Use the fact that **Saturn's average radius is 58,232 km**.

a = ______________________ Saturn radii = ____________________ km

Q3.2: What is the period of Titan's orbit? **p =** ______________________ days

Q3.3: Now solve for the Mass of Saturn using NVK3L. In the units of km and days, G=4.98 x 10^{-10} $km^3/(kg \cdot day^2)$.

M = ___________________________ **kg**

Q3.4: How many times Earth's mass is Saturn? ($M_{Earth} = 6.0 \times 10^{24}$ kg).

Q3.5: Calculate Saturn's density. Use the radius from Q3.1. How does it compare to the density of water (1 g/cm^3)?

Name: ______________________________

MASS OF SATURN | Post-Lab = please hand in at the start of next lab

1. **Metacognition:** We are always trying to improve the ASTR1030 labs and value your input on where they are succeeding and where they could be improved. Please briefly describe, what did you like and/or not like about this lab? What main idea(s) do you feel you learned from this lab?

Names: ____________________________

____

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(you must have 2-3 people per lab group. Identify your partner(s) but still turn in your own work)

SURFACE AND ATMOSPHERE OF MARS

Pre-Lab = please complete *before* coming to lab

Read briefly through the entire lab. Check this box ☐ to indicate that you have done so. Jot down any questions that occurred to you while skimming through the lab description.

SURFACE AND ATMOSPHERE OF MARS

SYNOPSIS: The purpose of this lab is to allow you explore the surface of a planet in our solar system, and relate real spacecraft observations of the surface both to the basic geologic processes that we discuss in lecture and to computer simulations of the Martian atmosphere. NASA's *Mariner 4* was the first spacecraft to fly by Mars, in 1965, and the Soviet *Mars 2* spacecraft was the first Mars orbiter. At least 25 missions have successfully visited Mars over the past 5+ decades, collecting a wealth of observations about how the planet works. These observations, coupled with observations made from telescopes here at Earth, help us to understand not only the geology of the red planet, but also how its atmosphere behaves. The data are also used when developing models of for how the atmosphere is structured, moves, and evolves.

EQUIPMENT: Computer with internet connection, web browser, and JMARS software installed.

Getting Started

Here's how to get your computer up and running:

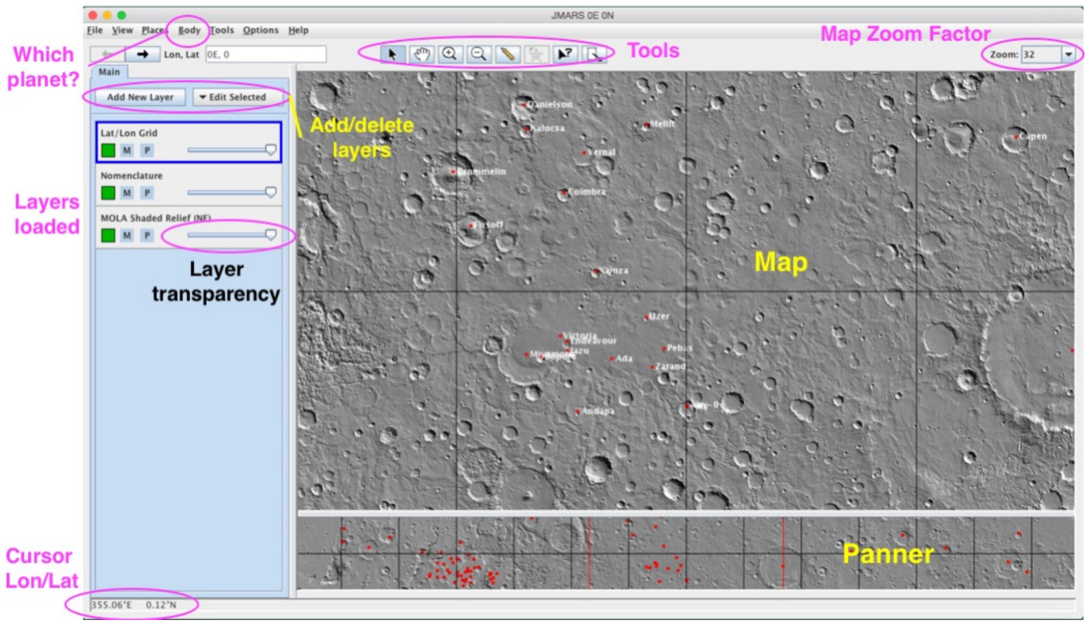
(1) To complete most sections of the lab you will need to use the JMARS software. JMARS should be installed on computers in the computer lab and accessible by clicking on an icon on the desktop. If you are working on your own computer then you can also download the JMARS software from Arizona State University (https://jmars.asu.edu/).

(2) Part V requires that you use results from a computer model of the Martian atmosphere. The model, made in France, is called the LMD GCM (Laboratoire de Météorologie Dynamique General Circulation Model). To access these results, click on the icon for the Mars Climate Database on the desktop of one of the computers in the computer lab. If working from home, or no icon is available, then launch an internet browser and go to http://www-mars.lmd.jussieu.fr/mcd_python/

Part I. Fiddling around with JMARS

JMARS is a software package used by the scientific community to explore data returned from planetary spacecraft missions. This software is used to conduct "real" science that results in published articles in scientific journals. But it is free for anyone to download and explore.

I.1 Get JMARS running. Once you've started JMARS you should see a screen like this:



I.2 Spend a few minutes learning to get around in JMARS by trying the following features. Initial each item when complete:

Use the pan tool to center a feature of interest ____________

Zoom in and out to see all of Mars, or specific features ____________

Turn the names of features on and off ____________

Figure out where your cursor is located in longitude and latitude ____________

Reorder the layers (by dragging them), and play with the transparency ____________

Use the measure tool (and look at the bottom left) to measure distances ____________

Add a layer of your choosing. Delete a layer. ____________

Reset the layers (under the "Body" pulldown menu) when done ____________

Part II. Surface Geology of Mars

Zoom out so that you can view Mars more or less globally. Make sure the MOLA Shaded Relief (NE) layer is loaded and visible. Turn the Nomenclature layer off to get a better view of the planet, turning it on only when you want to see the names of craters.

II.1 Does Mars look like a planet that has been geologically active? Support your answer.

II.2 Does Mars look like it has been more geologically active than Earth over time? Support your answer.

II.3 Find an example of a *volcanic* feature on Mars. Report the longitude and latitude of the feature. Finally, report the approximate area of the feature, in km^2. Show your work when computing the area. **Zoom in as necessary**.

Lon/Lat: ____________________ Area: ________________________

II.4 *Tectonics* involves deformation (bending, folding, stretching, breaking) of a planet's crust. Identify an example of a tectonic feature on Mars. Show your work when computing the feature's area in km^2. **Zoom in as necessary**.

Lon/Lat: ___________________ Area: ______________________

II.5 Identify an example of an *impact* feature on Mars. Show your work when computing the feature's area in km^2. **Zoom in as necessary**.

Lon/Lat: ___________________ Area: ______________________

II.6 *Erosion* is a process that breaks down and transports crustal material, usually by wind, precipitation, ice, or liquid. Erosional features can result from removal of crust, or from accumulation. Identify an example of an erosional feature on Mars. Show your work when computing the feature's area in km^2. **Zoom in as necessary**.

Lon/Lat: ___________________ Area: ______________________

MOLA (in the name of the layer) stands for Mars Orbiter Laser Altimeter, an instrument on the Mars Global Surveyor spacecraft of the late 1990's and early 2000's. The instrument measured the elevation of the Martian surface at every point by shining a laser at the surface and timing how long it took to receive a reflected signal.

Load the MOC Atlas 256ppd layer. MOC stands for Mars Orbiter Camera, an instrument also on the Mars Global Surveyor spacecraft which took images of Mars at visible wavelengths.

II.7 What differences do you see in the two different layers described above?

II.8 In which layer is it easiest to identify geologic features? Why do you think that is?

Part III. Martian Topography

Now add and make visible the MOLA Shaded Relief / Colorized Elevation layer. This is identical to the MOLA layer you should already have loaded, but now the map is colored by the elevation of the surface, with warm colors indicating high elevations and cool colors indicating low elevations.*

III.1 What is the major geological difference between "high" elevations and "low" elevations? Ignore the large volcanic region (Tharsis) when answering this question.

III.2 How are high elevations distributed, geographically?

*While pretty, a "rainbow" colorbar is not the best choice for display of scientific data. Rainbow colorbars can be misleading ("warm colors mean warm places on the planet, right?"), are difficult for colorblind people to interpret, and have been widely demonstrated to lead to misinterpretation of data.

III.3 What are the maximum and minimum elevations you can find in the map? What are the longitude and latitude of these locations?

High point: ____________________ Lon/Lat: _____________________

Low point: ____________________ Lon/Lat: _____________________

III.4 Compute the "vertical relief" on Mars, which is the maximum difference in elevation. Express your answer as a fraction of the radius of Mars (3400 km). Comment on whether Mars is "rough" or "smooth", in your opinion.

Teach yourself to create a line plot of elevation along a line segment. Do this by double clicking on the MOLA colorized elevation layer, then clicking on "chart" at the bottom of the pop-up window that appears. We recommend also clicking "Dock me" at the bottom of the pop-up window. Make sure you are using the "Selection" tool. Now, when you double click in the map, you can draw a line segment that shows the elevation along the segment you just drew.

III.5 About how deep is Valles Marineris (288° E, 11° S), the largest canyon in the solar system? Compare your answer to the 1.86 km depth of the Grand Canyon.

III.6 Use the tools you've developed to estimate the *volume* of Olympus Mons (227° E, 18° N), the largest volcano in the solar system. Use units of km^3. Show your work below.

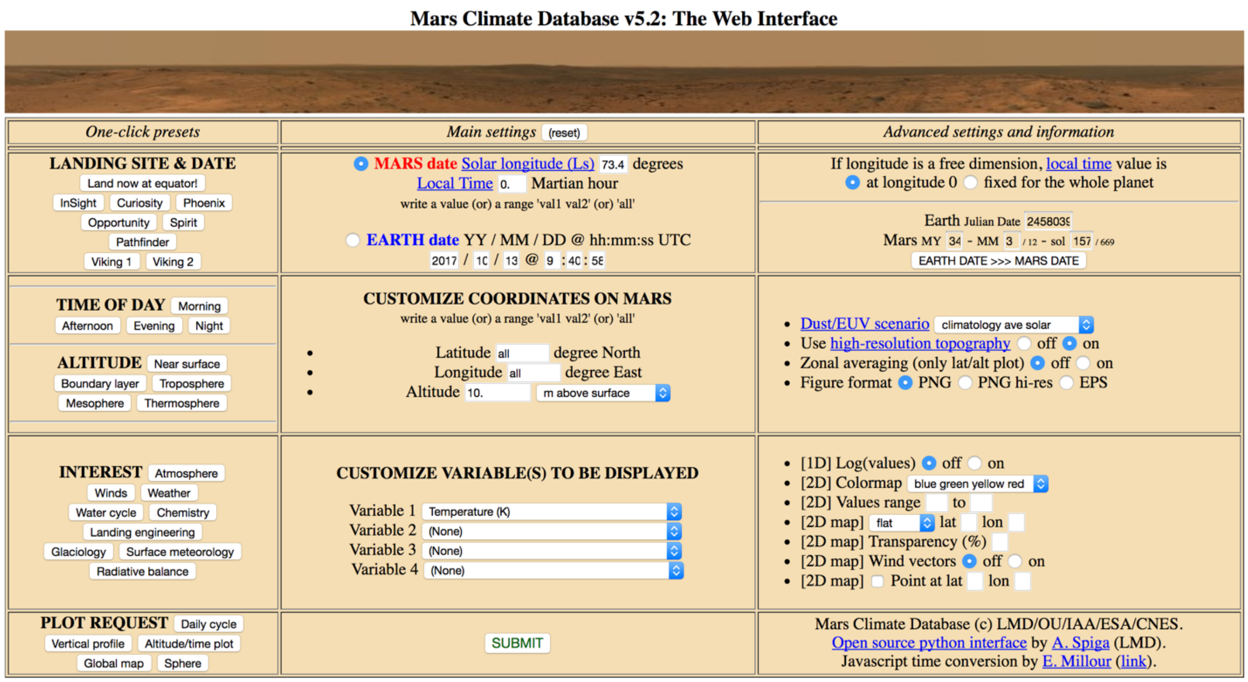
Part IV. Planetary Exploration

IV.1 Pick something fun to do with JMARS. This could include investigating a different layer on Mars, investigating a different planet, or something else! Tell us about (1) what you did, and (2) what you learned as a result.

Part V. Martian Atmosphere

Now let's turn our attention to the atmosphere of Mars. Go to the interface for the LMD GCM (atmospheric model), as described on the first page of the lab.

You should see an interface like this:



- In the middle column of the interface, set the Mars date to Solar longitude of 0 degrees.
- In the middle column of the interface, set the altitude to 0 m above the surface
- In the middle column of the interface, choose variables 1-3 as Temperature, Pressure, and density
- In the right column of the interface, choose the colormap to be blue or gray.
- Click "Submit" to obtain maps showing the results from the model.

V.1 Which longitude corresponds to midnight?

V.2 What is the most important factor controlling the pressure of the atmosphere at the surface of Mars?

V.3 How are surface pressure and density related to each other, qualitatively?

Now, without closing the tab you just created, go back to the web interface for the model and change the Ls to 90 and click submit. Then do the same for Ls=180 and 270. You should now have 4 open tabs with model results.

V.4 Based on the maps, which Ls corresponds to the northern summer solstice? Why do you conclude that?

V.5 Surface pressure is a measure of the weight of the atmosphere above the surface – or how much atmosphere a planet has. Does Mars have more atmosphere in northern summer or southern summer?

V.6 Where might the extra atmosphere come from? (Hint: You might want to look back at the MOC layer from JMARS)

Finally, using the button in the left column of the interface, create vertical profiles of temperature and pressure at 40 degrees south latitude at noon on Ls=270.

V.7 How does pressure vary as a function of height, qualitatively?

V.8 How does temperature vary as a function of height, qualitatively?

V.9 How does the temperature profile change when we change to Ls=90? Why?

Name: ______________________________

SURFACE OF MARS | Post-Lab = please hand in at the start of next lab

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SPECTRAL BARCODES°

Name: ____________________

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Read briefly through the entire lab. Check this box ☐ to indicate that you have done so. Jot down any questions that occurred to you while skimming through the lab description.

1) When electrons move down in energy levels, are they gaining or losing energy? If gaining, where did this energy come from? If losing, where does the energy go?

2) How does an incandescent light bulb differ from a fluorescent light bulb? Should you expect their spectra to look different?

3) How can a spectrum be used to identify an unknown gas? Why are spectra often referred to as 'fingerprints' of a gas?

Name: ____________________________

SPECTRAL BARCODES

"I ask you to look both ways. For the road to a knowledge of the stars leads through the atom; and important knowledge of the atom has been reached through the stars." - A.S. Eddington

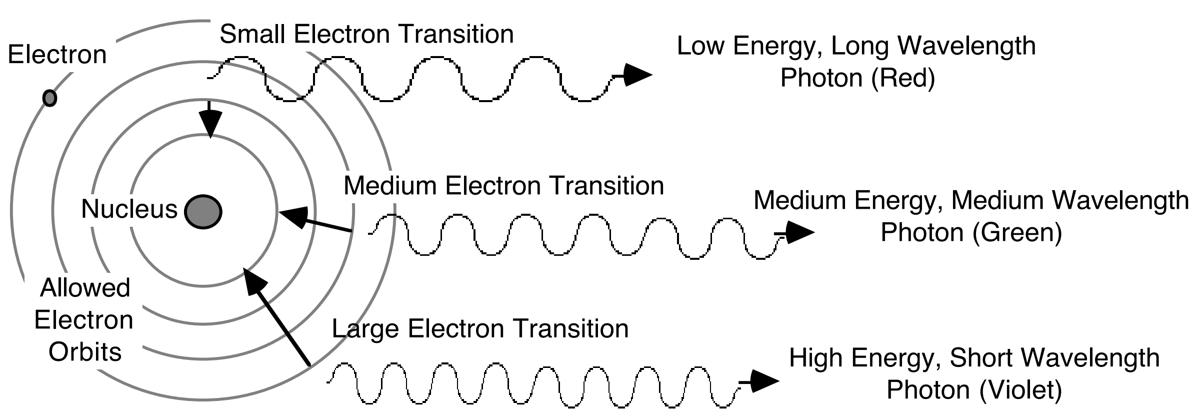
SYNOPSIS: Many objects in astronomy need to be studied from a distance by means of visible or invisible light (infrared; ultraviolet; etc.) What can we learn about astronomical objects from their light? What does light tell us about the chemical composition of the object that produced it?

EQUIPMENT: Hand-held spectroscope, spectrum tube power supply, helium, neon, nitrogen, air, and "unknown" spectrum tubes, incandescent lamp, heliostat.

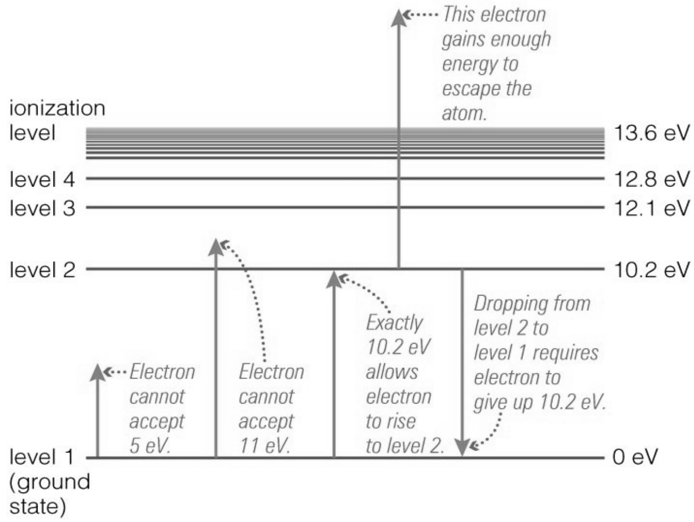
Reminder – What Is Spectroscopy?

Most of what astronomers know about stars, galaxies, nebulae, and planetary atmospheres comes from **spectroscopy**, the study of the colors of light emitted by such objects. Spectroscopy is used to identify compositions, temperatures, velocities, pressures, and magnetic fields.

An atom consists of a nucleus and surrounding electrons. An atom emits energy when an electron jumps from a high-energy state to a low-energy state. The energy appears as a **photon** of light having energy exactly equal to the *difference* in the energies of the two electron levels. Since each element has a different electron structure, and therefore different electron energy states, *each element emits a unique set of spectral lines.*



Part I – Electron Energy Transitions



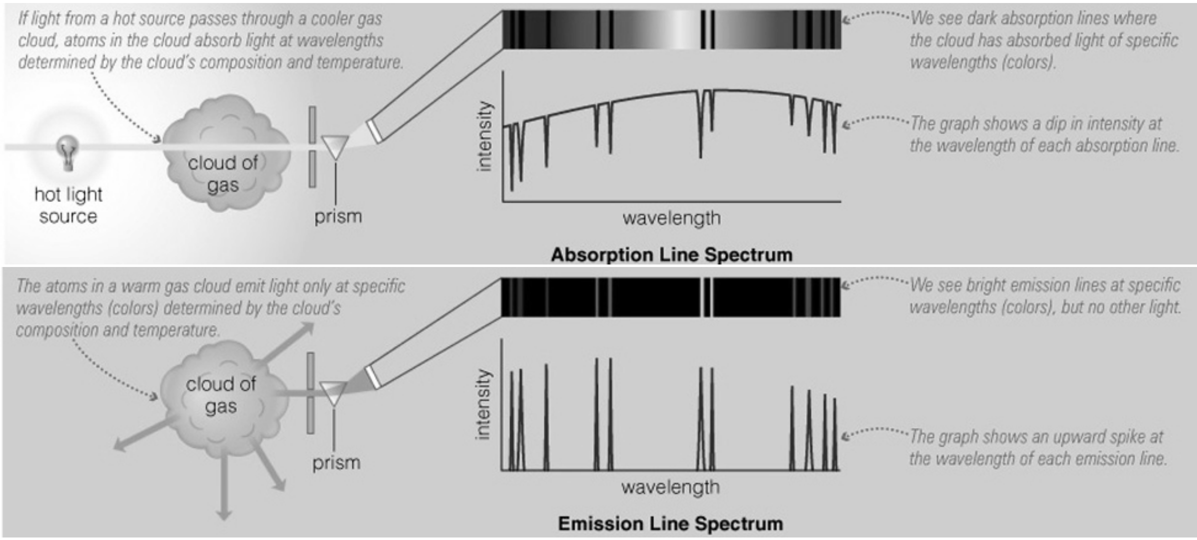
Energy levels for an electron in a hydrogen atom.

I.1 If an electron drops from *n=3* energy level to the *n=2* energy level of atomic hydrogen, how much energy is it getting rid of? Where does that energy go?

I.2 That particular *n=2* to *n=1* transition is called the "Balmer Hα line," and it's pretty famous in astronomy. Calculate the wavelength of light corresponding to that transition. Is this wavelength visible to our eyes?

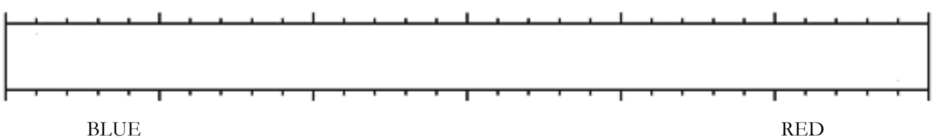
I.2 Pick a transition of energy levels that an electron could make that would *not* be visible to human eyes. What is the change of energy in that transition and to what wavelength does it correspond?

Part II - Continuum and Emission Line Spectra



Look at an ordinary lamp. Can you see all of the colors of the spectrum, spread out left to right? If the colors go up and down rotate your grating 90 degrees. You should see the familiar rainbow of colors you saw with the diffraction grating slide you used in the Light and Color Lab. Ask for TA/LA help if you don't see this.

II.1 Look through the spectroscope at the *incandescent lamp* (regular light bulb) and sketch the spectrum:

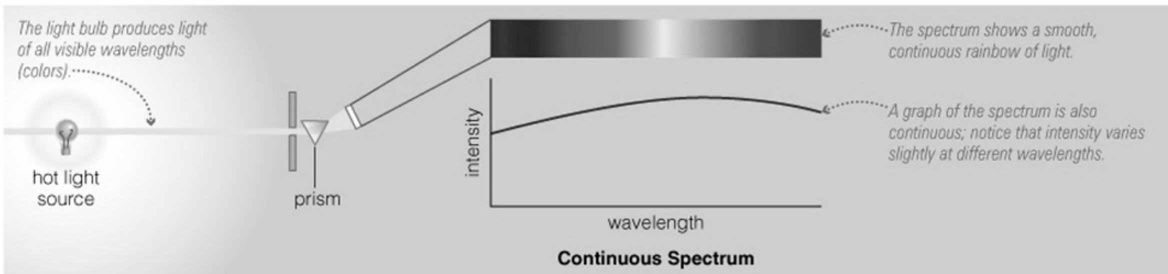


(a) Describe (in very basic terms) what you see.

(b) Are there any discrete spectral lines?

(c) What color in the spectrum looks the brightest? Or what color is in the middle of the bright part of the spectrum?

A *solid* glowing object, such as the filament of a regular light bulb, will not show a characteristic atomic spectrum, since the atoms are not free to act independently of each other. Instead, solid objects produce a **continuum spectrum** of light regardless of composition; that is, *all* wavelengths of light are emitted rather than certain specific colors.



Continuous Spectrum

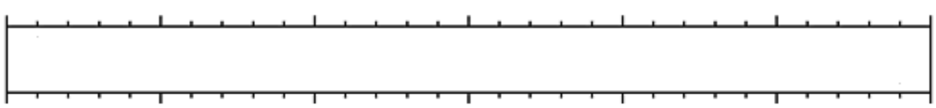
For each of these gases (Helium, Neon and Nitrogen):

- Install the element discharge tube in the power supply and turn it on
- Look through the spectroscope at the gas tube.
- Turn off the power supply before changing tubes.

> *CAUTION! The tubes are powered by 5000 volts! Do NOT touch the sockets when the power supply is on. The tubes also get very HOT! Let the tubes cool, use paper towels to handle them.*

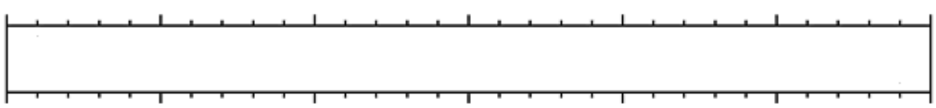
II.2 What colors do you see? Use colored pencils or crayons to sketch spectral lines of the light emitted by the element you are looking at on the frames below.

Helium



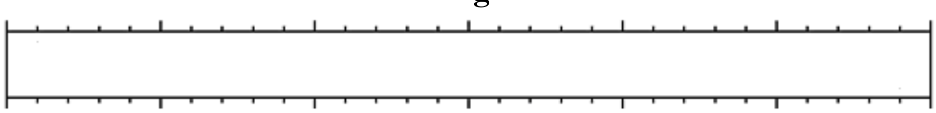
BLUE RED

Neon



BLUE RED

Nitrogen



BLUE RED

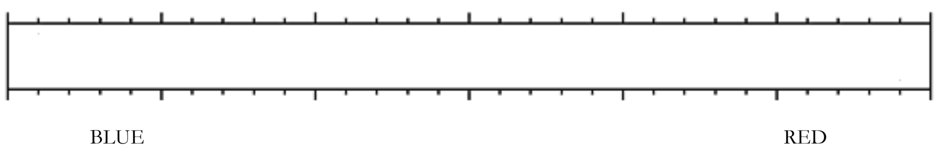
II.3 For each of these elements, how does the overall color of the glowing gas compare with the specific colors you see split apart in its spectrum?

- Helium
- Neon
- Nitrogen

II.4 Judging from the number of visible energy-level transitions (lines) you see in each, which element would you conclude has the more complex atomic structure: helium or neon? Explain.

Fluorescent lamps operate by passing electric current through a gas in the tube, which glows with its characteristic line emission spectrum. A portion of that light is then absorbed by the solid material lining the tube, causing the solid to glow, or **fluoresce,** in turn.

Point your spectroscope at a *fluorescent* light and sketch its spectrum.



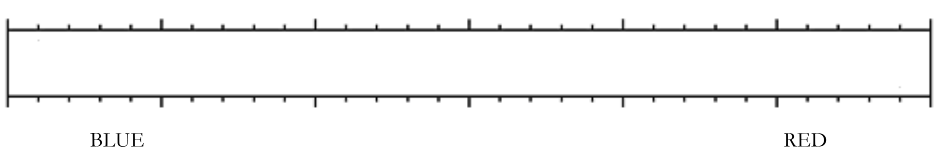
II.5 Which components of the spectrum originate from the gas?

II.6 Which components of the spectrum originate from the solid?

Part III – Identifying an Unknown Gas

Select one of the unmarked tubes of gas (it will be either hydrogen, mercury, or krypton). Install your "mystery gas" in the holder and inspect the spectrum.

III.1 What is the color of the glowing gas? Make a sketch of the spectrum and label the colors.



III.2 Identify the composition of the gas in the tube by comparing your spectrum to the spectra described in the tables on the next page. Explain which lines you used to make your identification.

The strongest lines are **bold**. Wavelengths in nm (10^{-9} m) are given to left of each color.

Hydrogen	
656	**Red**
486	**Blue-Green**
434	Violet
410	Deep Violet (dim)

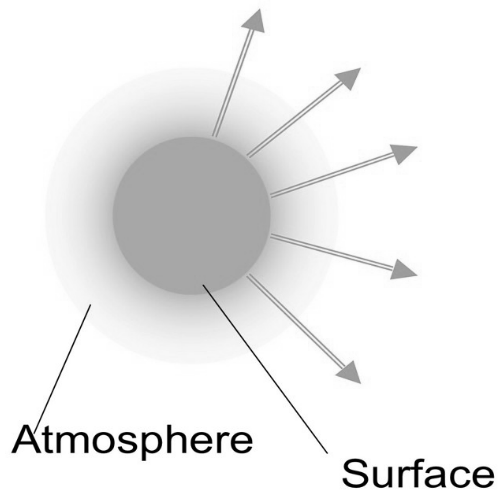
Mercury	
607	Orange
578	**Yellow**
546	**Green**
492	Blue-Green (dim)
436	**Violet**
405	Violet (dim)

Krypton	
646	**Red**
587	**Yellow-Orange**
557	**Green**
450	Violet (dim)
446	Violet (dim)
437	Violet (dim)
432	Violet (dim)
427	Violet (dim)

Part IV – Solar Spectrum

If the Sun is shining, the TA will use the Heliostat to bring up the solar spectrum. This involves using mirrors, lenses and a grating to pipe in sunlight from outside and to split the light by wavelength.

IV.1 What do you see? Describe the solar spectrum in terms of *continuous, emission* and/or *absorption* components.



IV.2 Based on the (extremely simplistic) model of the Sun above, which component of the spectrum comes from the Sun's surface? Which is due to its atmosphere?

IV.3 Your TA will also put light from a couple of gas tubes through the same optics that will produce emission lines above/below the solar spectrum. Can you identify these gases in the solar spectrum?

IV.4 How many lines of hydrogen can you find in the solar spectrum? (List which ones.)

IV.5 Have your TA identify the sodium absorption lines. What color are they in? What color are sodium (emission) lamps? *(Hint: Think back to the last Light and Color lab.)* Explain the reason for the similarity.

Name: ______________________________

SPECTRAL BARCODES | Post-Lab = please hand in at the start of next lab

1. **Metacognition:** We are always trying to improve the ASTR1030 labs and value your input on where they are succeeding and where they could be improved. Please briefly describe, what did you like and/or not like about this lab? What main idea(s) do you feel you learned from this lab?

COLLISIONS, SLEDGEHAMMERS, AND IMPACT CRATERS

Name: ____________________

Pre-Lab = please complete *before* coming to lab

Read briefly through the entire lab. Check this box ☐ to indicate that you have done so. Jot down any questions that occurred to you while skimming through the lab description. Your pre-lab here is slightly more substantial, to give more time for playing with sledgehammers during lab.

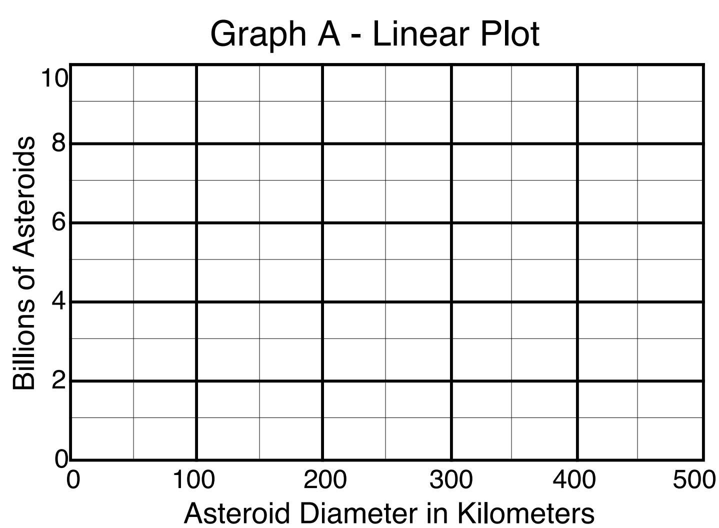
Size Distribution of Asteroids

Diameter	*Number*
500 km	Three (3×10^0)
250 km	Ten (1×10^1)
100 km	130 (1.3×10^2)
10 km	5,000 (5×10^3)
1 km	1,000,000 One Million (10^6)
0.05 km	10,000,000,000 Ten Billion (10^{10})

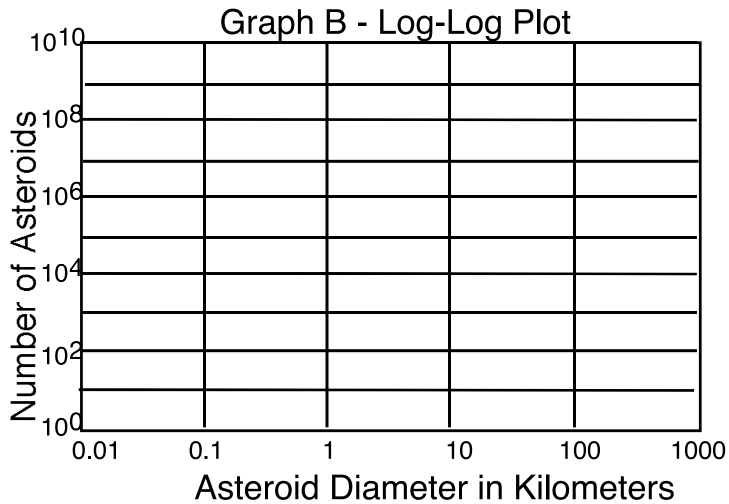
Power law distributions: Numbers and sizes of asteroids in the asteroid belt are not random, but rather exhibit a fairly well behaved and predictable pattern. For example, smaller asteroids are much more numerous than larger ones. Only three asteroids in the belt have diameters exceeding 500 km, yet twelve have diameters greater than 250 km, and approximately 150 asteroids are greater than 100 km across. Thousands of asteroids tens of kilometers in size have been catalogued. There are also uncountable numbers of smaller ones going all the way down to grain-sizes. The term given to this relationship between number and size in such a system is the **size distribution**.

1. Using Graph A (a **linear plot**), try to graph the size distribution of all of the asteroids in the list above.

   You will probably find that plotting these data on this simplistic scale is extremely difficult. The range of the numbers involved is simply far too large to conveniently be displayed in any meaningful manner on a linear plot.



2. Now, plot the same numbers again, but this time using the scale provided with Graph B (a **log-log plot**).



Graph B is a **logarithmic** plot, in which both the x- and y-axes are in increasing powers of 10. The use of logarithmic scales commonly enables you to accommodate the full range of the numbers involved.

3. In Graph B, how many orders of magnitude are there on (i) the x-axis? (ii) the y-axis? How does this compare to the rough number of orders of magnitudes you could *meaningfully* portray on the linear axes (Graph A).

Name: ______________________________

COLLISIONS, SLEDGEHAMMERS, AND IMPACT CRATERS

SYNOPSIS: The objectives of this lab are: (a) become familiar with the size distribution of particle fragments resulting from collisions; (b) compare that distribution with that of interplanetary debris found in the asteroid belt; and (c) relate the size distribution of craters on the Moon to the size distribution of fragments in the solar system.

EQUIPMENT: Sledgehammer, brick, denim cloth, sieves, plastic bags, buckets, scale, safety goggles, calculator, graph overlays.

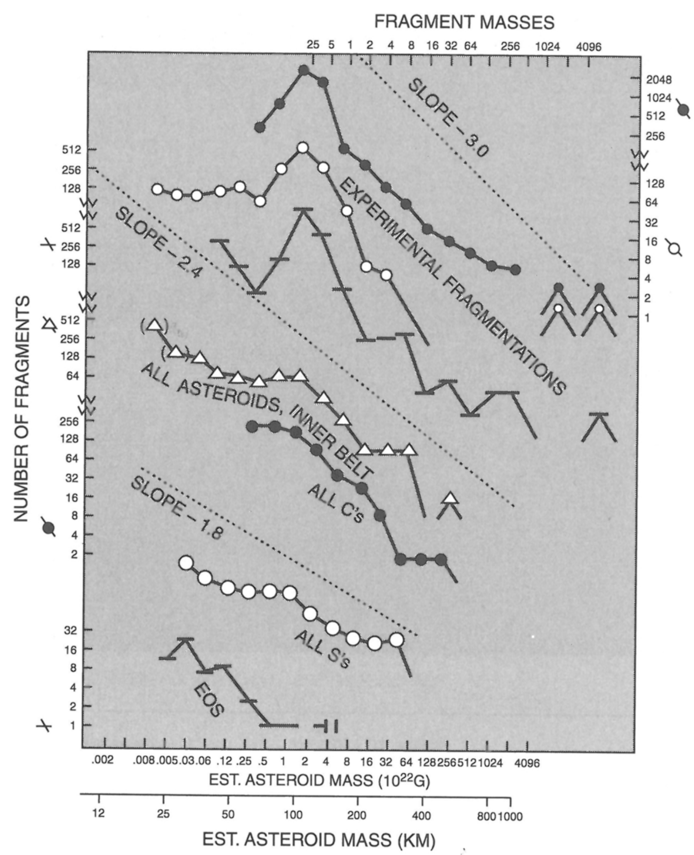
Part I – Size Distributions

The occurrence of asteroid sizes in the Solar System approximates a **power law** of the form

$$N = A\ R^{b}.$$

Here N is the number in a given radius (R) interval on the logarithmic scale, and A is a constant of proportionality. When plotted on a log-log plot, the distribution of objects that follow a power law behavior yield a *straight line*, the *slope* of which is equal to the power law exponent *b*. For objects in the asteroid belt, *b* has a value of approximately –2. This power law distribution of relative abundance persists over many orders of magnitude.

In the simplest terms, this mathematical relationship means that there are many more small fragments than large fragments resulting from disruptive collisions in the asteroid belt. The surprise is that such a general trend should be so precise that it holds true over objects differing in size by 100 million! The figure at right shows size distributions of impact-shattered rocks, both from the result of laboratory fragmentation experiments and the actual distributions of asteroids. The data points for the "All Asteroids, Inner Belt" population correspond roughly to b ≈ –2.



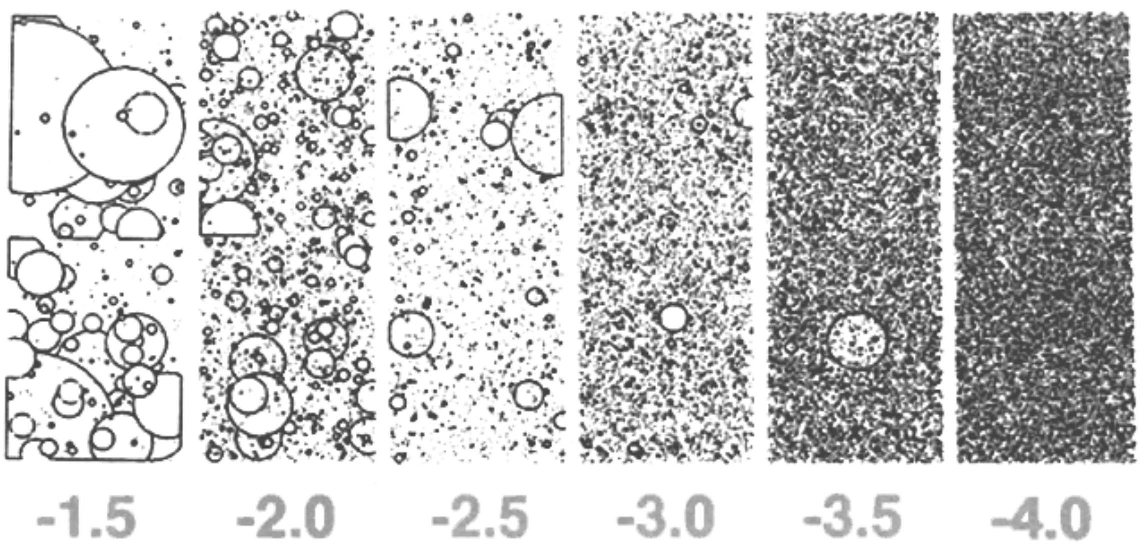
(From Moons and Planets, W. Hartmann, 4th Ed. 1999)

The power law distribution of the sizes of asteroids therefore suggests a **collisional fragmentation process**, the consequences of which are fascinating. The sizes continue getting smaller and smaller and the numbers continue to become greater and greater.

The equation $N = AR^b$ describes a relationship between the number of objects, *N*, and their radius, *R*. But what are the effects of changing the values of the parameters *A* and *b*?

I.1 If the value of *b* becomes "more" negative, does the slope get more or less steep?

Below is a diagram made using a computer simulation of different size distributions for a range in the values from b = –1.5 to –4.0.



I.5 Which value of b creates relatively more big particles? Which creates relatively more small particles? Explain why, in terms of the *slopes* of the particle size distributions.

Part II. Destructive Learning

You will test the hypothesis that asteroid size distributions are the result of collisional processes by simulating such collisions for yourself, using a brick as a rocky asteroid and a sledgehammer to provide the impact(s).

It is anticipated, however, that you will be producing far too many small "asteroids" to count one at a time. To overcome this limitation, it is possible estimate their number, N, by calculation, using the **density** of your original "asteroid" as a guide.

Density is a measure of the amount of mass in a given volume, and it is measured in units of mass (grams or g) per volume (cubic centimeters or cm^3). Water, for comparison, has a density of exactly 1 g/cm^3. Thus, one cubic centimeter of water would weigh one gram if placed on a metric scale. Moreover, if we had a container of water that weighed 100 grams, we would know that we had a volume of 100 cubic centimeters of water.

II.1 Use the metric scale to measure the mass (in grams) of your brick "asteroid":________

II.2 Measure the sides of the soft-brick:

Length (cm) =

Width (cm) =

Height (cm) =

II.3 Calculate the density of the soft-brick:

Now take your sledgehammer, brick "asteroid," goggles and denim cloth, and find a safe place outside for a smashing good time!

Step 1 Wrap the brick in one sheet of cloth, and spread the other out on the ground. Place the wrapped brick in the middle of the spread-out sheet. (*Note:* The cloth containing the brick will not last for more than a few hits before it rips; its purpose is to hold the pieces together as well as possible so that you won't lose any. Be careful so that any pieces that do come out stay on the other sheet.)

Step 2 Now, the fun part: smash your brick! You will most likely have to hit the brick about 4 to 6 times (representing 4 to 6 "collisions" with other asteroids) to ensure that you end up with enough small pieces for your analysis. (*Note:* The largest pieces you will be interested in are only one inch across.) ***Each member of your group should hit the brick at least one time. The person hitting the brick MUST be wearing the goggles!***

Step 3 Being careful not to lose any of the pieces, fold up the cloth sheet and bring your sample fragments to the sieves.

Part III. Sorting and Counting Your Fragments

Be sure to read and follow EACH step carefully. There are a total of five sieves of differing sizes: 2.54 centimeter diameter sieve, 1.27 cm, 0.64 cm, 0.32 cm, and 0.16 cm. The idea is to separate your material—by following the steps below—according to these sizes. You will begin by putting all the material into the *largest* sieve, thereby separating out the biggest pieces. Pieces larger than 2.54 centimeters in diameter will stay in the sieve while everything else will fall through. You will then use the *second* largest sieve, and so on down to the smallest.

III.1 Place the pan below the 2.54 cm sieve. Slowly pour the material into the sieve, gently agitating the sieve as you go. You may need to do a little at a time if there is too much material for the sieve. (Note: Do not be too rough with the sieves. By excessively shaking the sieve you may inadvertently cause unwanted further grinding.)

III.2 Separate your largest pieces—these will not be used in the analysis.

III.3 With the remaining material, use the next-sized sieve to separate out the next-largest pieces and place these in a baggie. These fragments are the largest you will consider in your analysis. Keep track of this material by placing a small piece of paper in the baggie that records the sieve size.

III.4 Repeat the process for each sieve in descending size order. Discard all material that falls through the smallest sieve.

III. 5 Weigh the material in each baggie with your balance scale. Record the results in column 6 of the table in III.7. (Note: You do not need to empty the material onto the scale. Instead, simply put the baggie on the scale. The small mass contribution from the baggie is negligible.)

The next step is to create a plot of the number of objects versus sieve size. Of course, one way to do this would be to count each fragment within each baggie. However, because we know the densities of our brick "asteroid" fragments, there is a much simpler way to estimate these numbers.

You have measured the *total* mass of all the objects in a certain size range (column 6 of the table). If you divide this number by the mass of a *single* object of that size, the result will be an estimate of the total number of fragments within that range (and within your baggie).

Assuming that each particle is approximately spherical in shape, and also that the average size of the particles in a baggie is halfway between the two sieve sizes that yielded the sample (the mean particle size from column 5 above).

III.6 Write down a formula for the mass of a single spherical object in terms of its density (rho) and radius (R):

III.7 Given your answer to III.6, calculate the mass of one object having a size equal to the average size of a particle collected in each sieve using the equation given on the previous page. In column 7 of the table, enter the mass of a single representative particle in each of your four bags. Show one example calculation below.

III.7 You now have enough information to estimate the total number of particles in each bag. Enter the estimated total number of particles in each size range in column 8.

Smashing Brick Data Table

1	*2*	*3*	*4*	*5*	*6*	*7*	*8*
Bag #	*Former Sieve Size (cm)*	*Current Sieve Size (cm)*	*Mean Particle Diameter (cm)*	*Mean Particle Radius* R *(cm)*	*Mass of Bits(+Bag) (grams)*	*Mass of One Particle* M_1 *(grams)*	*Number of Particles* N
1	2.54	1.27	1.91	0.95			
2	1.27	0.64	0.96	0.48			
3	0.64	0.32	0.48	0.24			
4	0.32	0.16	0.24	0.12			

III.8 To check the validity of your approach, actually *count* the number of fragments in one (or more) of the baggies. (You might choose the baggie containing the fewest and largest fragments, but that choice is up to you). Record your answer below and compare it with your estimate from column 8.

Part IV. Plotting and Analyzing Your Results

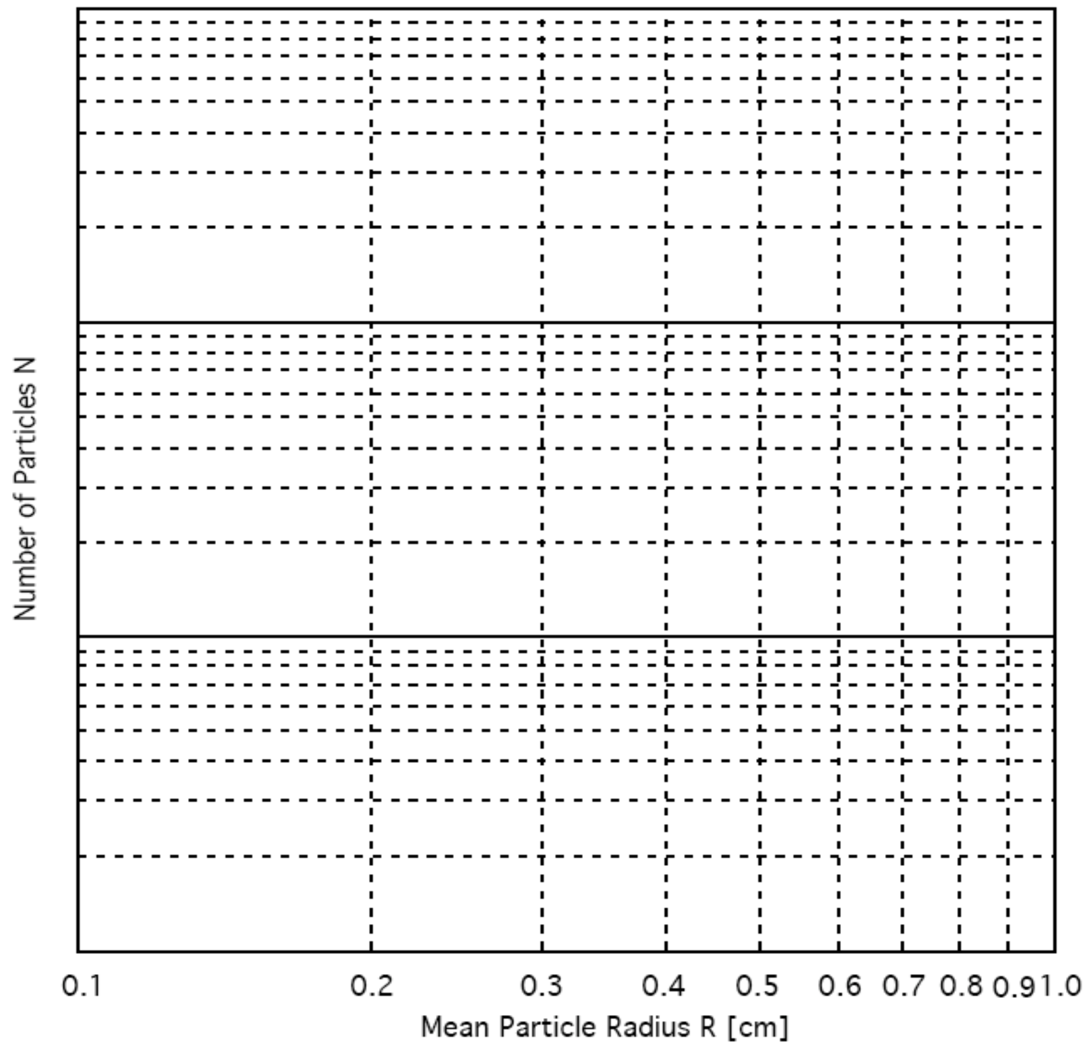
IV. 1 Using the log-log plot of the graph on the next page, plot your results showing the number of particles N versus the mean particle radius. (*Hint:* You will need to use your own data to come up with labels for the Y-axis. Remember this is a log-log plot!) Does your data form *roughly* a straight line? If not, can you think of any reason why not?

IV.2 Using the transparencies for different exponents (values of *b*) and overlaying them on your plot, find the one which best matches your plot. (Make sure you keep the X and Y axes of the overlay parallel to the X and Y axes of your graph, then judge by eye which slope seems closest to your data.) What is your estimate of *b* for your brick fragments?

IV.3 How does this compare with the value of *b* for the asteroid power law distributions? Does your power-law distribution from smashing particles tend to match those of asteroids? If your power-law distribution differs from that of the asteroids, explain whether your smashing experiment yielded too many large-sized, or too many small-sized particles.

IV.4 Compare your plot and your value of *b* to plots from other groups in your lab. What similarities and differences do you notice?

Smashing Bricks: The Asteroid Collision Simulation



IV.5 How do you think your plot would have changed if you had smashed the brick several more times? Would the slope change? Would the curve shift up or down? Explain your reasoning.

When finished, ***please clean up your lab station****. Replace all materials so the area appears as when you began. Have your TA check your station before you leave.*

Name: ______________________________

COLLISIONS | Post-Lab = please hand in at the start of next lab c

Asteroid Impacts on the Earth

Because both the Moon and the Earth occupy the same general region of the solar system, it is reasonable to assume that both have been bombarded by similar numbers and sizes of space debris. The only difference is that impacts on the Earth have been moderated somewhat by the atmosphere, and most of Earth's craters have been obliterated by geological activity (erosion, volcanism, and tectonics).

The surface geology of the Moon, on the other hand, has remained fairly undisturbed (except for impacts) since the last maria-building lava flows, which ended roughly 3.5 billion (3.5×10^9) years ago. Based on crater counts, approximately 30 objects 1 km in diameter have hit each 100,000 km^2 of the lunar surface in the last 3.5 billion years.

We can, therefore, use the lunar crater record to estimate the numbers and sizes of impacts that have occurred in the past on the Earth. The radius of the Earth is 6400 km.

1. Compute the surface area (in km^2) of the Earth.

2. How many times bigger is the Earth's surface area than the crater-counting standard area of 100,000 km^2 on the Moon?

3. So, in the same corresponding period of time that craters have accumulated on the Moon, about how many 1 km diameter impactors (asteroids or comets) have hit the Earth?

4. How frequently do impacts of 1 km objects occur on the Earth? (express this as an average number of years between impacts)

5. What is the approximate probability that a 1 km diameter object will strike the Earth in your *lifetime*? Such an impact would be pretty devastating for Earth, so it might help to check – does your answer make sense? If not, you may have made an error along the way…go back and check!

6. Based on the slope of your logarithmic graph (or assuming *b* = *-2* if you can't remember what you found in the lab), how frequently would you expect 1m size objects to impact the Earth?

7. What is the probability that a 1m diameter object will impact the Earth during your lifetime? Does your answer make sense?

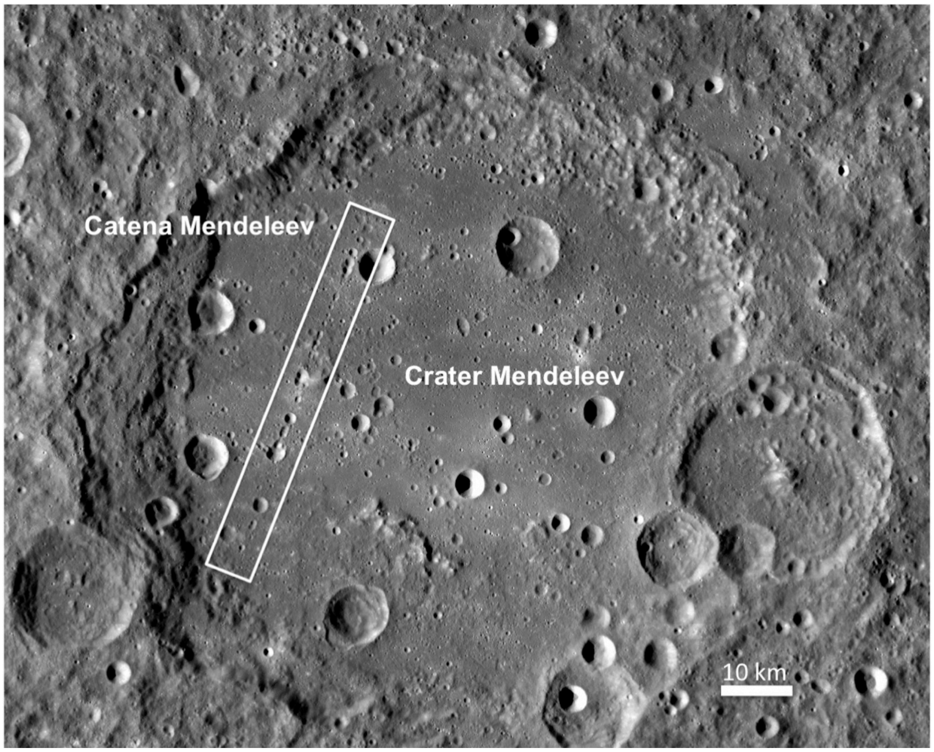


Image from the Lunar Reconnaissance Orbiter (launched in 2009) currently orbiting the Moon. How do you think the chain of craters called Catena Mendeleev might have been formed?

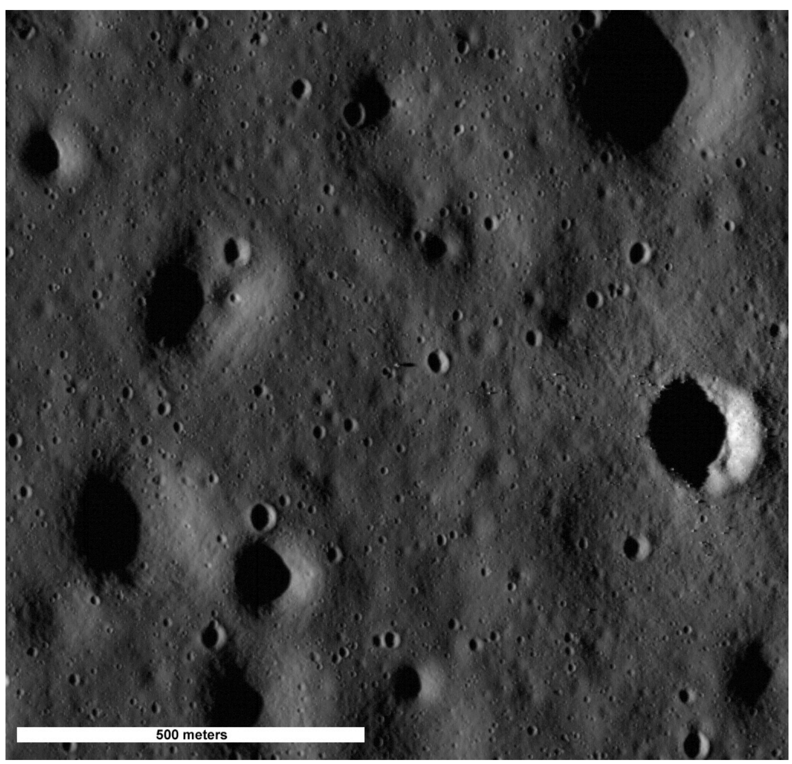


Image from the Lunar Reconnaissance Orbiter (launched in 2009) currently orbiting the Moon. That little dot in the middle of the bottom picture (casting a shadow) is the Apollo 11 Lunar Module!

Name: ______________________________

TRANSITING EXTRASOLAR PLANETS

Pre-Lab = please complete *before* coming to lab

Read briefly through the entire lab. Check this box ☐ to indicate that you have done so. Jot down any questions that occurred to you while skimming through the lab description.

1. Summarize Newton's form of Kepler's 3rd Law **in words**. Then state the law mathematically, explaining the meaning of each symbol in the equation.

2. Let's imagine an exoplanet transiting one of the closest stars to Earth. Draw a diagram that shows the exoplanet and its host star, during transit, as seen from Earth. (Pretend you have a magically good telescope so you can actually resolve the image.)

3. Draw another diagram, this time from above, showing the relative arrangement of that exoplanet, its host star, its orbit, and the Earth. If you have to make any approximations to sizes or scales to fit on the page, please explain what those are.

Name: ____________________________

TRANSITING EXTRASOLAR PLANETS

LEARNING GOALS: Detecting planets around other stars is a very challenging task. What is the transiting planet method of detection? What can we learn about extrasolar planets using this method?

EQUIPMENT: Lamp, ruler, Lego™ orrery (with a variety of detachable planets), light sensor, laptop (or other) computer with LoggerLite™ software, modeling clay (available from your lab instructor for the optional section at the end), calipers

In March 6, 2009, the NASA *Kepler* spacecraft successfully launched and began monitoring the brightness of more than 100,000 stars. The Kepler mission is NASA's first mission with the precision capable of finding Earth-size and smaller planets around other stars. In this lab, you will discover how Kepler's instruments work and what we can learn about extrasolar planets.

In this lab, you will occasionally be asked to predict (as a real scientist would) the outcome of an experiment before you try it. Make these predictions BEFORE moving on to the experiment itself. You will not be marked down if your predictions are wrong.

Unfortunately, the Lego orrery does not simulate a true solar system since it does not exactly follow Kepler's 3rd Law ($P^2=a^3$). Keep this in mind.

Part I. Setting up your Kepler Simulation

The transit detection method is an *indirect* detection method in that it is not directly detecting the planet itself but rather the planet's interactions with its central star. By detecting a repeating dimming of the star's brightness, scientists can infer that a planet is orbiting around the star and occasionally blocking some of the star's light from the telescope.

In this lab, the light sensor will simulate both the *Kepler* telescope and its light detecting hardware. The lamp will simulate the star and the Lego orrery will be configured to simulate various planets moving around that star.

The first thing you'll need to do is place the star in the middle of the orrery. Adjust the lamp so the bulb is over the middle shaft of the orrery. Be sure the base of the lamp is not blocking the path to the light sensor. Next you'll need to align the light sensor so it is pointing directly at the center of the light source.

I.1 Using your ruler, measure the height of the star and adjust your light sensor to the same height. Record the height of your star and light sensor: ________________

I.2 Next you'll need to make sure your light sensor is pointed directly at the center of the light source. Align the sensor roughly by eye, and then use LoggerLite software (if the software is not already running, ask your TA or LA to help start the program) to do it more accurately. Explain how you can use the LoggerLite software to align your sensor.

I.3 Record the value for the peak brightness of your lamp: ____________________
(The light bulb itself might show some small variability. As long as it's not periodic, this will not affect your measurements. Most real stars actually show some variability.)

Part II. Measuring the Effect of Planet Size

II.1 Affix a medium-sized planet to the middle arm of the orrery. Try to get the height of the planet to be the same height as the center of your star and your light sensor. Turn on the orrery motor and start the LoggerLite data collection. Describe the results.

II.2 (a) Suppose that your planet was ½ the diameter of your star. What percent of the star's light would you *predict* that planet would block?

(b) As seen from a distance, planets and stars look like circles. Draw a planet and a star on top of each other below with the planet having a diameter that is ½ that of the star (use circles, don't do a 3-dimensional drawing). To help make the point even clearer, temporarily pretend the star and planet are squares, with the smaller one ½ the width of the larger (start the square "planet" in the corner of the square "star"). Draw accurately!

Star & Planet Star & Planet (drawn as a squares)

(c) Based on your drawing, how does the *area* of the star compare to the area of the planet? (You should be able to give actual numbers here, not just bigger or smaller).

II.3 (a) Using the clamps provided, measure <u>your</u> star (in the orrery) and record its size here (be careful not to break your bulb!): ____________

(b) Now measure <u>your</u> planet and record its size here:________________

(c) What is the ratio of the diameters?_________________

(d) What is the ratio of the areas?________________

II.4 What percent of the star's light do you predict the planet will block?

II.5 (a) Use the experimental setup to measure the percentage of the light that is actually blocked. *Show your work*. (You can use the "Examine" button in LoggerLite to get the exact y-value at any point on your graph. Be sure to run your orrery for *at least* two complete orbits of the planet.)

(b) How well does your result agree with your predictions?

(c) What might be the cause(s) of any differences? Show your prediction and result to your TA or LA before you proceed.

II.6 Replace the medium planet with a different sized planet, run the orrery, and describe the results. Compare your results to those you found in the previous question.

Part III. Measuring the Effect of Planet Orbital Distance

III.1 Predict the effect of changing the orbital distance of the planet and record your prediction. Be as specific as possible. (*Hint: The spacecraft is called the Kepler spacecraft…*)

III.2 Move the planet to a different position, run the orrery for at least two orbits, and describe your results.

III.3 How well do your results agree with your prediction? If they disagree, what might be the cause(s) of any differences?

III.4 How is our lab setup different from real exoplanets around real stars observed from a real telescope on Earth? What effect could these differences have on our measurements?

Part IV. Measuring a Complex Planetary System

Split your lab group into two teams. Each team will take one turn acting as the extrasolar system creators and one turn acting as the *Kepler* Science Team. Fill in the appropriate sections when it is your turn to act as that team.

Extrasolar System Creators: Place the cardboard divider between your teams so the *Kepler* Science Team cannot see the orrery. Using the various planet choices, create a solar system consisting of up to three planets. You don't *have* to use all three but try to make it challenging!

IV.1 Record the sizes for the three planets you chose in the table below.

Planet	Size
1st Planet (closest to star)	
2nd Planet	
3rd Planet	

IV.2 In the space below, draw your *prediction* for what the Kepler light curve will look like. Explain, in words, your prediction.

When you are ready, turn on the orrery and tell the *Kepler* Science Team to begin their analysis.

Kepler Science Team: Your job will be to act as the scientists analyzing Kepler's data here on Earth. Without seeing the orrery, you will need to determine what kind of solar system the Kepler satellite has discovered.

IV.3 In the space below, make a sketch of the detected light curve. You might need a few minutes of data to recognize the full pattern. (If you want, you can print out your light curve and attach it to your lab write-up.)

IV.4 Based on the detected light curve, what are the sizes and distances of the planets around the system you've detected? Be as specific as possible (can you guess exact sizes?) Explain your reasoning.

IV.5 Once you've completed your analysis, check with the other team to see what actual planets were used. Was your analysis correct? If not, why not?

Now switch roles with the others; create a new solar system, and let the others analyze it. If you were the Kepler astronomers, you are now the creators, and should go back and fill in IV.1 and IV.2.

Part V. Summary

V.1 What are the difficulties that might be associated with detecting planets using the transit method? There are several answers to this question; you should list at least two for full credit.

V.2 For what types of extrasolar planets would the transit method work best? Large planets or small planets? Planets close to their host star or far from their host star? Highly eccentric orbits or circular orbits? Stars close to Earth or far away? Explain your reasoning.

V.3 If you had a spectrograph instead of a light sensor, how could this method be used to tell if a transiting planet had an atmosphere?

Part VI. Detecting Earth-size Planets

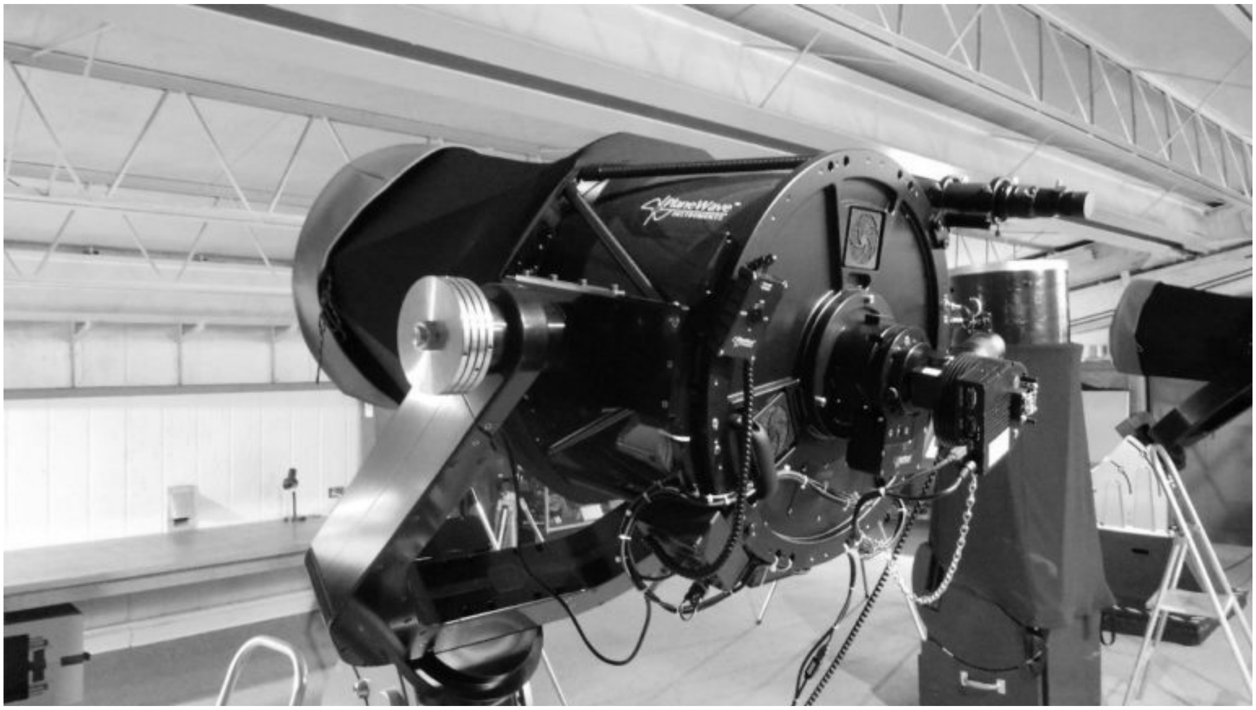
VI.1 Earth's radius is 6,400 km and the Sun's radius is 700,000 km. Calculate the percentage of the Sun's light that Earth would block during a transit.

VI.2 The *Kepler* spacecraft is able to measure individual brightness changes of as little as 0.002%, and by averaging together many measurements can sometime reach a measurement precision of 0.0002% (two parts per million)! What do you estimate to be the smallest brightness change you could reliably detect with *your* light sensor and model?

If you wish to explore the concepts a little further, your TA has modeling clay available. Create your own planets and predict what the light curve will look like. Some outcomes may surprise you!

Please clean up your lab station before you leave and leave it setup for the next lab group.

NIGHTTIME OBSERVING PROJECTS



Artemis and Apollo are the two 20-inch optical telescopes at Sommers-Bausch Observatory deck. (Photo by Gary Garzone, BASS)

Messier 13, a bright Globular Star Cluster in Hercules imaged with the CCD camera on Apollo, the western SBO 20" telescope.

KNOWING CONSTELLATIONS & BRIGHT STARS*

SYNOPSIS: In this self-paced lab, you will teach yourself to recognize and identify a number of constellations, bright stars, planets, and other celestial objects in the current evening sky.

EQUIPMENT: A planisphere (rotating star wheel) or other star chart. A small pocket flashlight may be useful to help you read the chart.

Part I: Preparation, Practice, and Procedures

Part II contains a list of 30 or more celestial objects that are visible to the naked eye this semester. You are expected to learn to recognize these objects through independent study, and to demonstrate your knowledge of the night sky by identifying them.

Depending upon the method chosen by the course instructor, you may have the opportunity to identify these objects one-on-one with your lab instructor during one of your scheduled evening observing sessions. Alternatively, there may be the opportunity to take an examination over these objects in Fiske Planetarium near the end of the semester. In addition, there may be the opportunity to do both, in which case the better of your two scores would be counted.

If you are given the option, and if you wait until the end of the semester to take the verbal quiz and then are clouded out, you have no recourse but to take the Planetarium exam. Do not expect your TA to schedule additional time for you. If you have not taken the oral test and are unable to attend the special exam session at Fiske (or it is not offered), you will not receive credit for this lab! *"Poor planning on your part does not constitute an emergency on our part."*

You can learn the objects by any method you desire:

- Independent stargazing by yourself or with a friend.
- Attending the nighttime observing sessions at the Observatory and receiving assistance from the teaching assistant(s) or classmates.
- Attending Fiske Planetarium sessions.
- All of the above.

Observing and study tips:

- It is generally to your advantage to take the nighttime verbal quiz if it is an option, since *you* control the method and order of the objects to be identified. In addition, it is easier to orient yourself and recognize objects under the real sky rather than the synthetic sky of the planetarium, since that is the way that you learned to recognize the objects in the first place.
- Your textbook may come with access to planetarium software, or you can download free software from http://stellarium.org/. These tools can be a great way to explore the night sky even during the daytime and/or from the comfort of your desk. Then you can go outside to see how the real thing compares with the simulation.
- If you like a hands-on aid, we recommend using the large, 10" diameter Miller planisphere available from Fiske Planetarium: it is plastic coated for durability, is easiest to read, and includes sidereal times. The smaller Miller planisphere is more difficult to use but is handier to carry. Other planispheres are available from the bookstore or area astronomy stores.

- To set the correct sky view on your planisphere, rotate the top disk until the current time lines up with the current date at the edge of the wheel. Planispheres indicate local "standard" time, not local "daylight savings" time. If daylight savings time is in effect, *subtract* one hour from the time before you set the wheel. (For example, if you are observing on April 15th at 11 p.m. Mountain Daylight Time, line up 10 p.m. with the April 15th marker).

- The planisphere shows the current appearance of the entire sky down to the horizon. It is correctly oriented when held overhead so that you can read the chart, with North on the chart pointing in the north direction. The center of the window corresponds to the zenith (the point directly overhead). When you face a particular direction, orient the chart so that the corresponding horizon appears at the bottom. As with all flat sky maps, there will be some distortion in appearance, particularly near the horizons.

- Be aware that faint stars are difficult to see on a hazy evening from Boulder, or if there is a bright moon in the sky. On the other hand, a dark moonless night in the mountains can show so many stars that it may be difficult to pick out the constellation patterns. In either case, experience and practice are needed to help you become comfortable with the objects on the celestial sphere.

- Learn relationships between patterns in the sky. For example, on a bright night it may be virtually impossible to see the faint stars in the constellation of Pisces; however, you can still point it out as "that empty patch of sky below Andromeda and Pegasus". You can envision Deneb, Vega, and Altair as vertices of "the Summer Triangle", and you might think of Lyra the harp playing "Swan Lake" as Cygnus flies down into the murky pool of the Milky Way.

- If you merely "cram" to pass the quiz, you will be doing yourself a great disservice. The stars will be around for the rest of your life; if you *learn* them now rather than just *memorize* them, they will be yours forever.

Good luck, good seeing, and clear skies!

Part II. Spring Naked-Eye Observing List

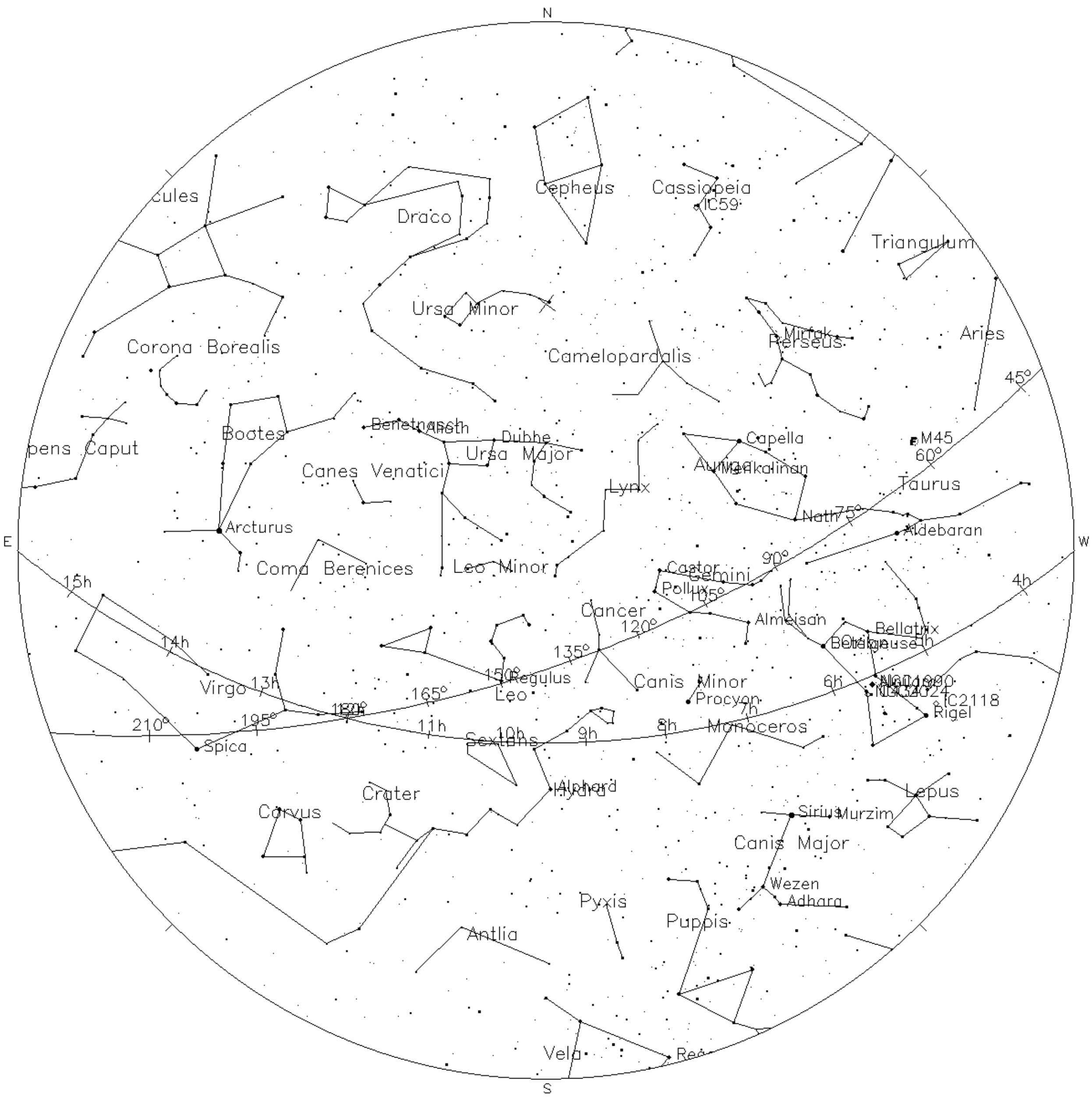
Different stars are visible at night at different times. Be sure to use the right list for the semester. If you're taking class in the fall, please skip ahead to the Fall Naked-Eye Observing List.

Constellations	Bright Stars
Ursa Major *(big bear, big dipper)*	Alcor & Mizar
Ursa Minor *(little bear, little dipper)*	Polaris
Cassiopeia *(queen, 'W')*	
Cygnus *(swan, northern cross)*	
Perseus *(hero, wishbone)*	
Taurus *(bull)*	Aldebaran
Auriga *(charioteer, pentagon)*	Capella
Orion *(hunter)*	Betelgeuse, Rigel
Gemini *(twins)*	Castor, Pollux
Canis Major *(big dog)*	Sirius
Canis Minor *(small dog)*	Procyon
Cancer *(crab)*	
Leo *(lion)*	Regulus
Boötes *(herdsman, cone, kite)*	Arcturus
Virgo *(virgin)*	Spica

Other Celestial Objects or Regions

Mercury, Venus, Mars, Jupiter, Saturn (and bright comets if present)	*(check the web resources at the beginning of this manual)*
Ecliptic or Zodiac	*trace its path across the sky*
Celestial Equator	*trace its path across the sky*
Pointer Stars	*how to locate Polaris*
Great Andromeda Galaxy	*fuzzy patch in Andromeda*
Pleiades *(seven sisters)*	*star cluster in Taurus*
Hyades *(closest star cluster)*	*near Aldebaran*
Great Nebula in Orion (M42)	*center 'star' in Orion's sword*

Spring Constellations
(with ecliptic and celestial equator shown)



The Boulder Night Sky
14 March, 10:00 pm

Part II. Fall Naked-Eye Observing List

Different stars are visible at night at different times. Be sure to use the right list for the semester. If you're taking class in the fall, please go back to the Spring Naked-Eye Observing List.

Constellations	Bright Stars
Ursa Minor *(little bear, little dipper)*	Polaris
Lyra *(lyre, harp)*	Vega
Cygnus *(swan, northern cross)*	Deneb
Aquila *(eagle)*	Altair
Cepheus *(king, doghouse)*	
Capricornus	
Aquarius	
Pegasus *(horse, great square)*	
Andromeda *(princess)*	
Pisces *(fishes)*	
Aries *(ram)*	
Cassiopeia *(queen, 'W')*	
Perseus *(hero, wishbone)*	
Taurus *(bull)*	Aldebaran
Auriga *(charioteer, pentagon)*	Capella
Orion *(hunter)*	Betelgeuse, Rigel
Gemini *(twins)*	Castor, Pollux

Other Celestial Objects or Regions

Mercury, Venus, Mars, Jupiter, Saturn (and bright comets if present)	*(check the web resources at the beginning of this manual)*
Ecliptic or Zodiac	*trace its path across the sky*
Celestial Equator	*trace its path across the sky*
Pointer Stars	*how to locate Polaris*
Great Andromeda Galaxy	*fuzzy patch in Andromeda*
Pleiades *(seven sisters)*	*star cluster in Taurus*
Hyades *(closest star cluster)*	*near Aldebaran*
Great Nebula in Orion (M42)	*center 'star' in Orion's sword*

Fall Constellations
(with ecliptic and celestial equator shown)

The Boulder Night Sky
15 October, 10:00 pm

OBSERVATIONS OF SATURN AND ITS MOONS°

MOTIVATION and SYNOPSIS: This is a three-part observational lab in which you will learn various aspects of observational astronomy, basic data analysis, and apply your results in a realistic way to determine some real aspects of planetary science.

In **Part 1**, you and a group of 2-3 other students will observe Saturn through the telescopes at the Sommers-Bausch Observatory (SBO) on campus both visually (through the eyepiece) and using a camera to record a detailed digital image.

In **Part 2**, each student will use software to perform some basic analysis of the image(s) you took at the telescope. You will measure four different quantities from your image and record them on a class data repository for future use. Part 2 is to be completed after your observing session (part 1). Your TA will give you a due date for Part 2.

Part 3 of this lab will be completed during one of your regular day-time lab sessions in late-October/early-November. You will use the data collected by the entire class (including your data points) to determine the orbit of Titan around Saturn to calculate the mass of the Saturnian system.

OBSERVING SCHEDULE: Saturn is only visible in the evening sky for the first half of the Fall 2019 semester. We have several observing nights scheduled during this period. To complete parts 1 and 2 of this lab, you MUST attend at least one of these observing nights***. Do not put this off until the last minute!*** Periods of bad weather and telescope scheduling can make your first-choice evening unavailable. Your grade on this lab is based on completing parts 1 and 2. Part 3 is counted as a regular day-time lab and can be completed even if you don't get to the telescope.

Pre-Lab = please complete *before* coming to lab

Read briefly through Part 1 of the lab. Check this box ☐ to indicate that you have done so. Time at the telescopes is limited, so knowing what to do is crucial to smooth operation. TAs/LAs will be there to assist, but you should expect to do much of the work with the telescope and computer yourself. Jot down any questions that occurred to you while skimming through the lab description.

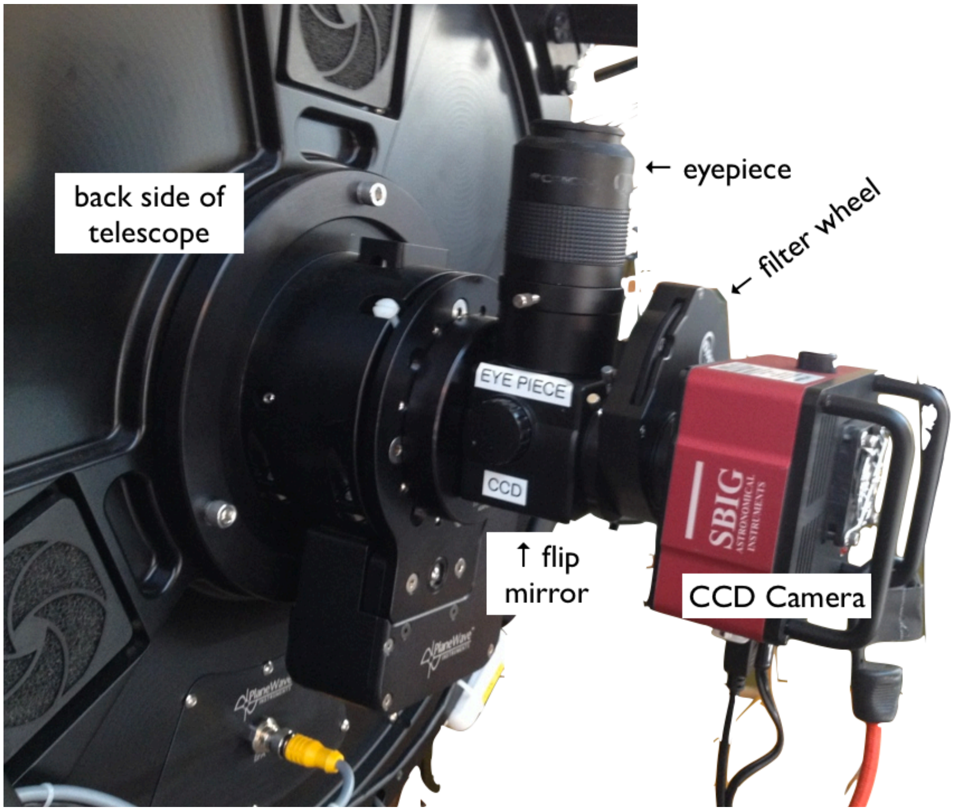
PART 1 - Observing Saturn

1.1 Point telescope to Saturn

1. At the computer terminal on the observing deck, locate the window for TheSkyX software. Find the vertical Telescope tab on left edge of the window.

2. In search box on left side, enter Object Name (Saturn) and click Find. (You can also find Saturn on the map and click on it. A red target symbol should appear on the sky map.

3. Check that the area around the telescope is clear and announce that you are slewing. When the area is clear, click Slew. Telescope should start to move and yellow target symbol should move over to cover the red symbol.



The back side of the 20-inch telescopes with the important parts labeled. The flip mirror allows light from the telescope to enter either the eyepiece or the camera. The filter wheel allows you to block different colors of light from the camera.

1.2 Visual observations:

1. A knob below the eyepiece controls a flip mirror that sends light to either the camera or the eyepiece. Make sure the knob is rotated counter-clockwise as far as it will easily move.

2. Look through the eyepiece (use the step ladder if necessary).

3. Do you see Saturn? Is it in focus? If not, ask a TA/LA to focus the system for you.

4. If Saturn is in the best focus, you should be able to see the *Cassini Division*, a gap between the inner (B) ring and outer (A) ring. You should also be able to see some of Saturn's larger moons. There may be up to five visible. *Titan* is the brightest of these, but you may also see *Rhea, Dione, Tethys, Enceladus,* and *Iapetus.*

5. Record your visual observations below: Make a sketch of what you saw through the eyepiece.

1.3 Camera observations:

1. Turn the flip mirror knob to send light to the camera (rotate the knob clockwise)

2. Click the Filter Wheel tab on left edge of the screen (near the Telescope tab).

3. Use the pull down in the Filter Wheel window to select the red (r) filter (it may already be there) and click Move Now to rotate the filter wheel to move the filter in place. This filter blocks most of the light from the telescope so the camera electronics are not overwhelmed (Saturn is bright!).

4. Click the Camera tab on left edge of screen.

5. Set Exposure Time to **0.1 second**.

6. Click Take Photo. An image should appear in another window in a few seconds.

7. In the image window, click the Histogram icon in the upper left. Move the brightness control sliders on the image to look for the same details you saw with your eye. Is the image in-focus? Is the image saturated? The brightest part of Saturn should be around 20,000 counts (but definitely BELOW 30,000 counts!) Seek help from your TA to check. Note, atmospheric turbulence blurs images and may wash out features. Take another image and see if it is better or worse than the first one.

8. Take a few images (at least 3) until you get one that is as clear and sharp as possible.

1.4 Save your data:

1. When you have gotten the sharpest image you can, save it to the disk (Please save ONLY your best image). In the FITS Viewer window (where the image is), click the Photo pull-down menu select Save As.

2. Save in the class Google Drive directory using a file name that identifies your team and the date (e.g., "Smith_Brown_Jones_sept14.FIT").

1.5 Reset the telescope for the next team:

1. Close your image window.

2. Move the telescope. In the Telescope control panel, Find "Vega", and click Slew.

3. Flip the mirror back to eyepiece mode (turn counter-clockwise) and make sure a bright star appears in the center of the field.

4. **You're done for the night. Congratulations!**

Part 2: Image Analysis

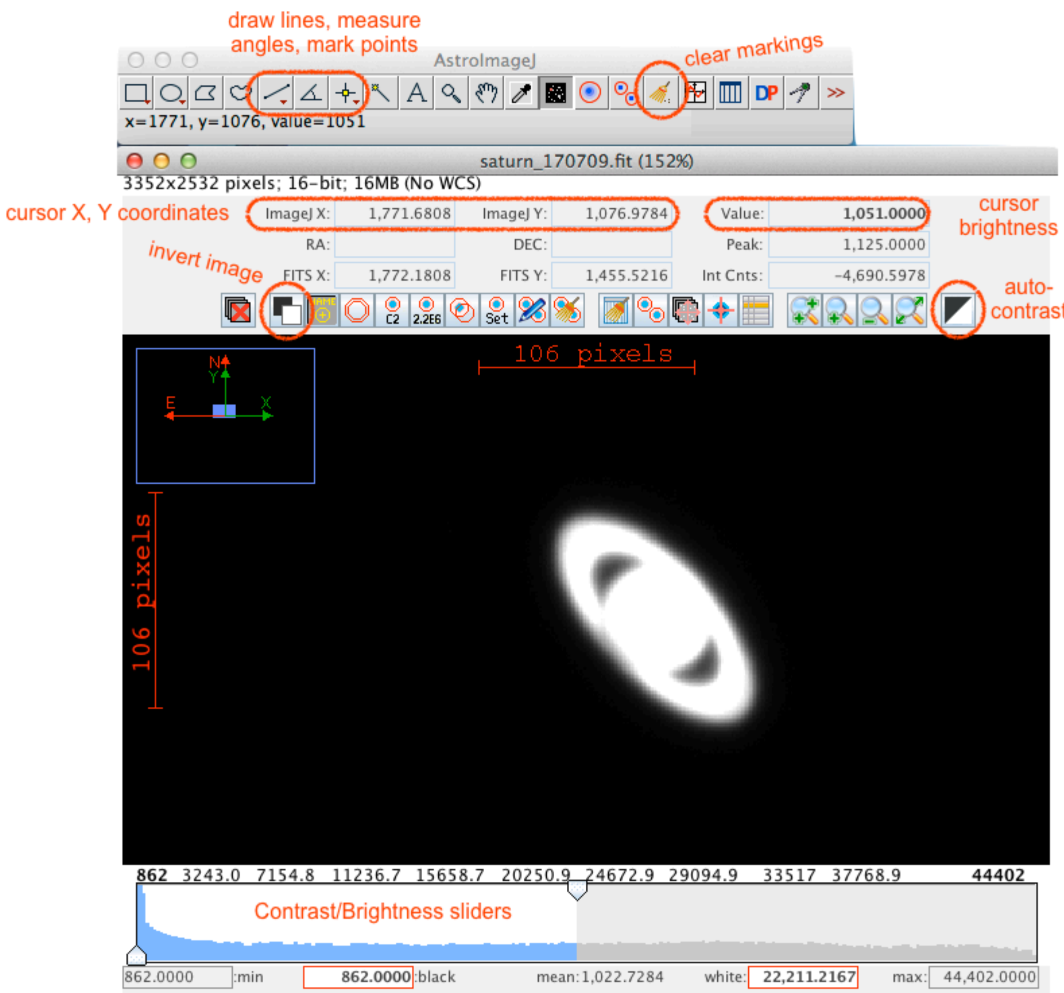
After your observing session, you'll need to do some analysis on your Saturn image. In this class, we will use the *AstroImageJ* software package distributed for free from the University of Louisville. It is a powerful-yet-intuitive code which will run on most computers.

2.1 Download *AstroImageJ*

Download whichever version of the software is appropriate from
http://www.astro.louisville.edu/software/astroimagej/

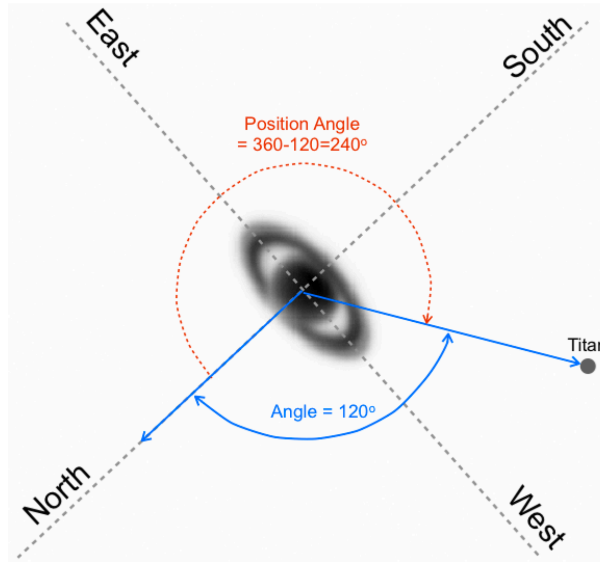
2.2 Copy your observations

Copy the *.fits* file from the class data directory to your local computer.
Open your image using *AstroImageJ* (File > Open)



Take a few minutes to familiarize yourself with the software. Try **zooming** in and out. Notice the X and Y coordinate boxes at the top of the image window and/or in the toolbar window.
Try sliding the bars around on the **contrast sliders**. Try inverting the image (switching black and white with the invert button). Do you see more of less detail on the planet? What setting allows you to see the moons best? Can you see both moons and ring details at the same time?

2.3 Identify coordinates in your image:



Adjust the image so you can clearly see as many details of the planet and rings as possible. It is currently northern summer at Saturn, so the planet is tilted with its northern hemisphere toward us. North corresponds to a line perpendicular to the ring semi-major axis. (In the example above, north points down and to the left.) East is 90 degrees clockwise from the north line. In this example, Saturn would be in the SW quadrant.

2.4 Diameter of Saturn and Rings

Use the line tool in the toolbar to measure the diameter of Saturn in pixels. The line tool will give you a length in pixels, but you can also record the X and Y positions of any two points and use the Pythagorean theorem to compute a length between points 1 and 2:

$$dist^2=(x_1-x_2)^2+(y_1-y_2)^2.$$

Divide this number by two so you have a measurement of Saturn's radius *in pixels*. You will use this number to convert the rest of your measurements into units of Saturn radii.

Saturn diameter: ___________ pixels
Saturn radius: ___________ pixels

Now measure the diameter of the rings (east-west distance) and convert to Saturn radii using the number of pixels you just calculated for Saturn's radius.

Ring diameter: ___________ pixels = ___________ Saturn radii

2.5 Titan

Zoom out and adjust the contrast until you can easily see Saturn and its moon(s).

Identify which moon is Titan (it will be the brightest and probably the farthest from the planet as well). The website https://www.wwu.edu/skywise/saturn.html can help. Note that UT time is 6 hours ahead of Mountain Daylight Time. *(If you cross 24h be sure to add one day to the UT date!)*

Which quadrant does Titan lie in with respect to Saturn? (NE, SW, etc.) ____________

Use the same technique you used above to measure distance from the center of Saturn to Titan.

Titan distance: ___________ pixels = ___________ Saturn radii

Next, measure the "position angle" from Saturn to Titan. This is defined as the direction from the center of Saturn to Titan and measured in degrees east of north. Use the angle tool in *AstroImageJ* to measure the angle. Click Titan, then click the center of Saturn, then another point along the north axis of the planet. Remember that position angle is degrees east of north; if Titan is in the SW or NW quadrant, you'll need to subtract your angle measurement from 360 degrees.

Titan PA = _______________ degrees east of north

It's easy to get the sense of this number wrong, so check which quadrant Titan lies in and compare to the example figure and table. In the example, north lies in the lower left direction, so Titan is in the SW quadrant. The angle measurement from *AstroImageJ* is 120º. However, we must go from north through east, so PA=360º-120º=240º.

Quadrant	PA range
NE	0-90
SE	90-180
SW	180-270
NW	270-360

2.6 Upload your data

In one of the in-class labs, you will be combining your measurements with those of the rest of the class to determine some important parameters about Saturn. As part of this, you will upload your measurements to a form. See the class web page for instructions.

Include your **name** and that of others on your team (this information will be anonymized before presented to the whole class). Enter the Mountain Daylight Time **date** and **time** of your observations. Note **weather**, **equipment**, and any other information you think might be relevant.

Next, record the four quantities you calculated in sections 2.4 and 2.5:

- **radius** of Saturn (pixels)
- **diameter** of ring system (Saturn radii)
- Saturn-Titan **distance** (Saturn radii)
- Saturn-Titan **position angle** (degrees E of N)

2.7 Above and beyond (optional):

Can you identify any other moons in your image? If so, note them down in the comments section of the web form. The website https://www.wwu.edu/skywise/saturn.html will give you the positions of the five bright moons. Note that UT time is 6 hours ahead of Mountain Daylight Time. *(If you cross 24h be sure to add one day to the UT date!)*

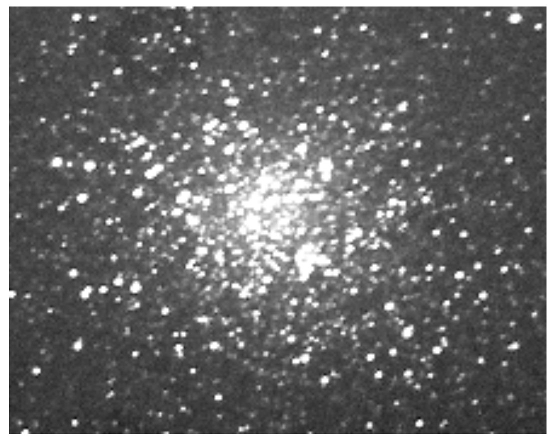
OBSERVATIONS OF DEEP-SKY OBJECTS*

SYNOPSIS: You will view and sketch a number of different astronomical objects through the SBO telescopes. The requirements for credit for telescope observing may vary depending on the requirements of your instructor. The following is given only as a guideline.

EQUIPMENT: Observatory telescopes, observing forms, and a *pencil.*

Be sure to dress warmly - the observing deck is not heated!

Part I. Observing Deep Sky Objects



The two main SBO observing telescopes (twin 20-inch telescopes) are both operated by computer. The user may tell the computer to point at, or the observer may specify the coordinates at which the telescope should point. Deep sky objects are easily selected from the catalog (or using the Find window).

Your instructor may point a telescope to at *least* one of each of the following different types of deep-sky objects (provided that weather cooperates, and the appropriate objects are visible in the sky at the time). Distinguishing characteristics to look for have been included in italics.

- **Double or multiple stars**. *Separation of the stars, relative brightness, orientation, and color of each component.*
- **Open clusters**. *Distribution, concentration, and relative brightness and color of the stars.*
- **Globular clusters**. *Shape, symmetry, and central condensation of stars.*
- **Diffuse nebulae**. *Shape, intensity, color, possible association with stars or clusters.*
- **Planetary nebulae**. *Shape (ring, circular, oblong, etc.), size, possible central star visible.*
- **Galaxies**. *Type (spiral, elliptical, irregular), components (nucleus, arms), shape and size.*

For *each* of the above objects that you observe:

I.1 In the spaces provided on the observing form, fill in the object's name, type, position in the sky (RA and Dec), the date/time you observed it, and the weather conditions. Make certain to note what *constellation* the object is in, because this information is almost essential when using the reference books.

I.2 Observe through the telescope and get a good mental image of the appearance of the object. You may wish to try averted vision (looking out of the corner of your eye) to aid you in seeing faint detail. Take your time; the longer you look, the more detail you will be able to see.

I.3 Using a pencil, carefully sketch the object from memory, using the circle on the observing form to represent the view in the eyepiece. Be as detailed and accurate as possible, indicating color, brightness, and relative size.

I.4 Include an "eyepiece impression" of what you observed: a brief statement of your impressions and interpretation. Feel free to draw upon comparisons (for example, "like a smoke ring," or "a little cotton ball," etc.). Express your own enthusiasm or disappointment in the view!

I.5 If you wish (or if your instructor has required it), research some additional information on your objects. The Observatory lab room has some sources, as does the Math-Physics Library. Specific useful books are *Burnham's Celestial Handbook*, the *Messier Album*, and various textbooks. Read about the object, and then provide any additional information that you find is particularly pertinent or interesting.

Part II. Planetary Observations

Most of the planets (other than the Earth) are readily observed with the SBO telescopes. The difficult ones are Pluto (tiny and faint) and Mercury (usually too close to the Sun). Provided that they are available in the sky this semester (consult the web resources at the beginning of this manual):

II.1 Observe, sketch, and research at least two of the solar system planets, as in paragraphs I.1 through I.5 above. Pay particular attention to relative size, surface markings, phase, and any special features such as moons, shadows, or rings. You may wish to use different magnifications (different eyepieces) to try to pick out more detail.

NAME:

Object Name
Object Type
Constellation
R.A.
Dec.
Date
Time
Telescope size
Sky Condition

Description of Observation

Additional Information

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Object Type
Constellation
R.A.
Dec.
Date
Time
Telescope size
Sky Condition

Description of Observation

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OBSERVATIONS OF THE MOON*

SYNOPSIS: You will investigate the Moon through telescopes and binoculars, identifying and sketching several of the lunar features.

EQUIPMENT: Sommers-Bausch Observatory telescopes and binoculars, lunar map, lunar observing forms, and a *pencil.*

Be sure to dress warmly - the observing deck is not heated!

Part I. Lunar Features

Listed below are several types of lunar features. Read the description of feature types, and identify at *least* one example of each, using either a telescope or binoculars. Locate and label each feature on one of the lunar outline charts below. (Note that one of the charts is presented in "normal view," which resembles the appearance of the Moon as seen through binoculars, while the other is a "telescope view," which is a mirror image of the Moon as it may appear through the telescopes. Use either or both charts at your convenience.)

You will encounter additional features not shown on the outline charts, such as small craters. Feel free to add them as you view them. Feature types marked with "T" are best seen through a telescope, while those marked with "B" can be seen with binoculars.

I.1 **Maria**: These are relatively smooth and dark areas formed by ancient volcanic eruptions that filled even older giant impact craters. The maria make up the "man-in-the-Moon." These were once thought to be seas, during the early days of the telescope. (B) *Shade in these dark patches with your pencil.*

I.2 **Craters with central peaks**: Many large craters have mountain peaks in their centers, which can reach 5 km in height. These peaks are produced by a rebound shock wave produced by the impact that formed the crater. (T)

I.3 **Craters with terraced walls**: As some large craters formed, their inner walls collapsed downward, pulled by gravity. This can happen several times, giving the inner crater wall a stair-stepped appearance. (T)

1.4 **Overlapping craters**: An impact crater may be partially obliterated by a later impact, giving clear evidence of which impact occurred earlier, and which occurred later. (T)

I.5 **Craters with rays**: Some younger craters have bright streaks of light material radiating from them. These rays are created by debris tossed out by the impact that formed the crater. Craters with bright rays are relatively "young" (less than 1 billion years old); the rays of older craters have been obliterated by subsequent geologic activity or impacts. Rays are most prominent near the time of the full Moon. (B)

I.6 **Walled plains**: A few very large craters have bottoms that are partially filled by mare lava. The appearance is that of a large flat area surrounded by a low circular wall. (T)

I.7 **Rilles:** Rilles are trenches in the lunar surface that can be straight or irregular. Although some of them look like dried riverbeds, they were not formed by water erosion, but rather by ancient flows of liquid lava. Straight rilles are probably geological faults, formed by ancient "moonquakes." (T)

I.8 **Mountains and mountain ranges**: The Moon's mountains are the remnant rims of ancient giant impact craters. Because of the Moon's low gravity and slow erosion, these mountain peaks can reach heights of 10 km. (B or T)

Part II. The Terminator

The **terminator** is the sharp dividing line between the sunlit and dark sides of the Moon's face.

II.1 If the Moon is not full on the night of your observations, carefully sketch the terminator on the chart. Include irregularities in the line, which give visual clues to the different heights in the lunar features (high mountains, crater edges, and low plains). (B)

II.2 Inspect the appearance of craters near the terminator, and those that are far from it. How does the angle of sunlight make the craters in the two regions appear different? In which case is it easier to identify the depth and detail of the crater? (If the Moon is full, look for craters near the edge of the Moon, and contrast with those near the center.) (T)

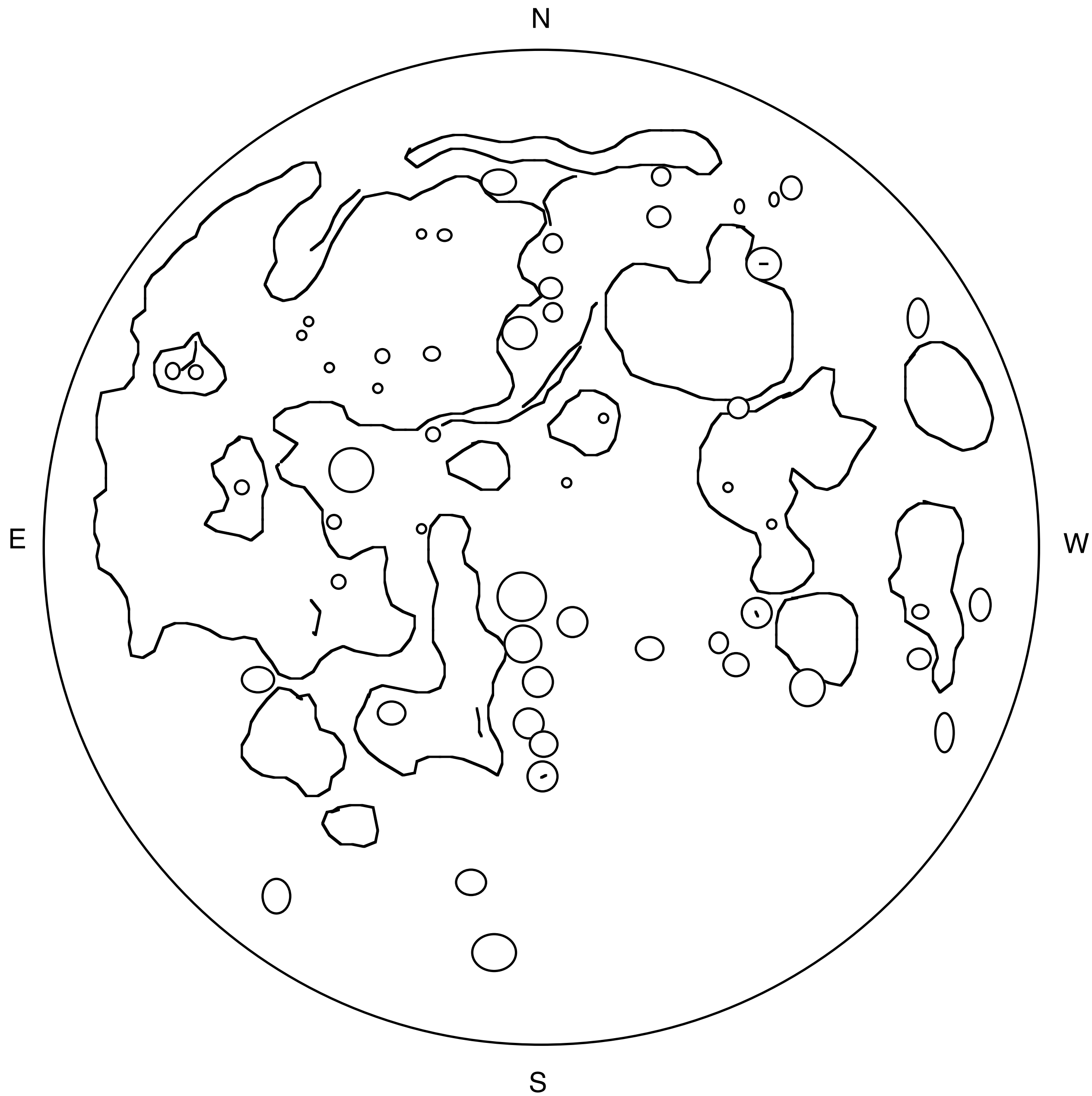
II.3 If you were standing on the Moon at the terminator, describe what event you would be experiencing.

Part III. Lunar Details

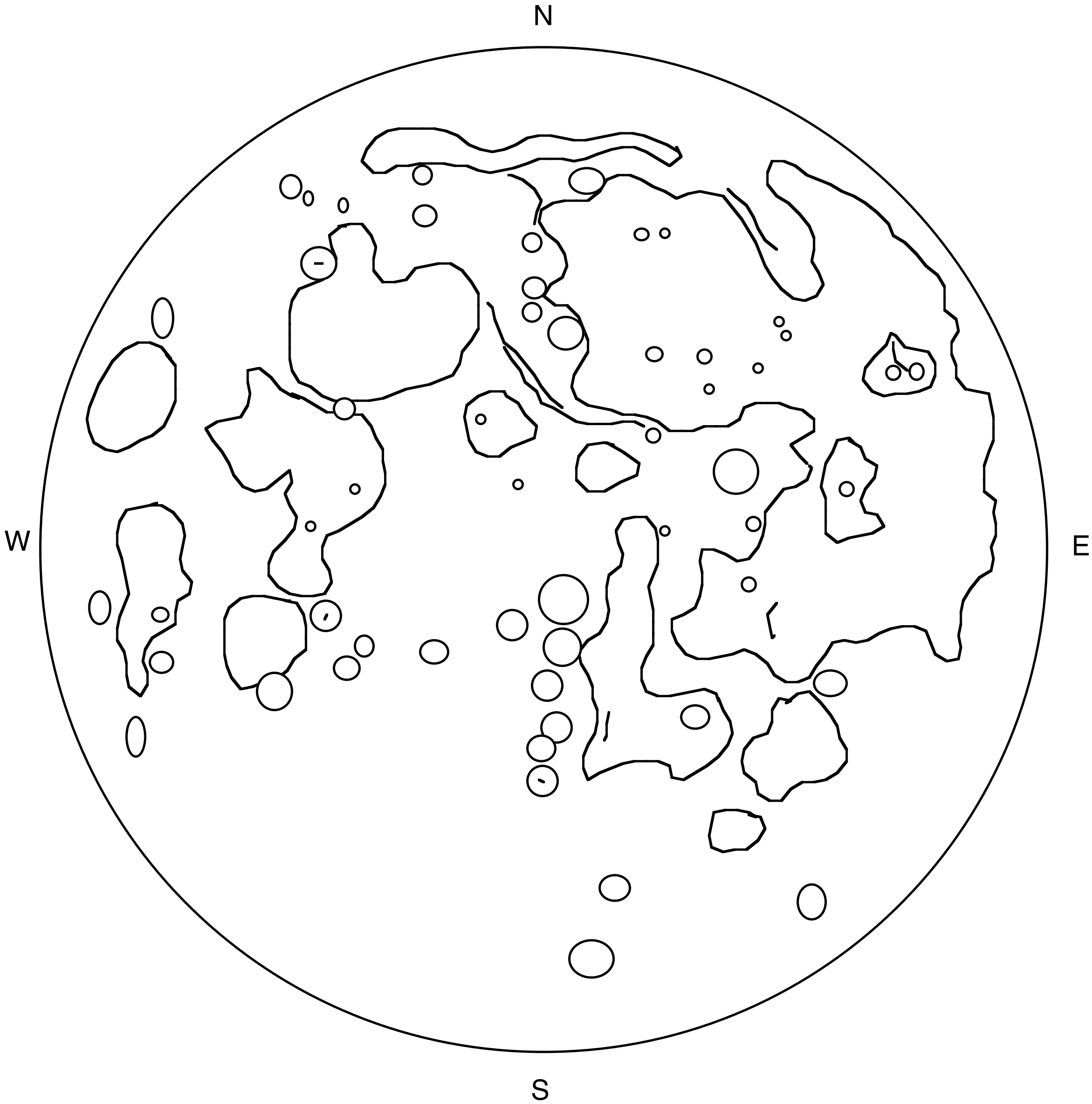
III.1 Select two lunar features of particular interest to you. Use the attached lunar sketch sheet to make a detailed pencil sketch of their telescopic appearance. Be sure to indicate their locations on an outline chart, so that you can later identify the features. (T)

Part IV. Lunar Map Identification

III.1 Compare your finished lunar outline charts and observing sheets with a lunar map and determine the proper names for the features you have identified and sketched.



Binocular View



Telescope (Inverted) View

North may not be "up" in the eyepiece

NAME:

Object Name	
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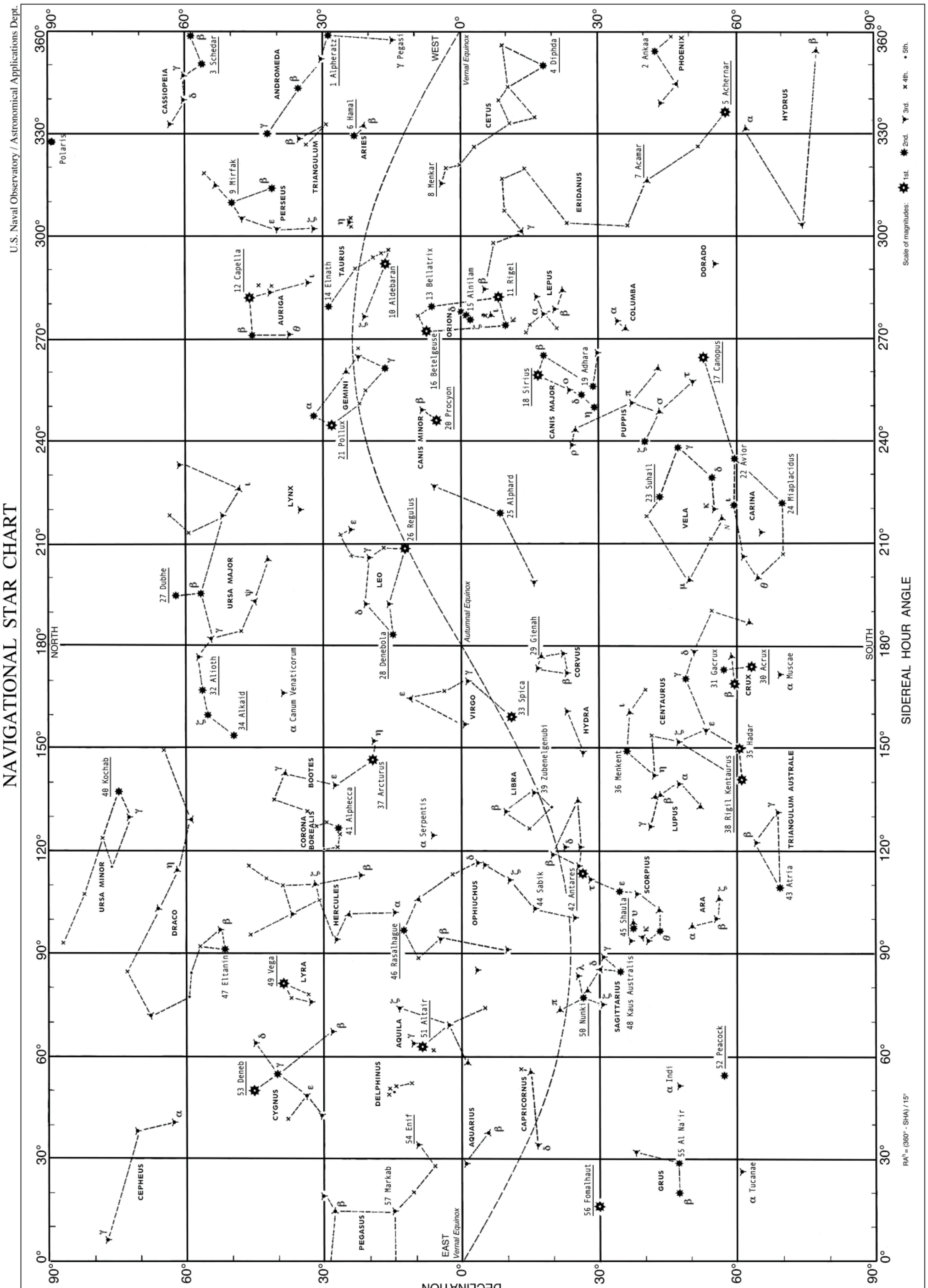
NAVIGATIONAL STAR CHART
U.S. Naval Observatory / Astronomical Applications Dept.
SIDEREAL HOUR ANGLE
DECLINATION
NORTH
SOUTH
EAST
WEST
Vernal Equinox
Autumnal Equinox
1 Alpheratz
2 Ankaa
3 Schedar
4 Diphda
5 Achernar
6 Hamal
7 Acamar
8 Menkar
9 Mirfak
10 Aldebaran
11 Rigel
12 Capella
13 Bellatrix
14 Elnath
15 Alnilam
16 Betelgeuse
17 Canopus
18 Sirius
19 Adhara
20 Procyon
21 Pollux
22 Avior
23 Suhail
24 Miaplacidus
25 Alphard
26 Regulus
27 Dubhe
28 Denebola
29 Gienah
30 Acrux
31 Gacrux
32 Alioth
33 Spica
34 Alkaid
35 Hadar
36 Menkent
37 Arcturus
38 Rigil Kentaurus
39 Zubenelgenubi
40 Kochab
41 Alphecca
42 Antares
43 Atria
44 Sabik
45 Shaula
46 Rasalhague
47 Eltanin
48 Kaus Australis
49 Vega
50 Nunki
51 Altair
52 Peacock
53 Deneb
54 Enif
55 Al Na'ir
56 Fomalhaut
57 Markab
Polaris
Scale of magnitudes: 1st. 2nd. 3rd. 4th. 5th.
RA° = (360° - SHA) / 15°

Exoplanet Travel Bureau

Some 40 light-years from Earth, a planet called TRAPPIST-1e offers a heart-stopping view: brilliant objects in a red sky, looming like larger and smaller versions of our own moon. But these are no moons. They are other Earth-sized planets in a spectacular planetary system outside our own. These seven rocky worlds huddle around their small, dim, red star, like a family around a campfire. Any of them could harbor liquid water, but the planet shown here, fourth from the TRAPPIST-1 star, is in the habitable zone, the area around the star where liquid water is most likely to be detected. This system was revealed by the TRAnsiting Planets and PlanetesImals Small Telescope (TRAPPIST) and NASA's Spitzer Space Telescope. The planets also are excellent targets for NASA's James Webb Space Telescope. Take a planet-hopping excursion through the TRAPPIST-1 system.

NASA's Exoplanet Exploration Program. Jet Propulsion Laboratory, Pasadena CA.
exoplanets.nasa.gov